Progress of underground nuclear astrophysics experiment (JUNA)

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CIAE/SUSTech

The 17th International Symposium on Origin of Matter and Evolution of Galaxies
Chengdu, China
September 8-11, 2024

Thanks NSFC, Yalong power, THU, CAS and CNNC

北京師範大學 BEIJING NORMAL UNIVERSITY
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中国原子能科学研究院 中核集团 CHINA INSTITUTE OF ATOMIC ENERGY
雅砻江流域水电开发有限公司 YALONG RIVER HYDROPOWER DEVELEPOMENT COMPANY, LTD.
南方科技大学 SOUTHERN UNIVERSITY OF SCIENCE AND TECHNOLOGY
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四川大学 SICHUAN UNIVERSITY
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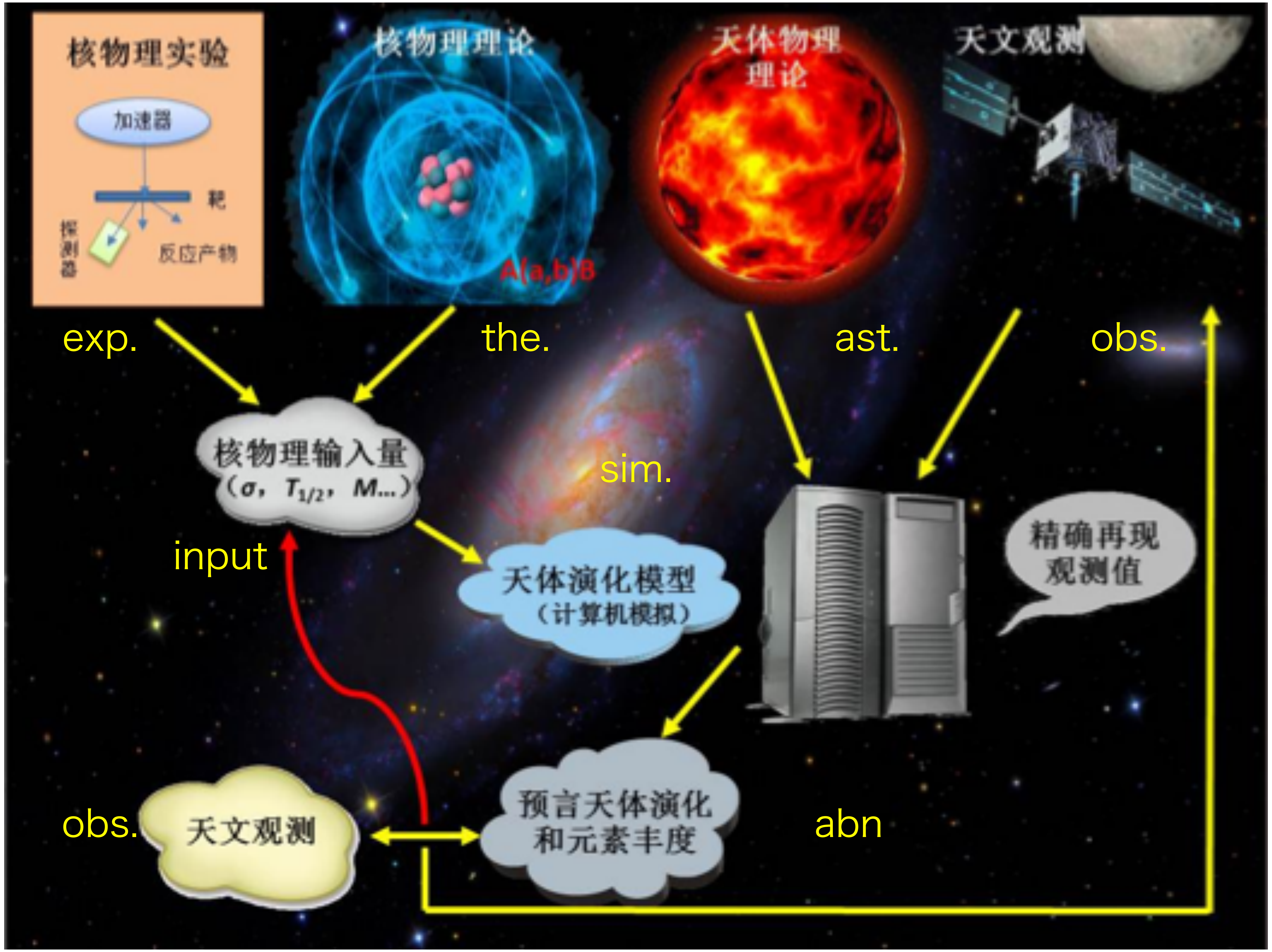
中山大學 SUN YAT-SEN UNIVERSITY
清華大學

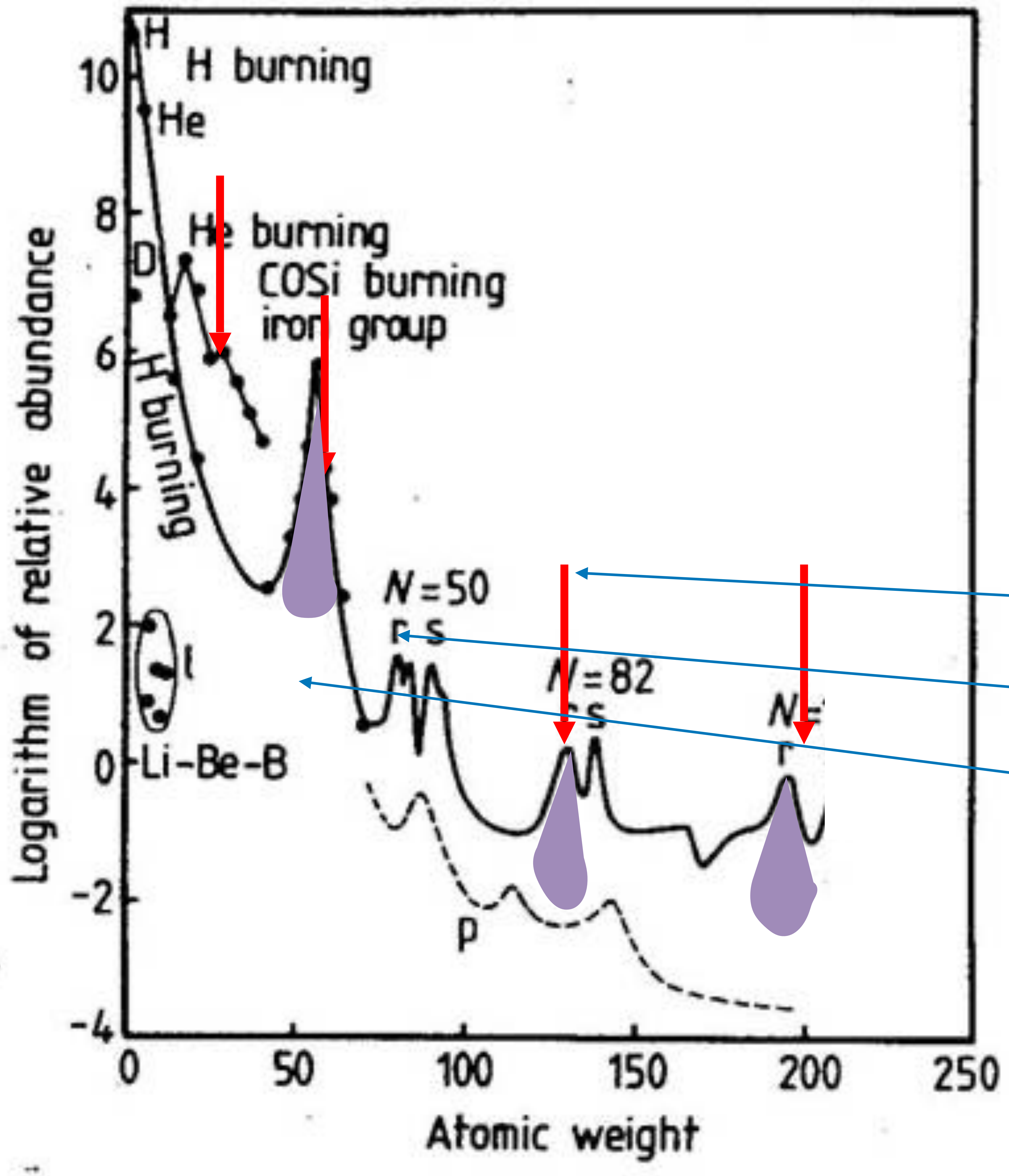
nuclear astrophysics

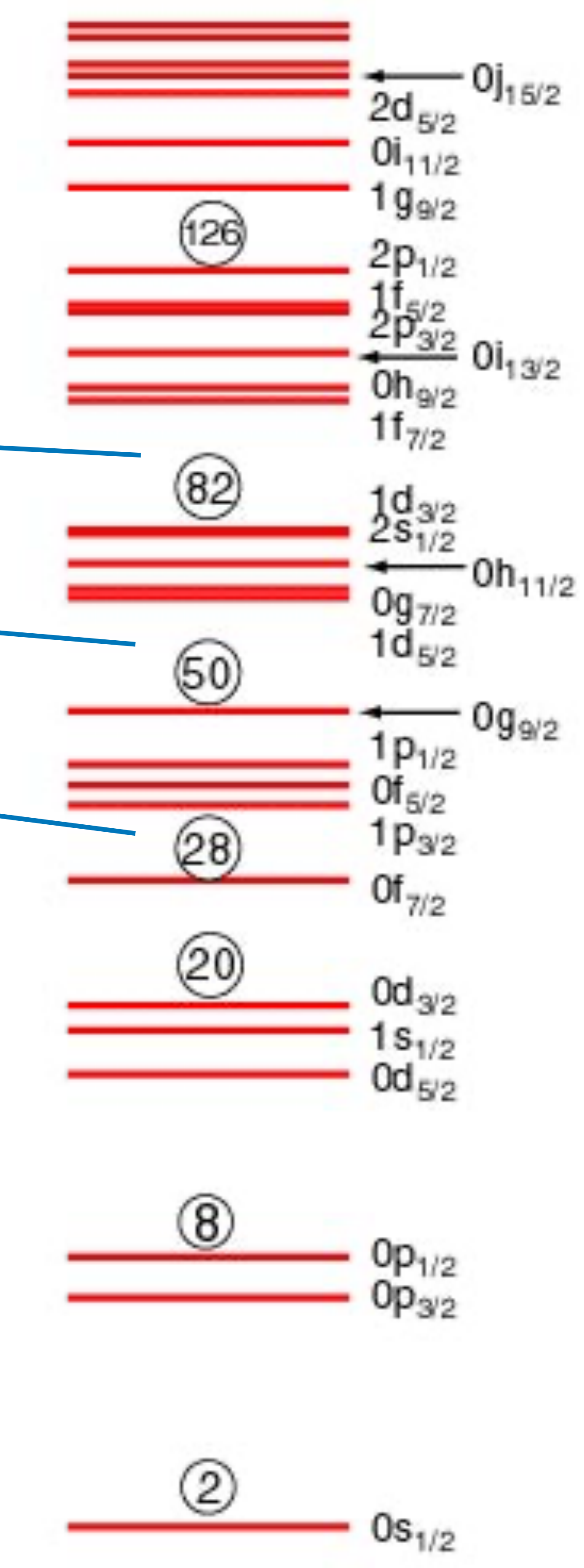


- NP, microscopic, 10^{-15} m, —>observation, cosmic, 10^{14} m, truly interdisciplinary
- For energy production and element synthesis in star

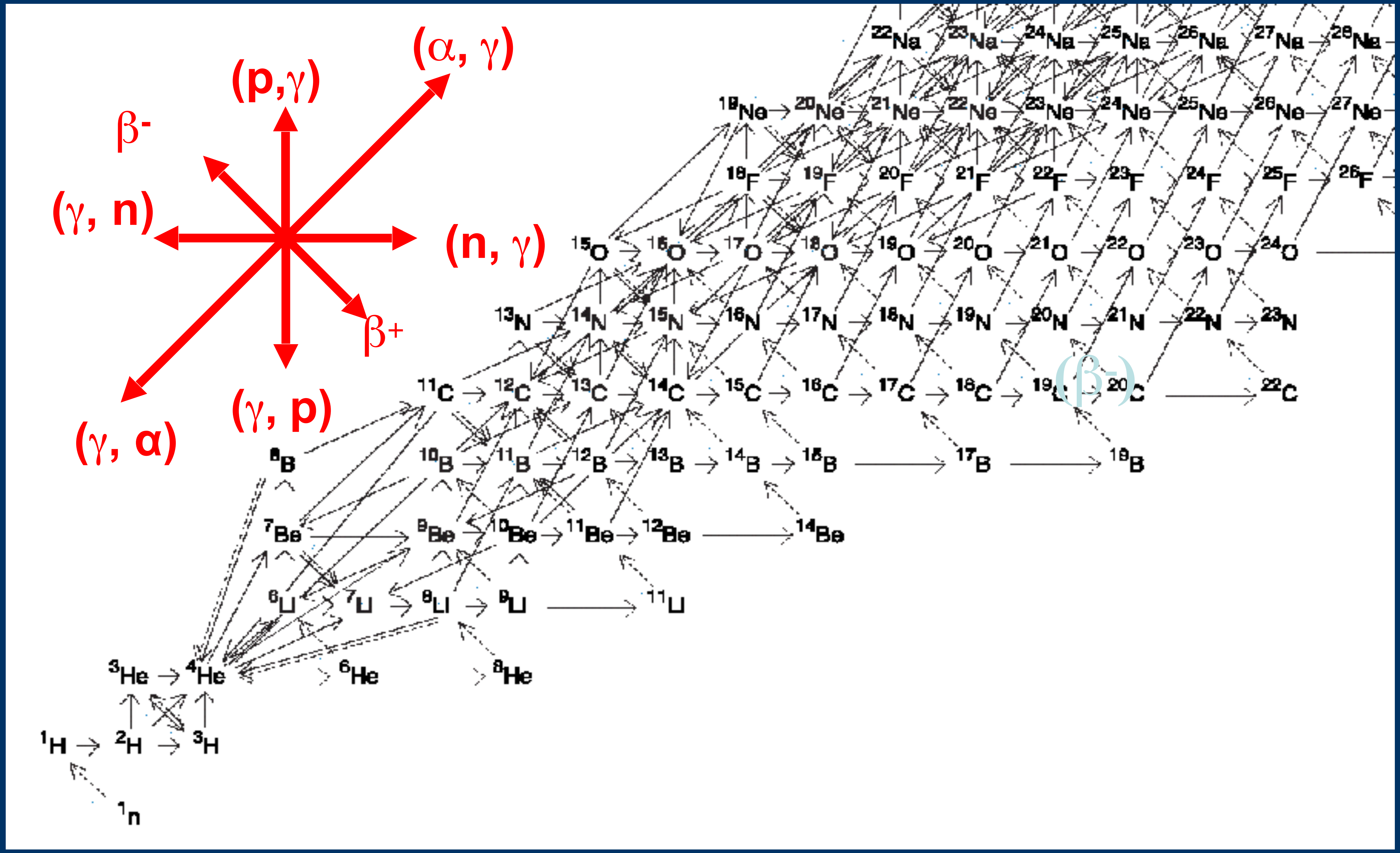
Nuclear Reactions: Alchemists in the Universe

Peaks are the birthmark of nuclear physics: the magic number of the nuclear shell model





Element synthesis network



Cross section

$$\frac{dY_i}{dt} = \sum_j N^i_j \lambda_j Y_j + \sum_{j,k} N^i_{j,k}\rho N_A <\sigma V>_{jk,i} Y_j Y_k + \sum_{j,k,l} N^i_{j,k,l}\rho^2 N_A^2 <\sigma V>_{jkl,i} Y_j Y_k Y_l$$

Decay half-life

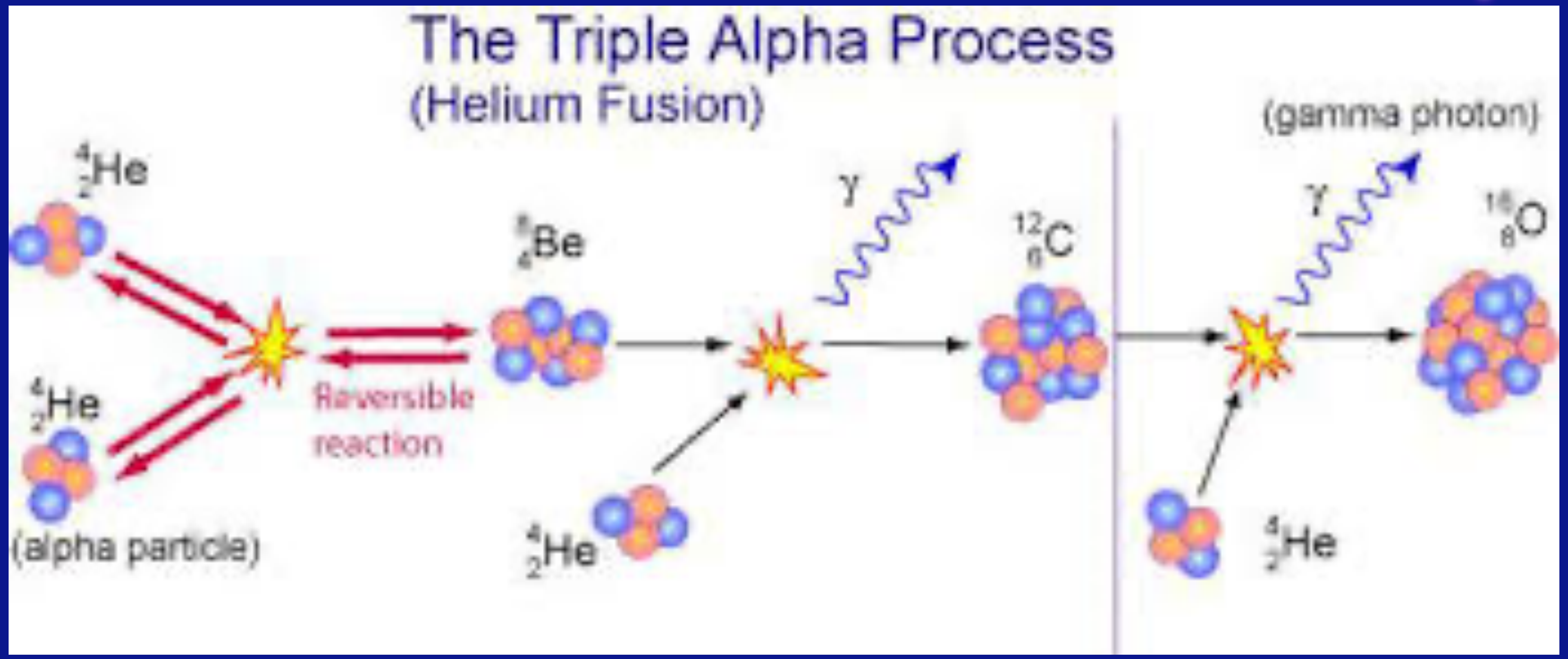
State of nuclei, fate of star

α 0+ α 7.65 0+

1954, Hoyle state, large amount of Carbon

^{4}He ^{8}Be 3α reaction bridge

Fred Hoyle (1915-2001)
APJS 1(1954)121



8.87 2-

regulator

α

7.12 1-

Adequate amount: for sun burn long, for human live

^{12}C

William A. Fowler (1911–1995)
Rev.Mod.Phys. 56 (1984) 149-179

$^{12}C(\alpha,\gamma)^{16}O$

Holy grail reaction

^{16}O

from “Claudon in the universe”

Nuclear astrophysics in China

JUNA experiment

Direct in Gamow window (underground)

RIBLL and LEAF experiment

Direct in higher energy

CIAE experiment

In-direct measurements

BRIF, CSR experiment

Nuclear decay

CSR experiment

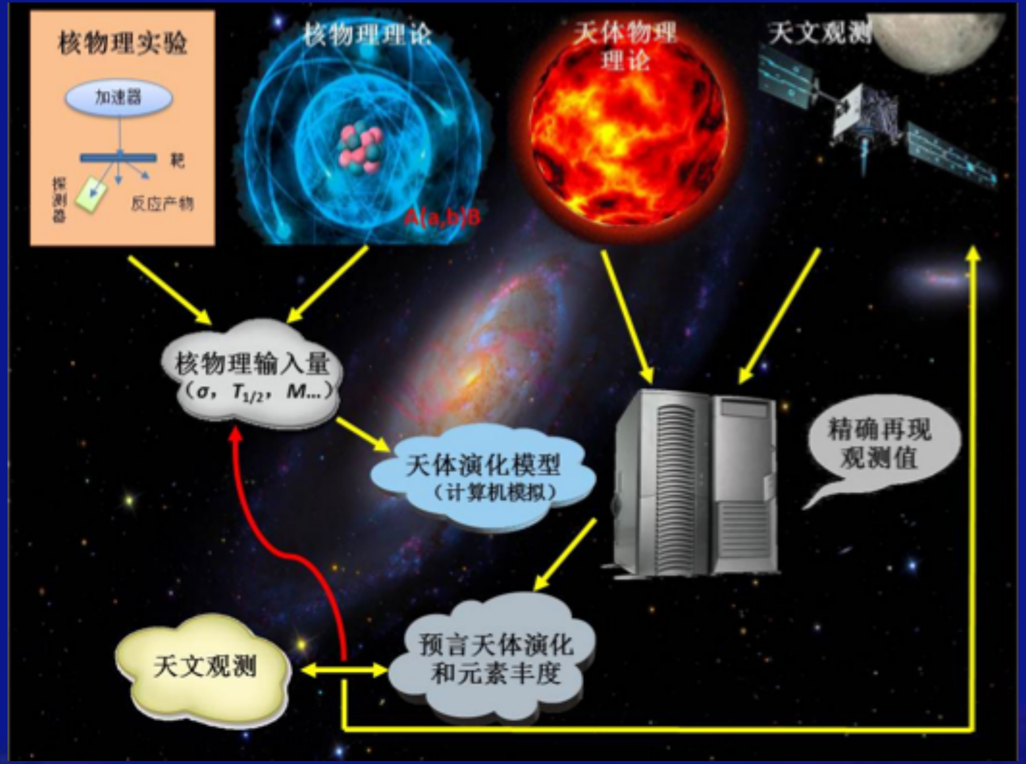
Nuclear mass

theoretical study

Nuclear astrophysics and sensitivity study

theoretical study

Shell model and mean field calculation



Shell model and mean field calculation

theoretical study

theoretical study

Nuclear astrophysics network calculation

Reaction rate database

Nuclear input database

Mass data compilation in China

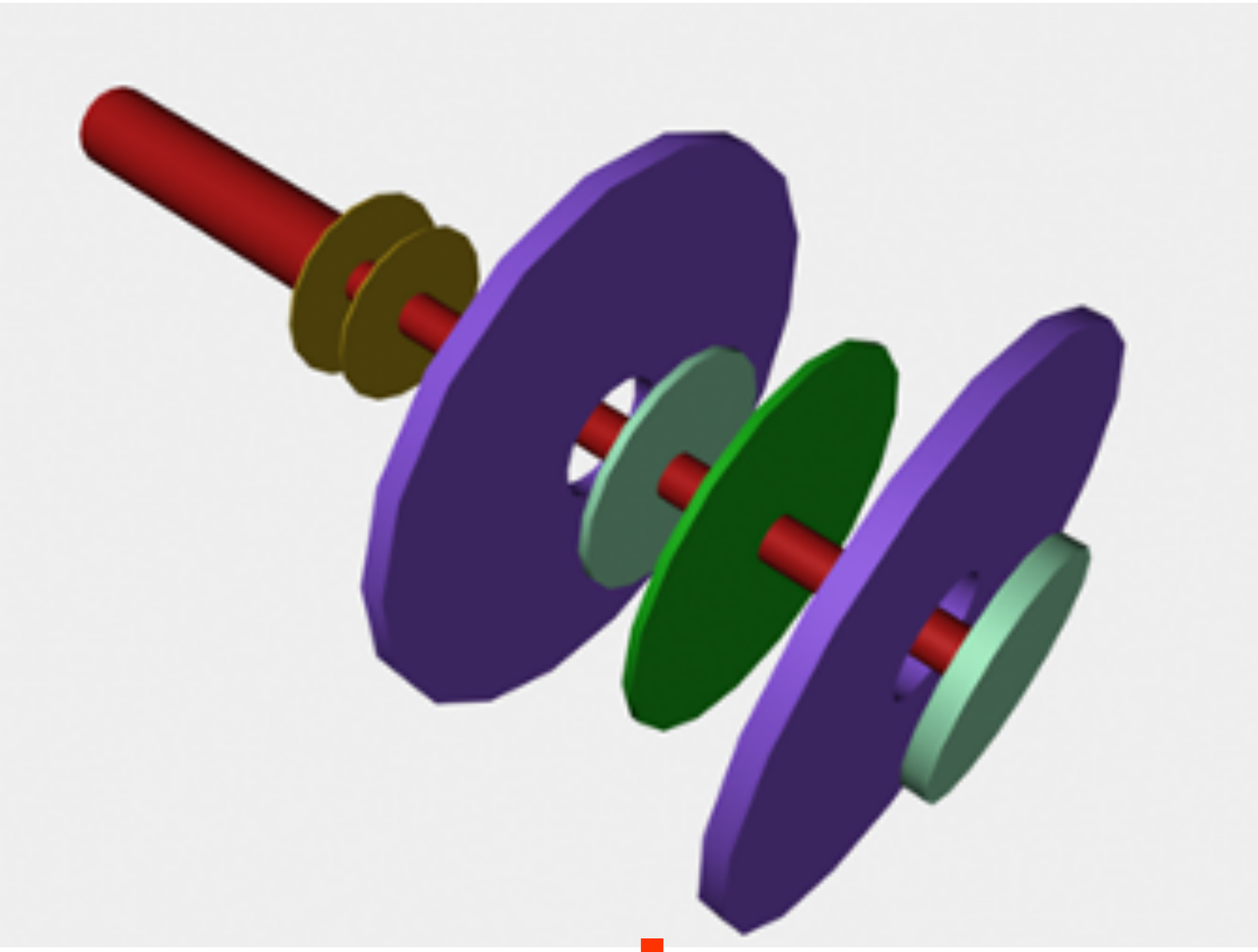
Mass and decay rate database

1996: new method for ^{8}B solar neutrino cross section

RIB production

X. X. Bai, WPL et al., NP A588(1995)273c

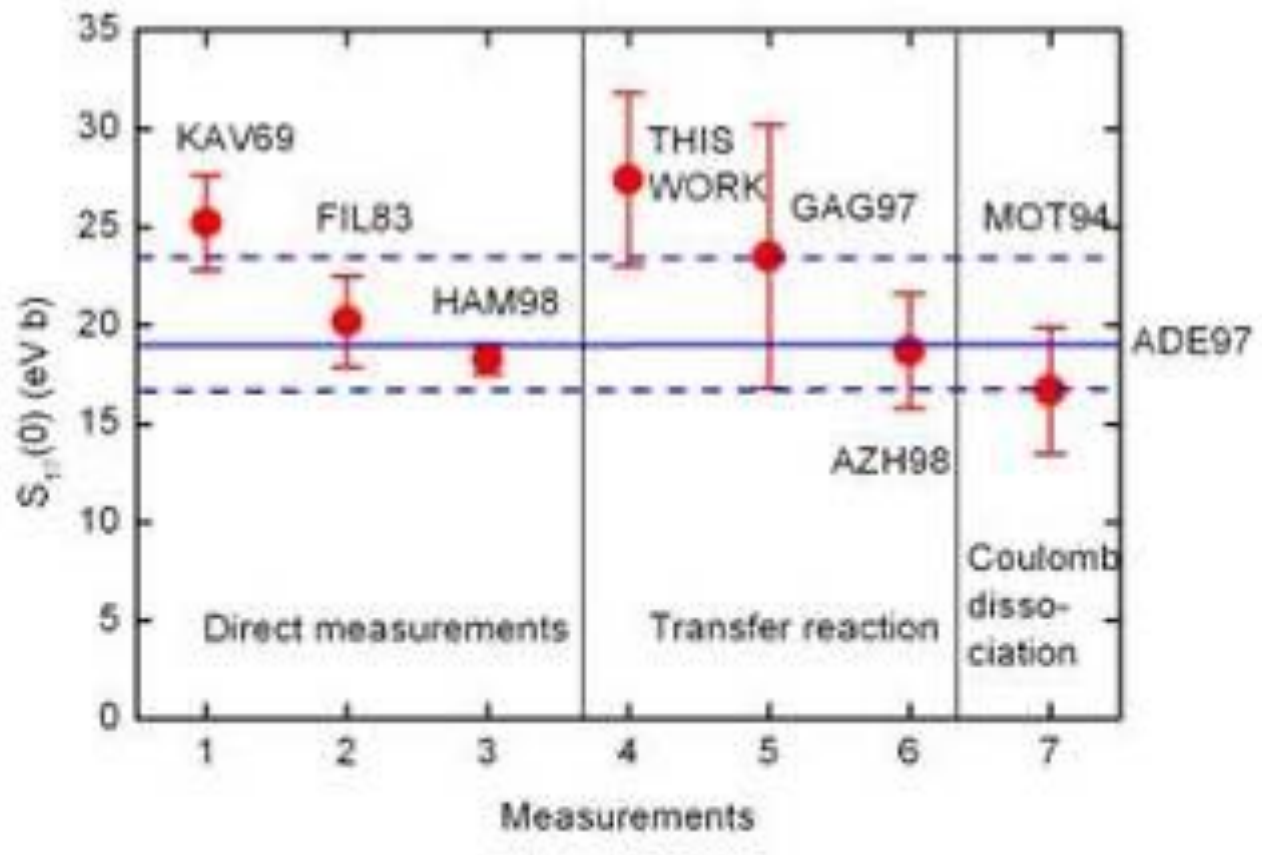
(d,n) or (d,p) measurement



$$\left(\frac{d\sigma}{d\Omega}\right)_{exp} - \left(\frac{d\sigma}{d\Omega}\right)_{CN} = \sum_{j_i j_f} (C^{d}_{l_i j_i})^2 (C^{^{12}N}_{l_f j_f})^2 \frac{d\sigma^{DW}_{l_f j_f l_i j_i}/d\Omega}{b^2_{l_i j_i} b^2_{l_f j_f}},$$

$$\sigma_t = \frac{16\pi}{9}\left(\frac{E_\gamma}{\hbar c}\right)^3 \frac{1}{\hbar v}\frac{e^2_{eff}}{k^2}\frac{(2j_f+1)}{(2I_1+1)(2I_2+1)} C^2_{\ell_f j_f} \times \left|\int_{R_N}^{\infty} r^2 dr\, f_{\ell j}(kr) W_{\eta,\ell_f+1/2}(2\kappa r)\right|^2,$$

ANC or Spec factor



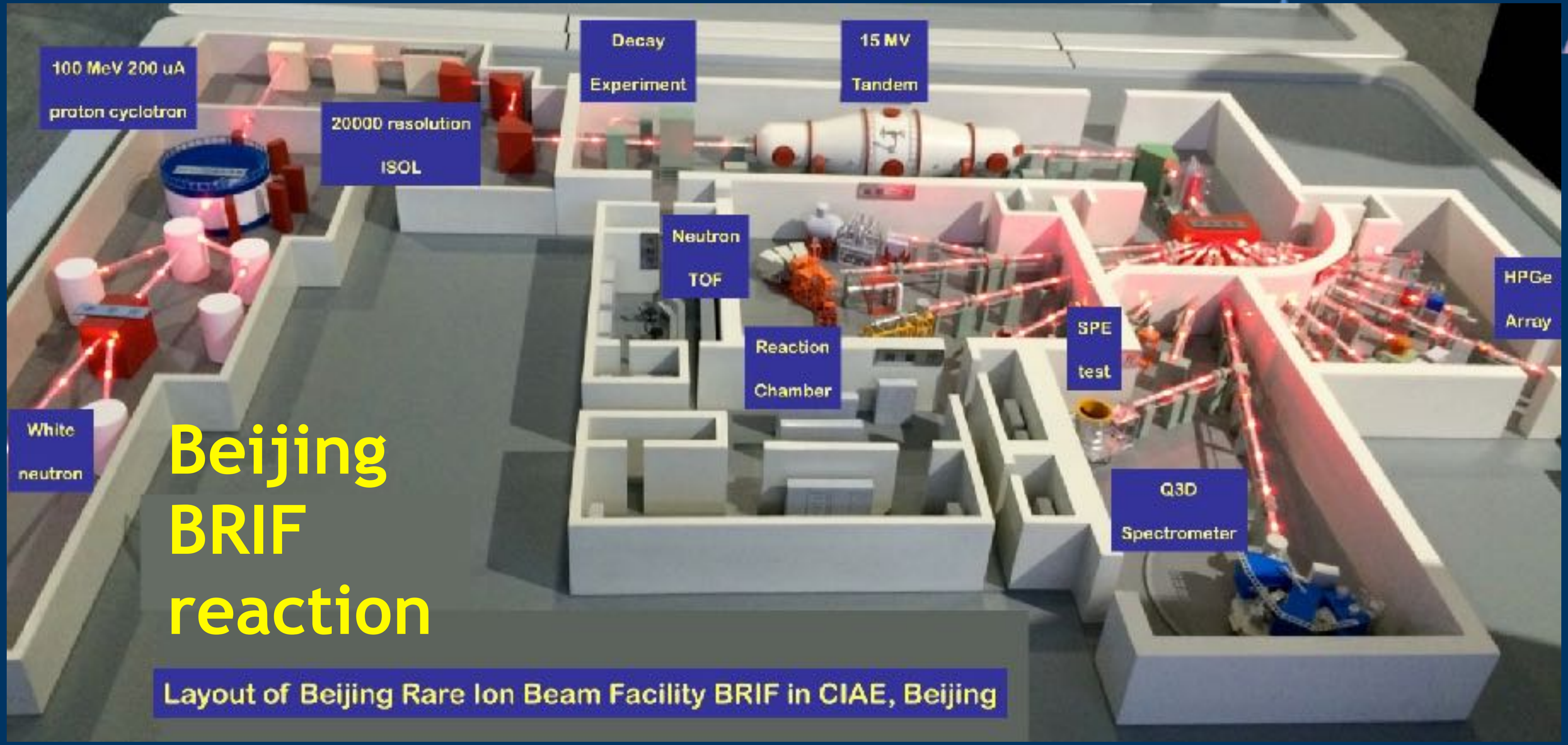
Astrophysical reaction rates

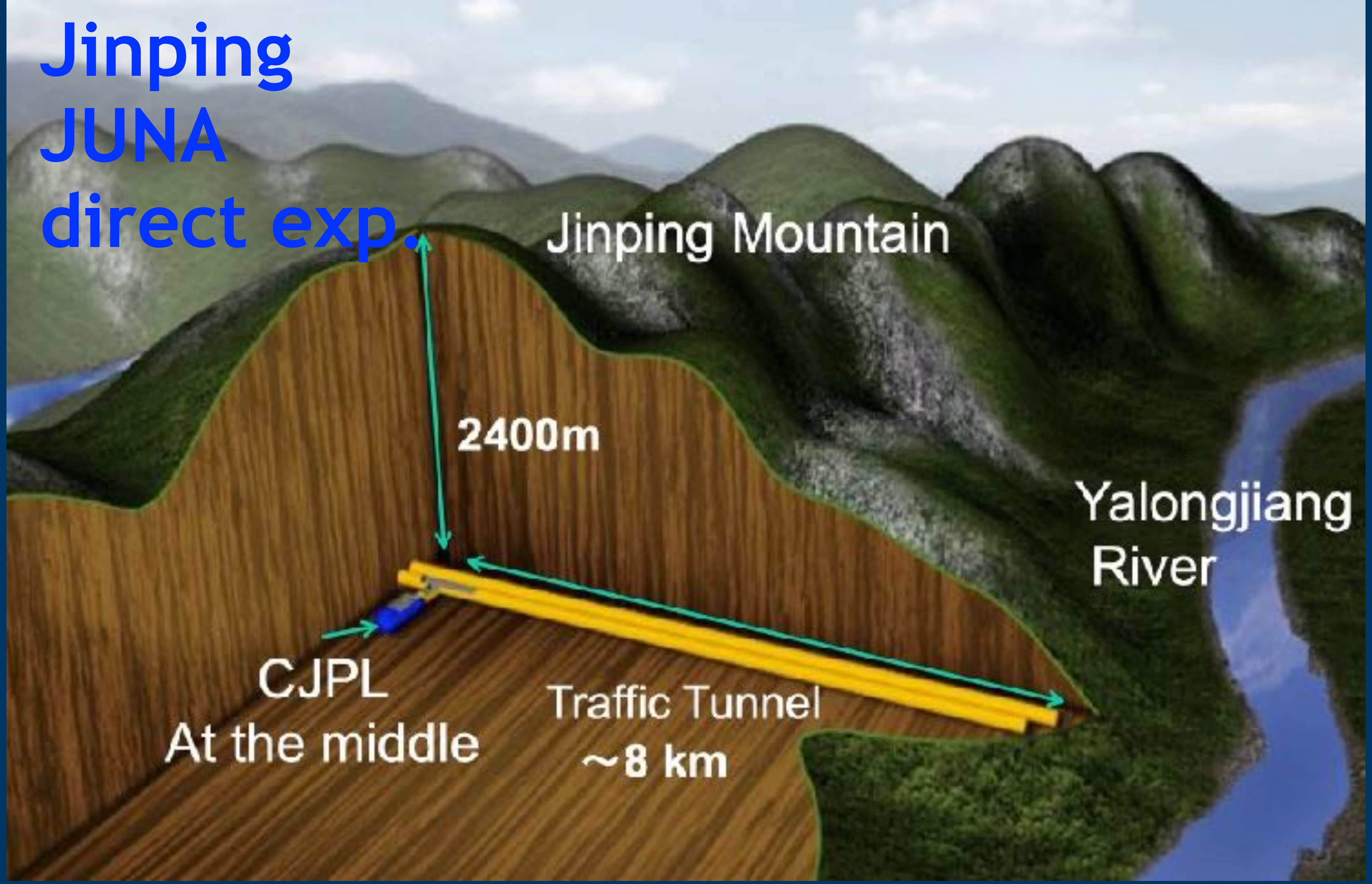
This paper describes an excellent experiment, one of the first examples where a radioactive ion beam has been used in inverse transfer reaction studies.

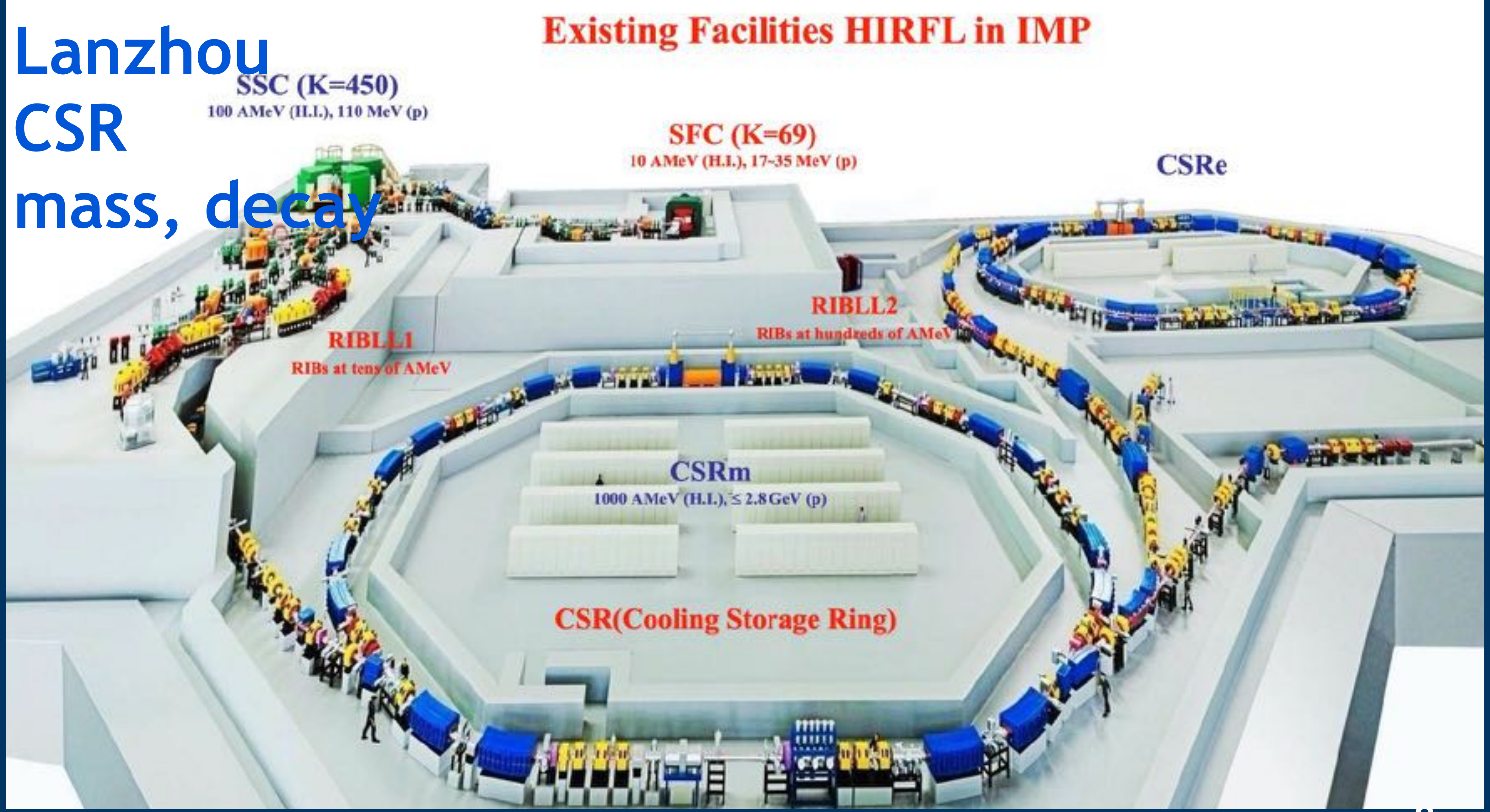
The topic is sufficient important that this paper should see timely and widespread exposure to the physics community in order to stimulate a board-based dialog.

WPL et al., PRL77(1996)611, 1st NP exp. paper in PRL in China

Major facilities in China

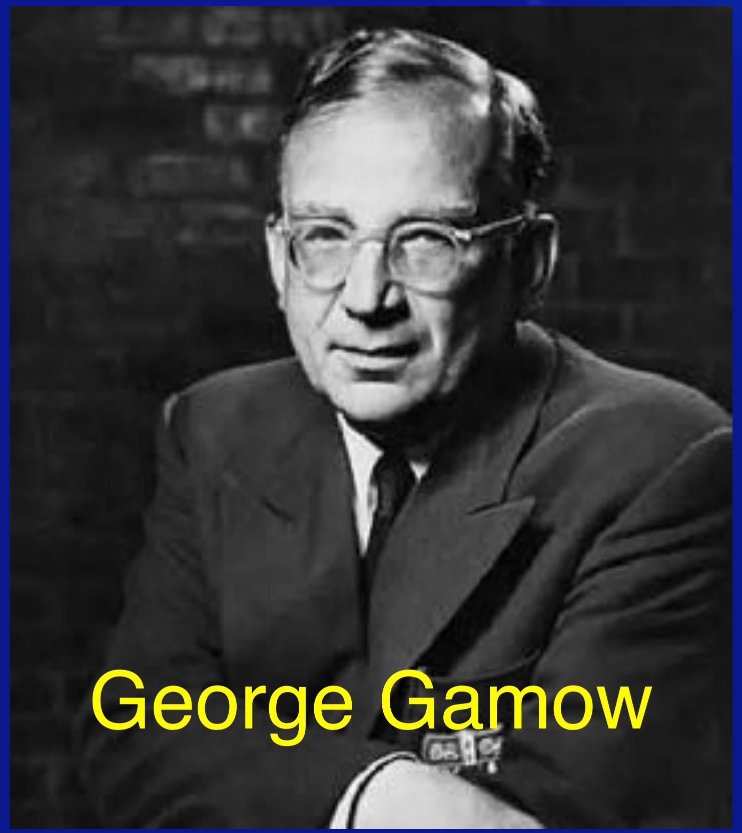






W. P. Liu, 2022, NuSYS

Gamow window



coulomb term

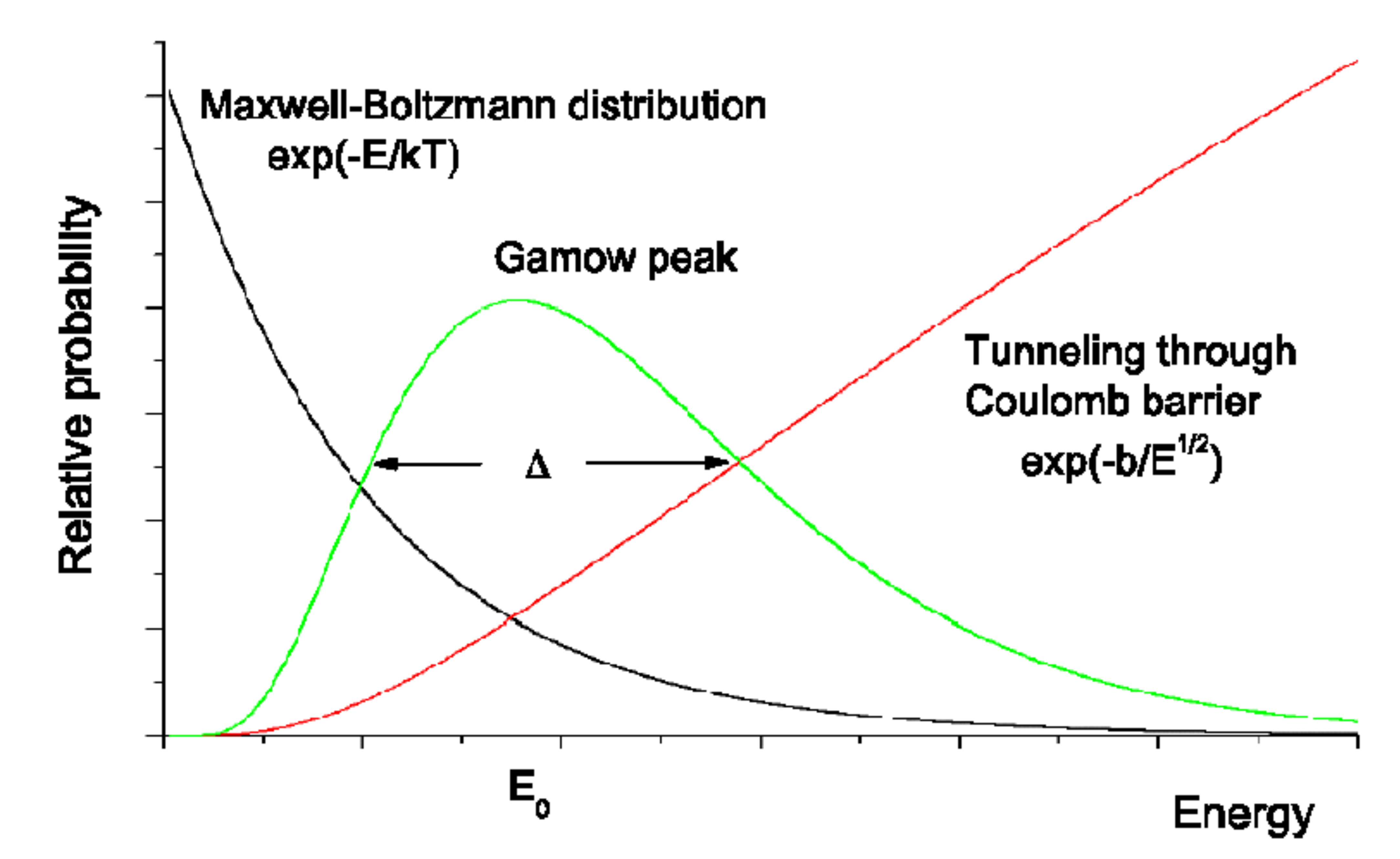
$$\sigma(E) = S(E)e^{-2\pi\eta}\frac{1}{E}$$

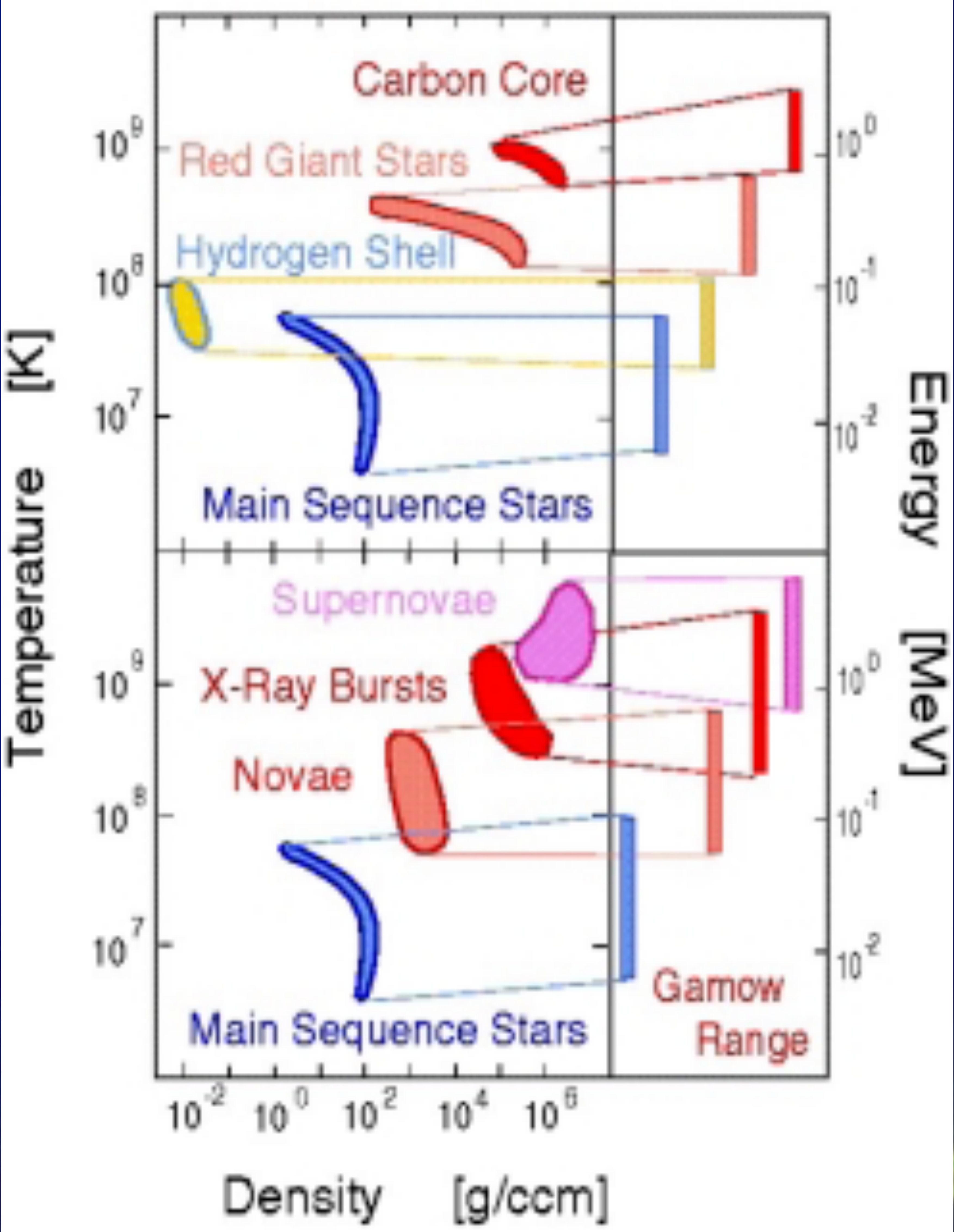
astrophysical s factor

$$\eta = 0.1575 Z_1 Z_2 \sqrt{M/E}$$

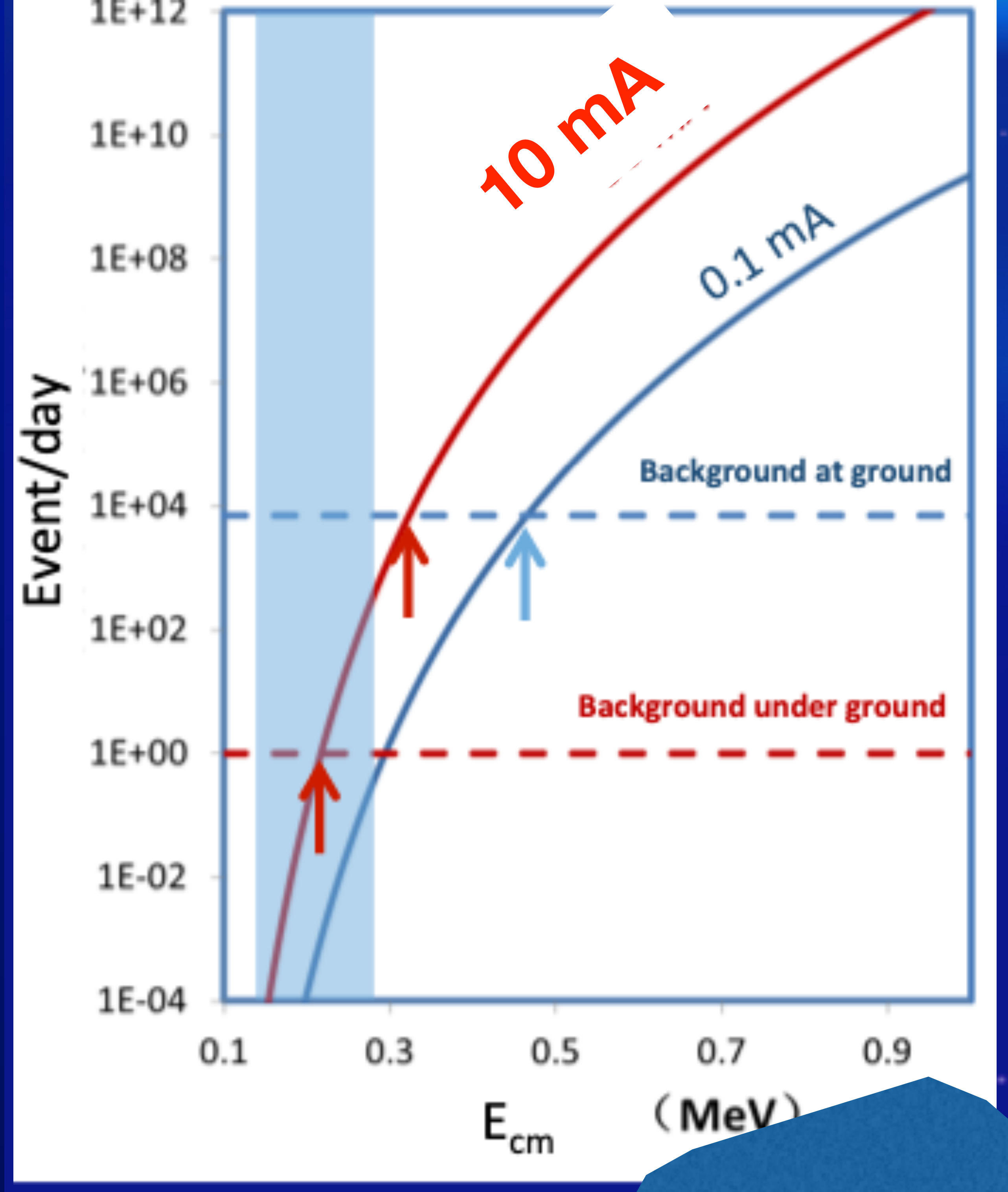
$$E_0 = 1.22(Z_1^2 Z_2^2 M T_6^2)^{1/3}\mathrm{keV}$$

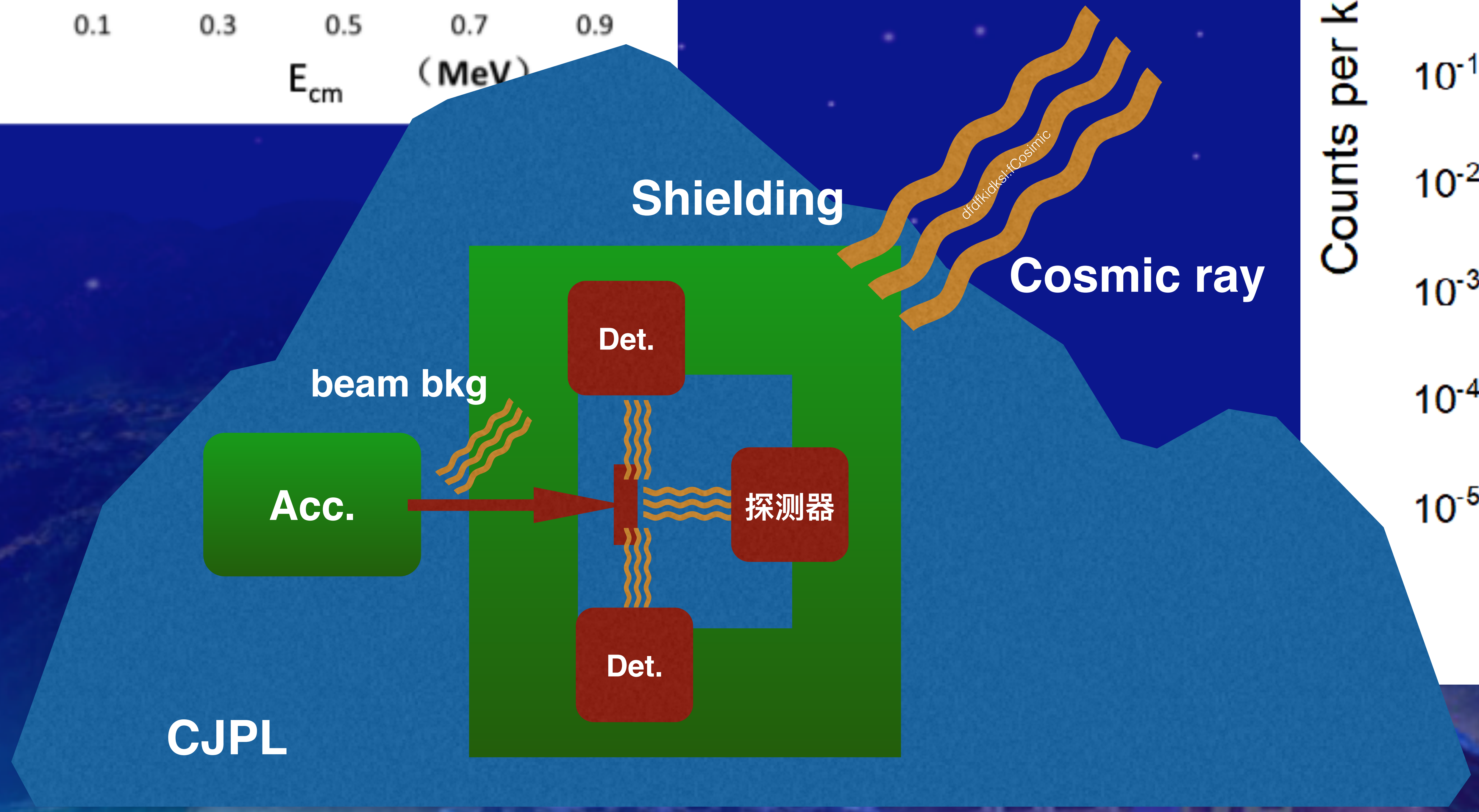
Gamow window

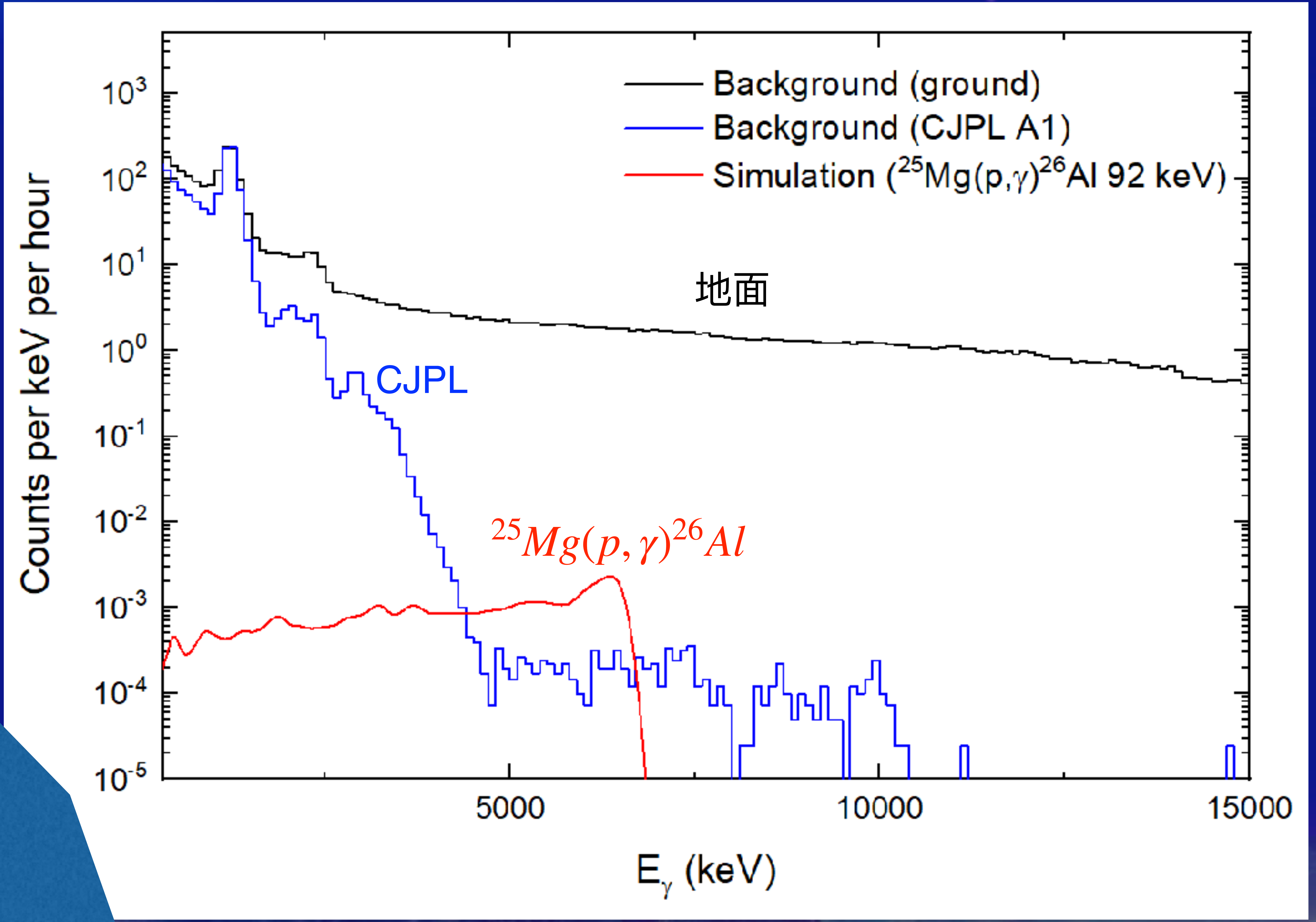




low bkg+high intensity

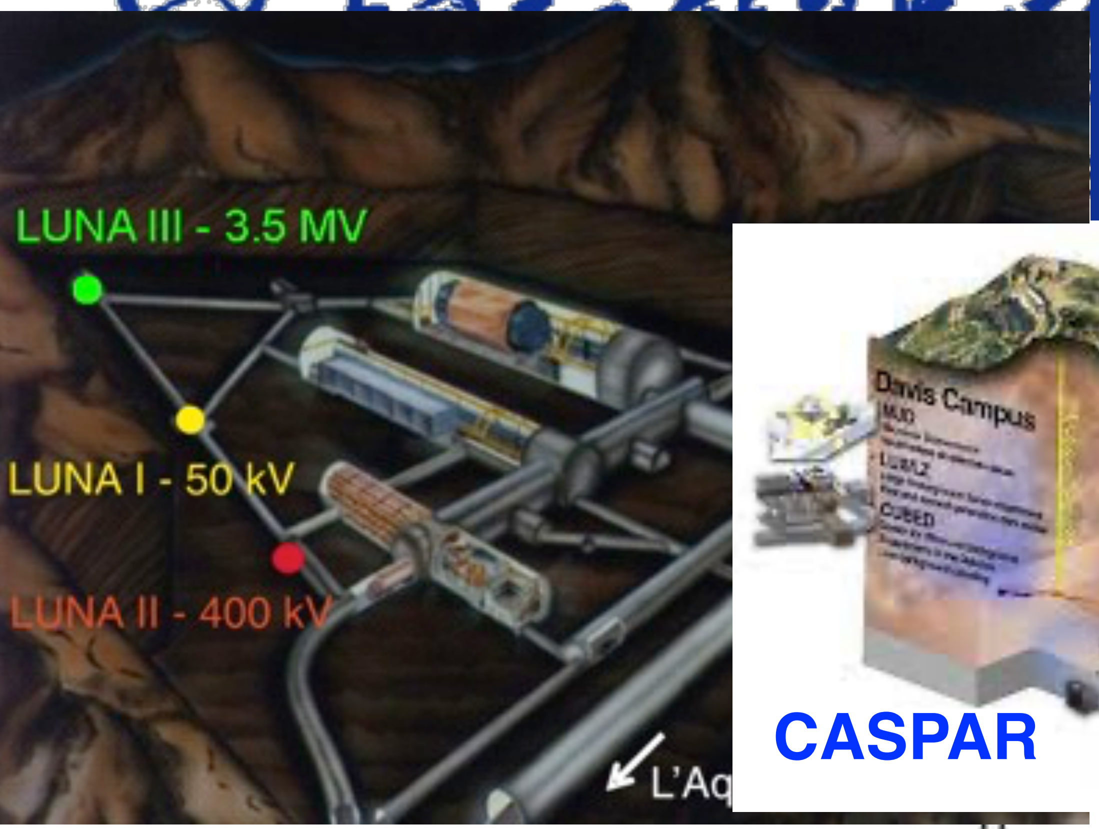


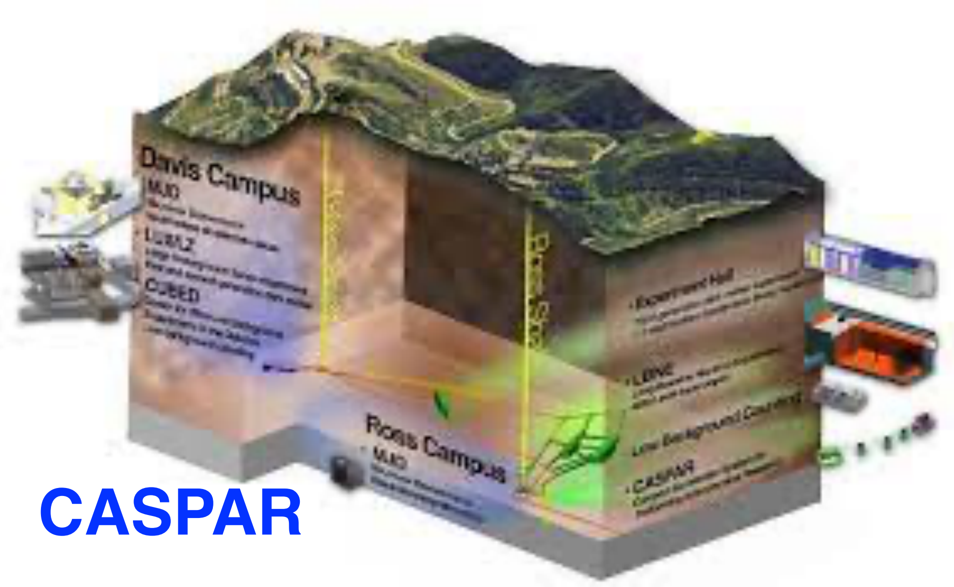


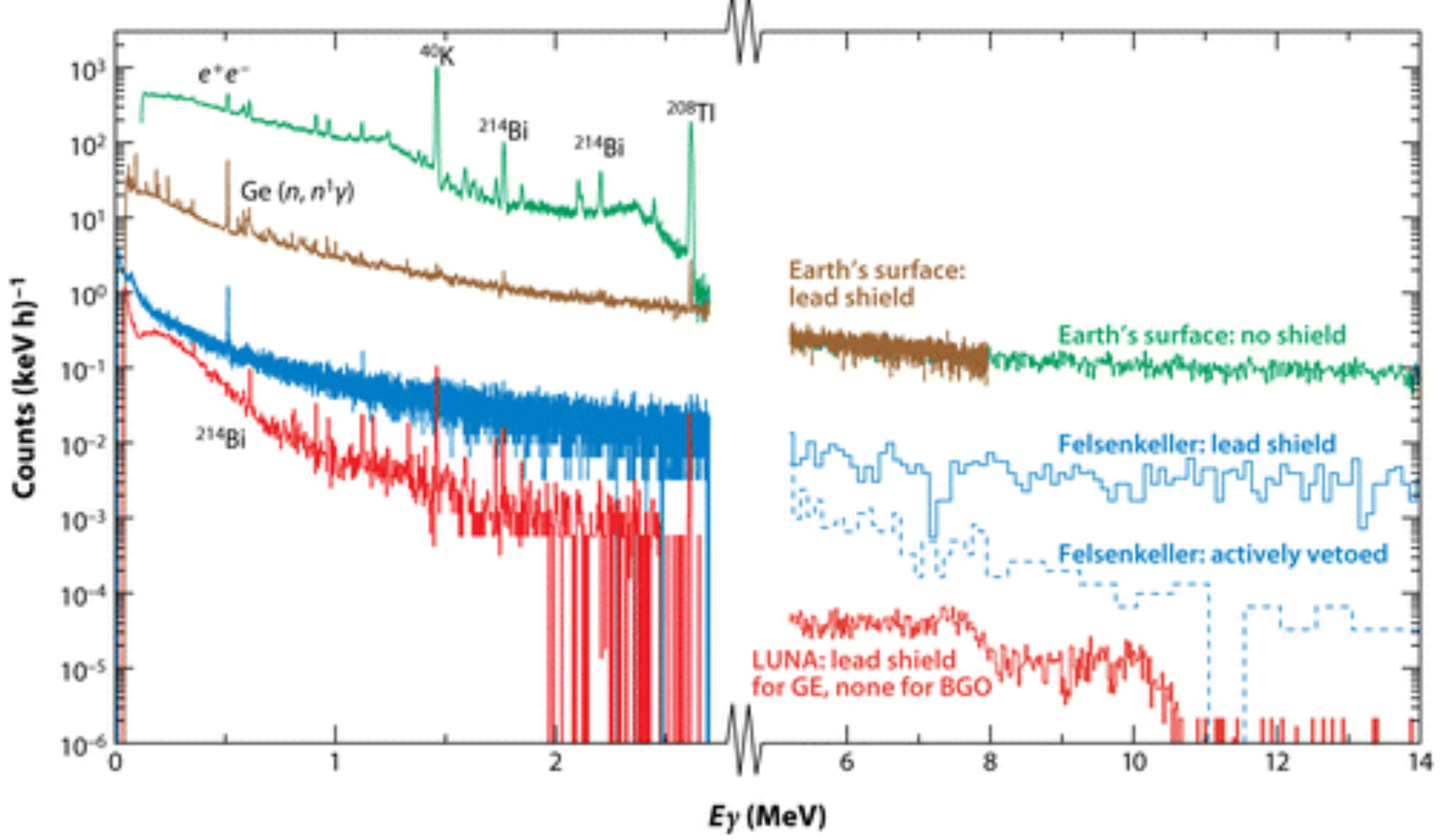


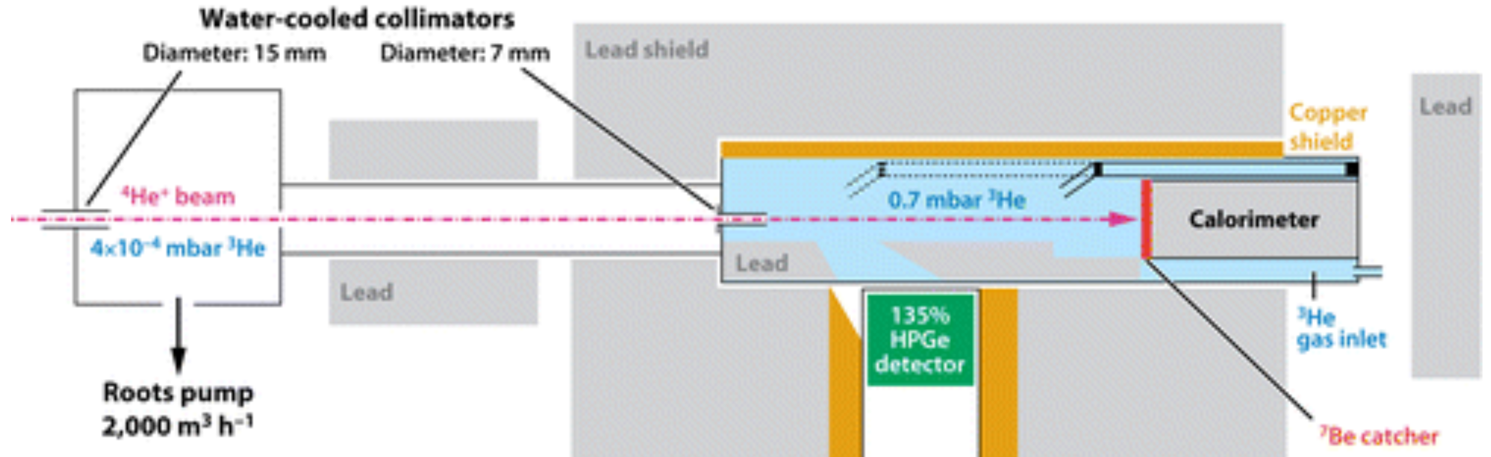
WPL et al., Sci. China 59(2016)2

LUNA, CASPAR and Felsenkeller





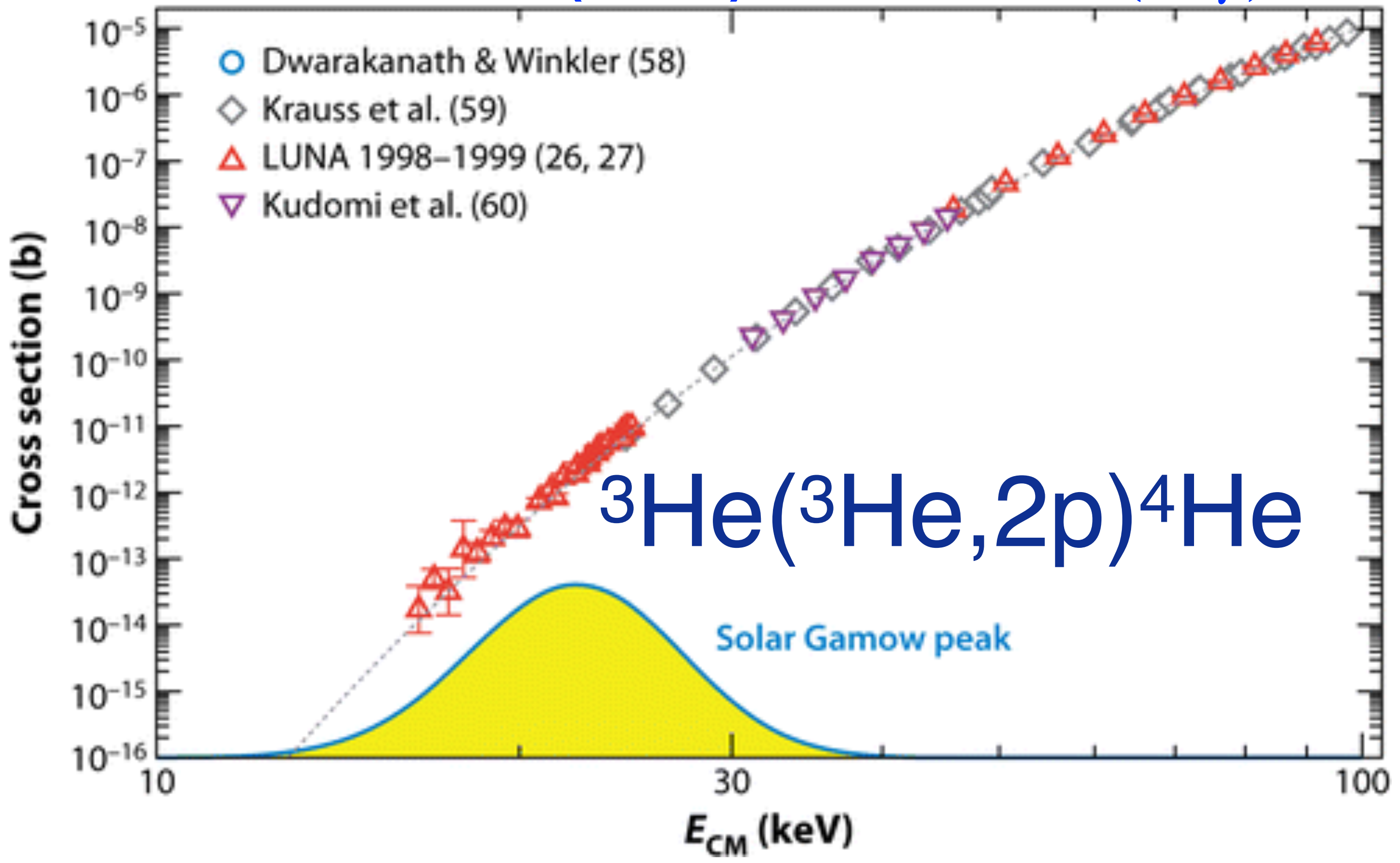


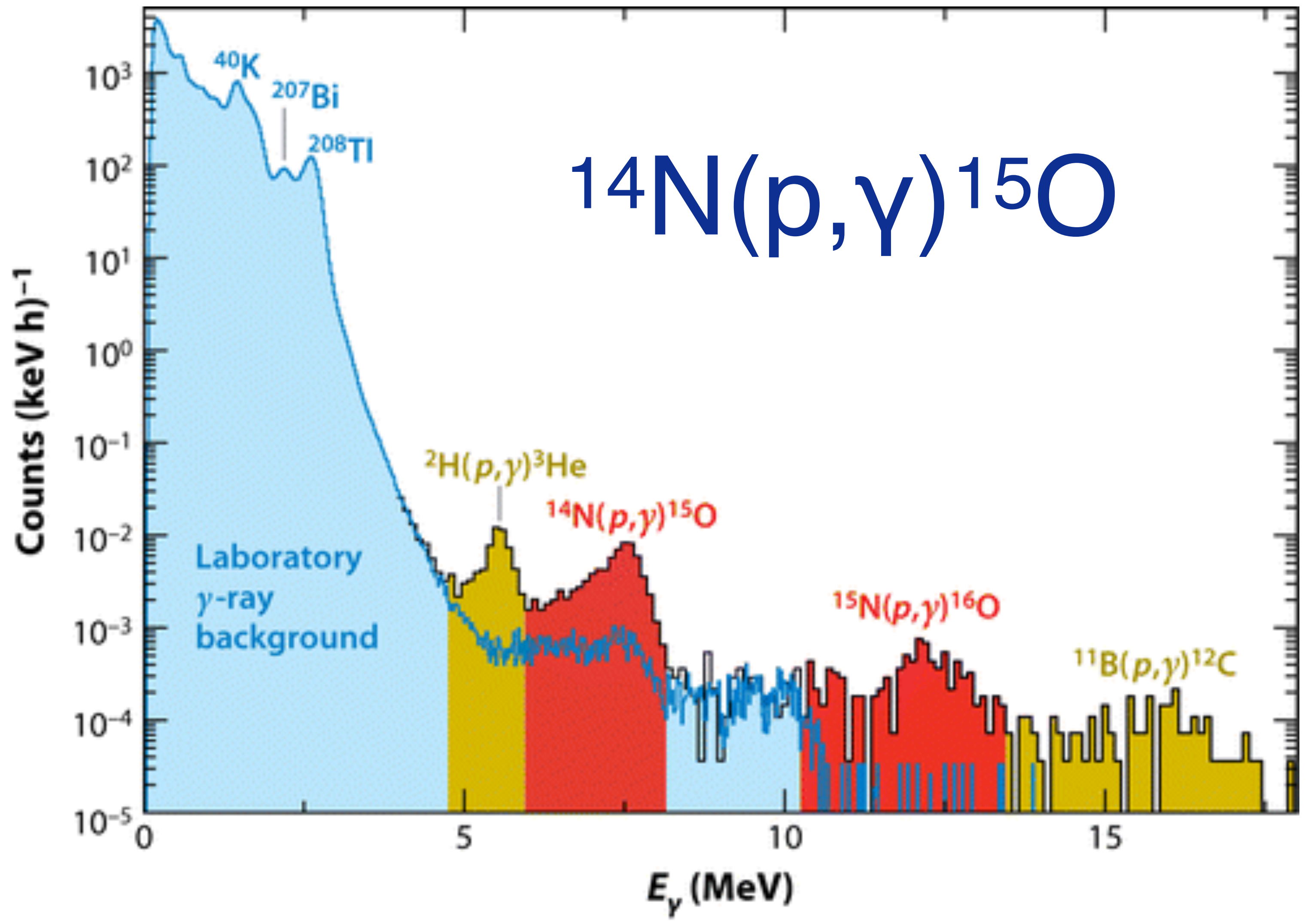


Broggini C, et al. 2010.
Annu. Rev. Nucl. Part. Sci. 60:53–73

CASPAR

F. Cavanna et al., PRL 115(2015)252501, $^{22}Ne(p,\gamma)^{23}Na$.
F. Ciani et al. PRL 127(2021)152701, $^{13}C(\alpha,n)^{16}O$
V. Mossa et al., Nature 587(2020)210 , $D(p,\gamma)^3He$
A. C. Dombos et al., PRL 128(2022)162701, $^{18}O(\alpha,\gamma)^{22}Ne$





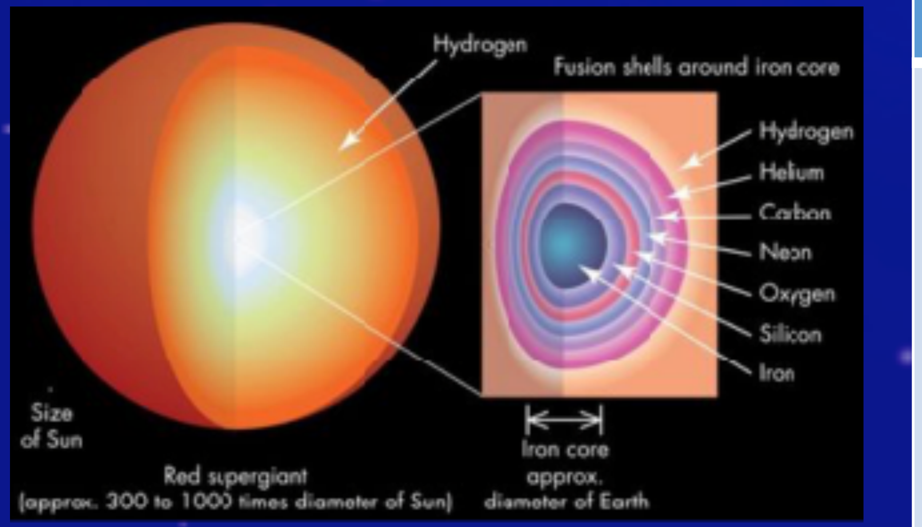
LUNA

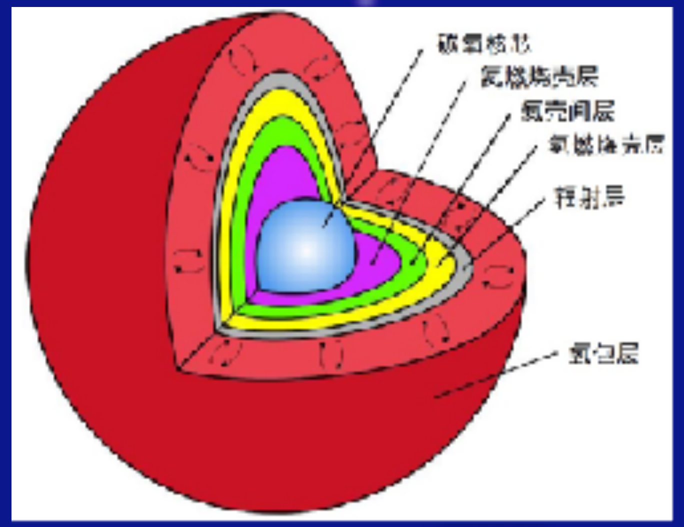
^{3}He(^{3}He,2p)^{4}He PRL82(1999)5205
^{2}H(^{3}He,p)^{4}He PLB482(2000)43
^{2}H(p,γ)^{3}He NPA 706(2002)203
^{3}He(α,γ)^{7}Be PRL 97(2006)122502
^{14}N(p,γ)^{15}O PLB 591(2004)61
^{15}N(p,γ)^{16}O PRC82, 055804(2010)
^{17}O(p,γ)^{18}F PRL 109, 202601(2012)
^{25}Mg(p,γ)^{26}Al PLB 707(2012) 60

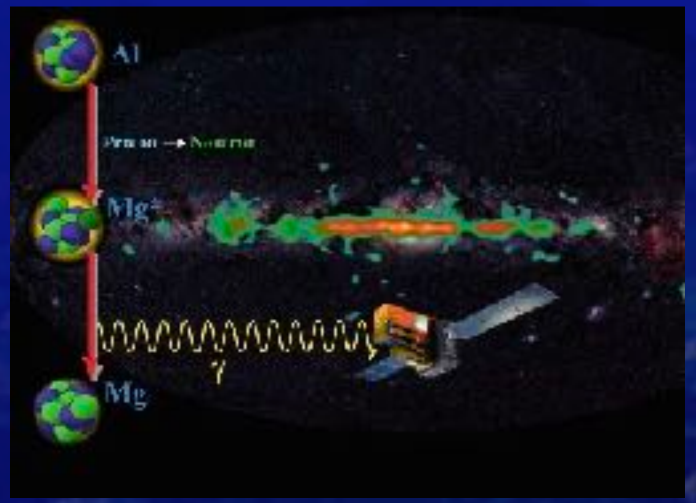
R. M. Gesuè et al., PRL 133(2024)052701, $^{17}O(p,\gamma)^{18}F$

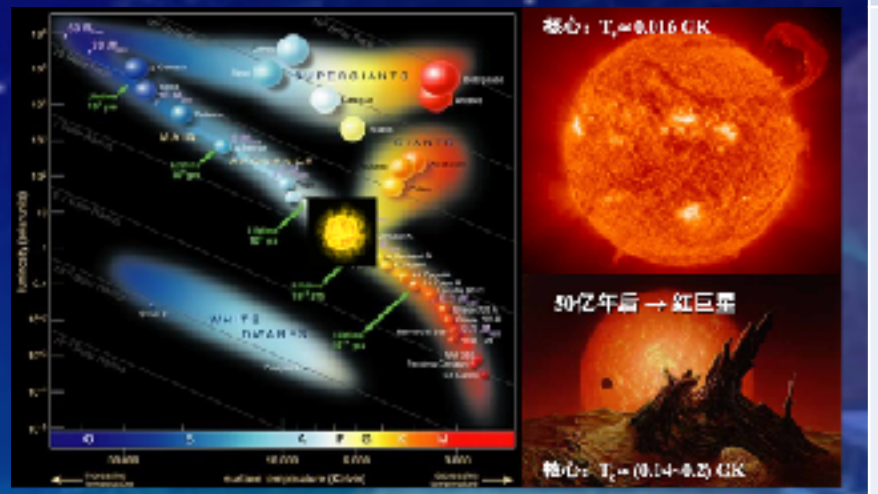
Felsenkeller
^{3}He(α,γ)^{7}Be, progress
^{2}H(p, γ)^{3}He, progress
^{12}C(p,γ)^{13}N, progress
^{12}C(α, γ)^{16}O, plan

Uncertainty remained for key reactions

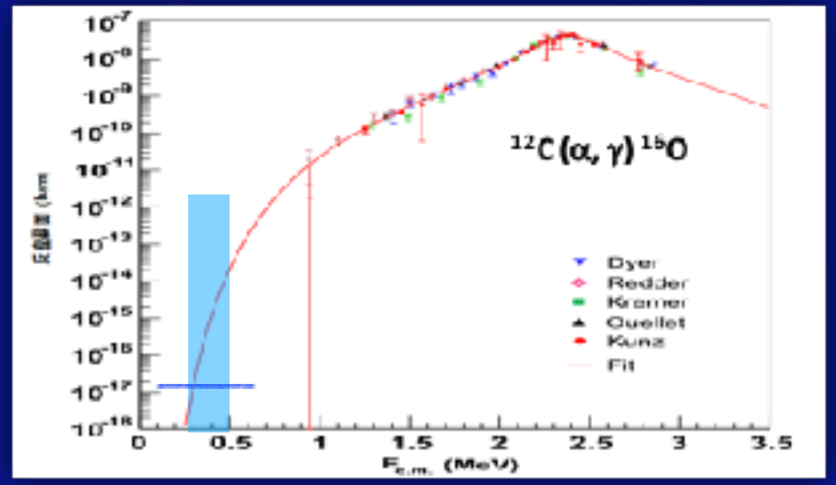




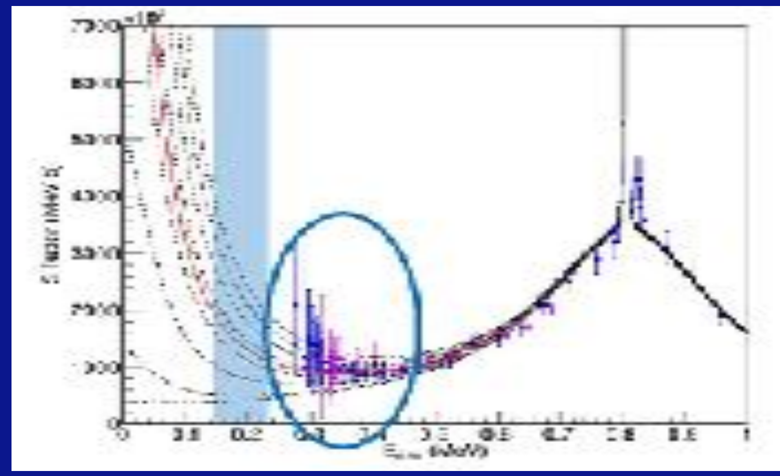




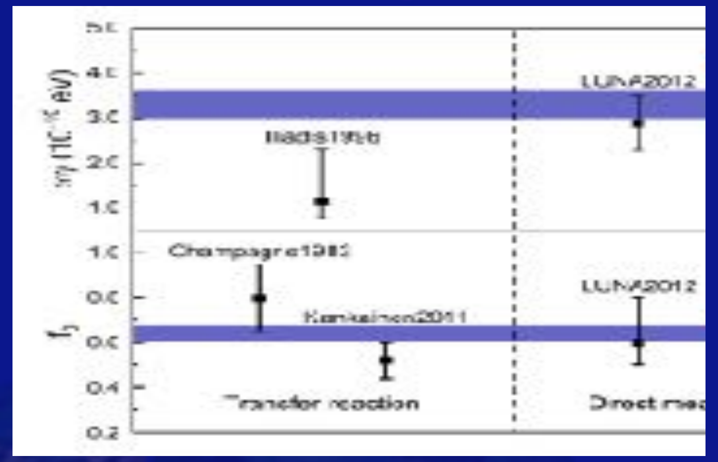
Physics	Reaction	Current	Desired
Massive star	$^{12}C(\alpha,\gamma)^{16}O$	60% 890 keV	20% 220-380 keV
s-process neutron source	$^{13}C(\alpha,n)^{16}O$	60% 230 keV	10% 140-230 keV
Galaxy ^{26}Al source	$^{25}Mg(p,\gamma)^{26}Al$	20% 92 keV	5% 50-300 keV
F aboundace	$^{19}F(p,\alpha)^{16}O$	80 % 189 keV	5 % 50-250 keV



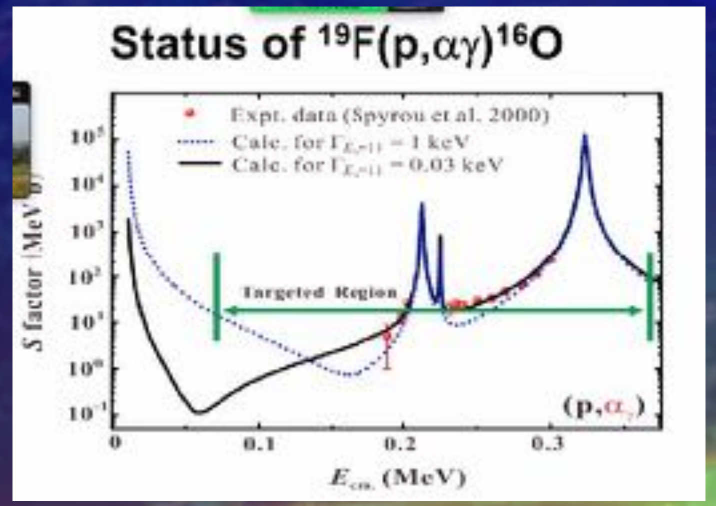
R. J. deBoer et al., RMP vol. 89, 2017



Y. P. Shen, B. Guo, WPL, PPNP 119(2021)103857



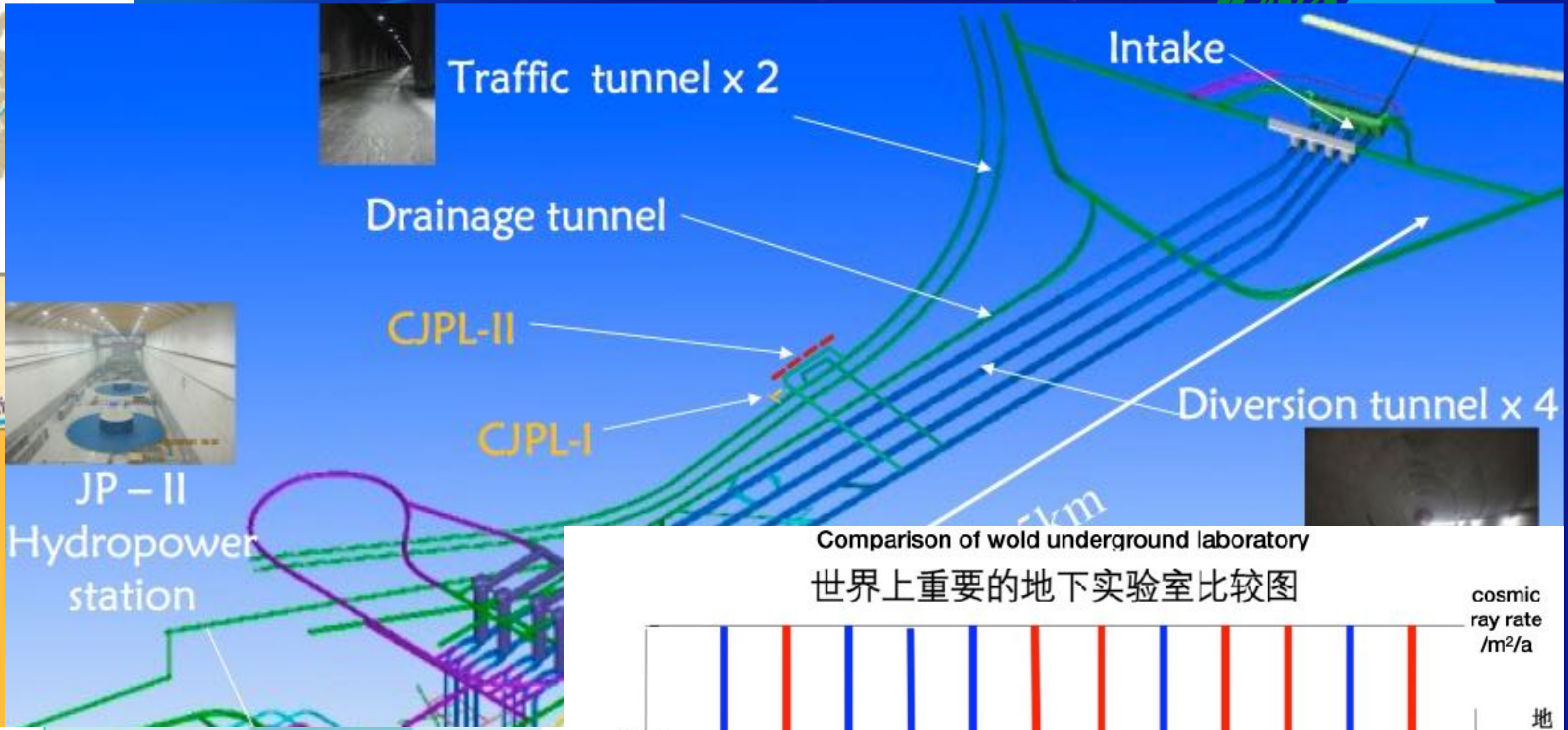
G.F. Ciani et al. PRL 127(2021)152701

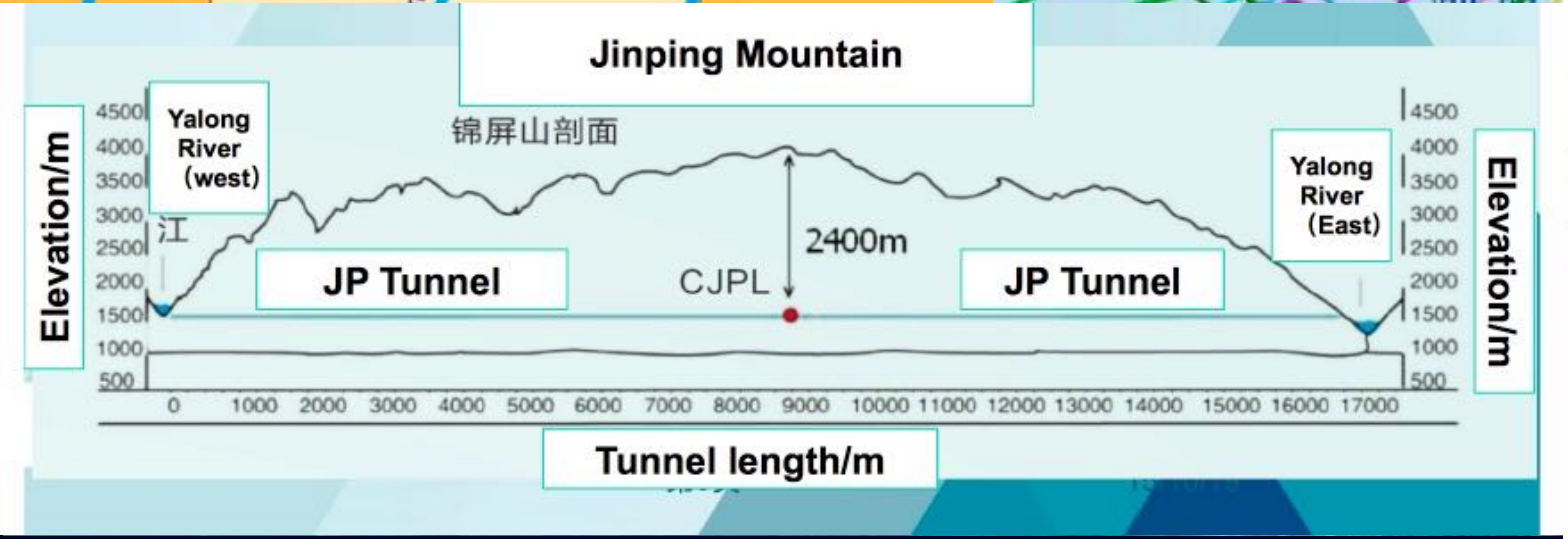


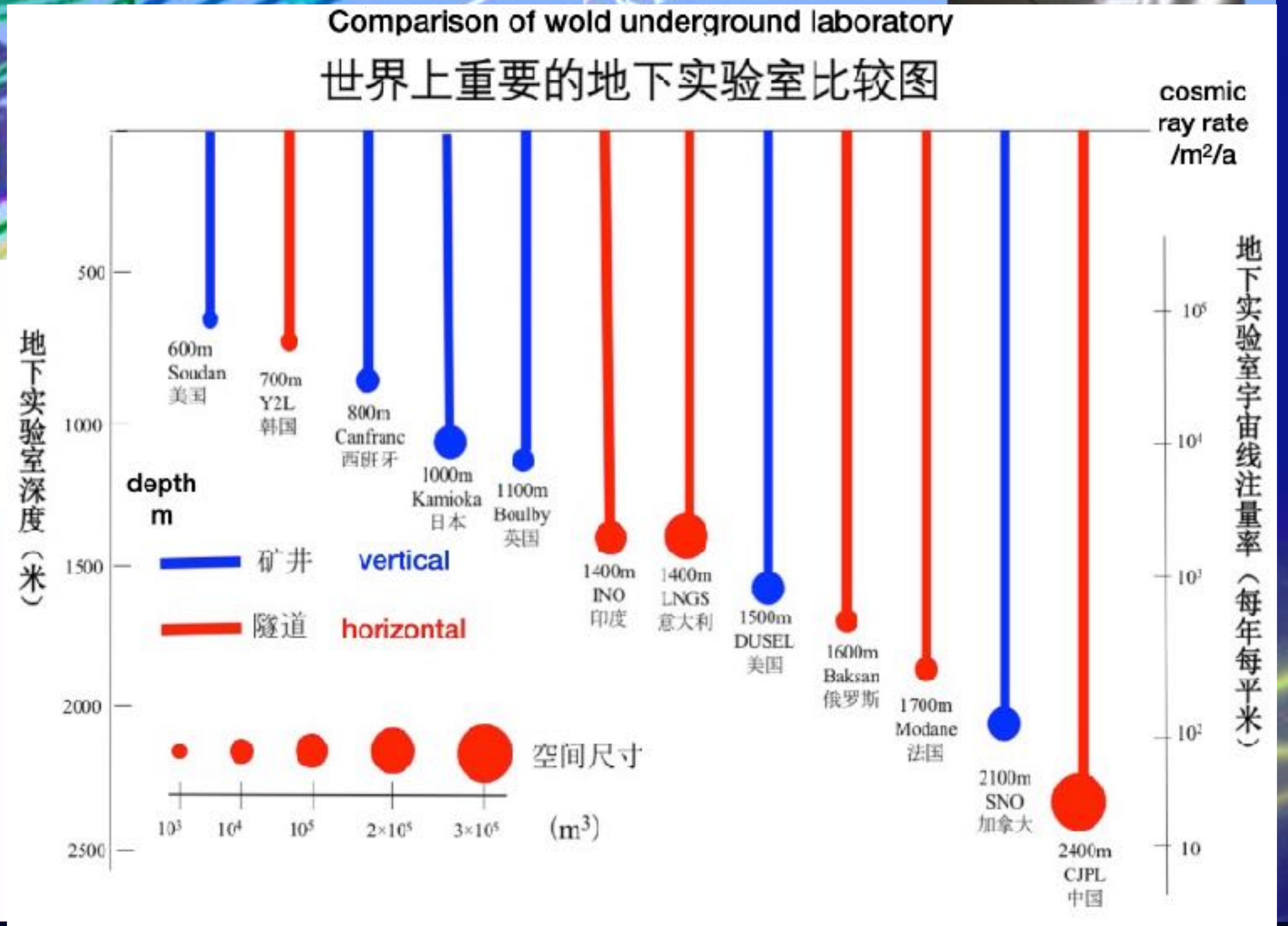
J. J. He et al., Sci. China Phys 59 (2016) 652001

WPL et al., Sci. China 59(2016)2

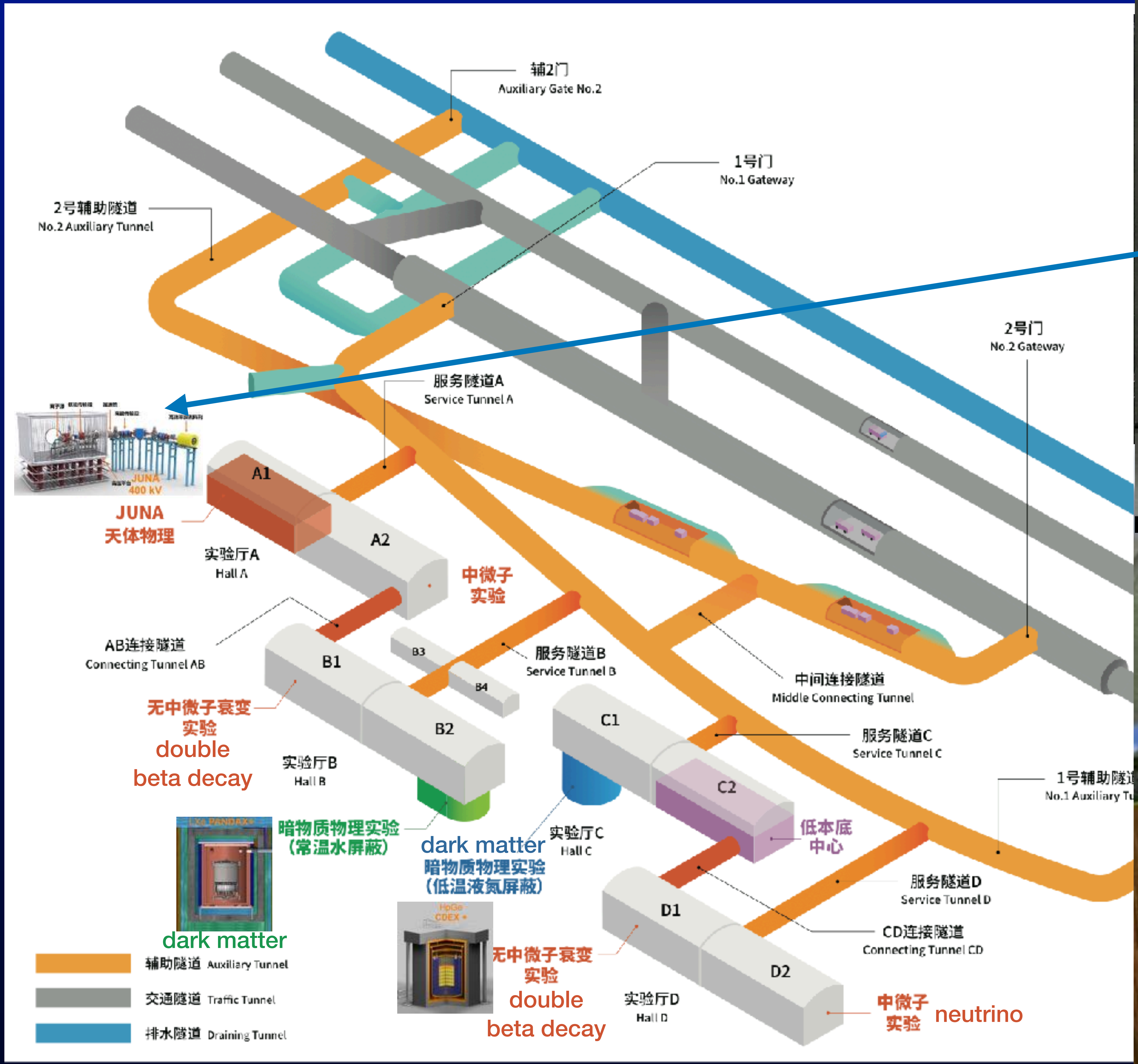
China Jinping: CJPL







Most silent location: CJPL



JUNA dream team

JUNA

Group leader

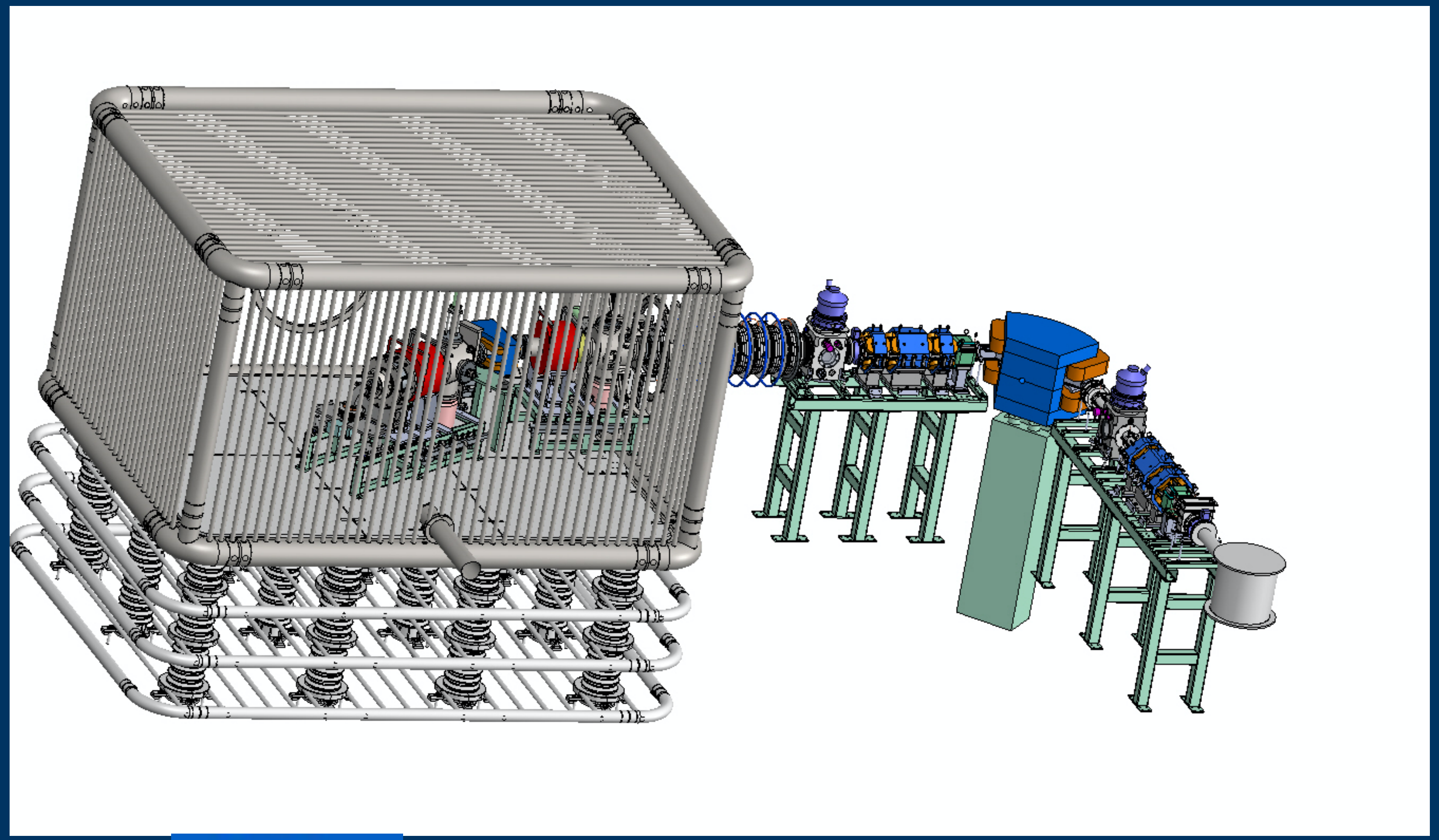
Weiping Liu

$^{12}C(\alpha,\gamma)^{16}O$

Yangping Shen, CIAE

Jun Su, BNU

PI



Bing Guo

$^{22}Ne(\alpha,n)^{25}Mg$

CIAE

Shuo Wang

$^{14}N(p,\gamma)^{15}O$

SDU

Xiaodong Tang

$^{13}C(\alpha,n)^{16}O$

Ion source IMP

Zhihong Li

$^{25}Mg(p,\gamma)^{26}Al$

CIAE

Jun Su, BNU

Jianjun He

$^{19}F(p,\alpha)^{16}O$

BNU

Gang Lian

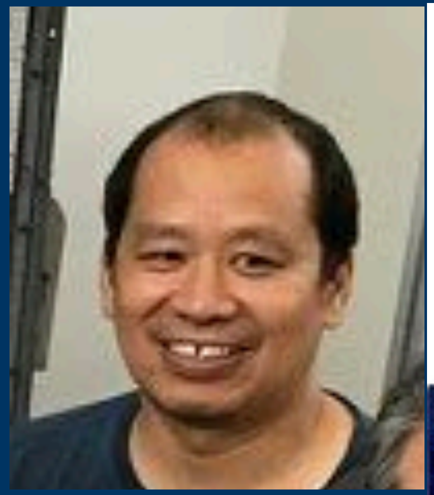
Lab. exp. sup.

CIAE

中国原子能科学研究院 CHINA INSTITUTE OF ATOMIC ENERGY

中国科学院近代物理研究所 Institute of Modern Physics, Chinese Academy of Sciences

SDIC 国投集团 雅砻江流域水电开发有限公司 YALONG RIVER HYDROPOWER DEVELOPMENT COMPANY, LTD.



Bao Quncui,CIAE

Liangting Sun, IMP

Ion source and acc.

Acc. installation

Arjun Li

A1 construction

Hongwei Yang

Site support

Xiaopan Cheng

Acc. operation

Long Zhang

Supported by the National Natural Science Foundation of China, Grant No. 11490560, 2015

WPL et al., Sci. China 59(2016)2

JUNA Milestone

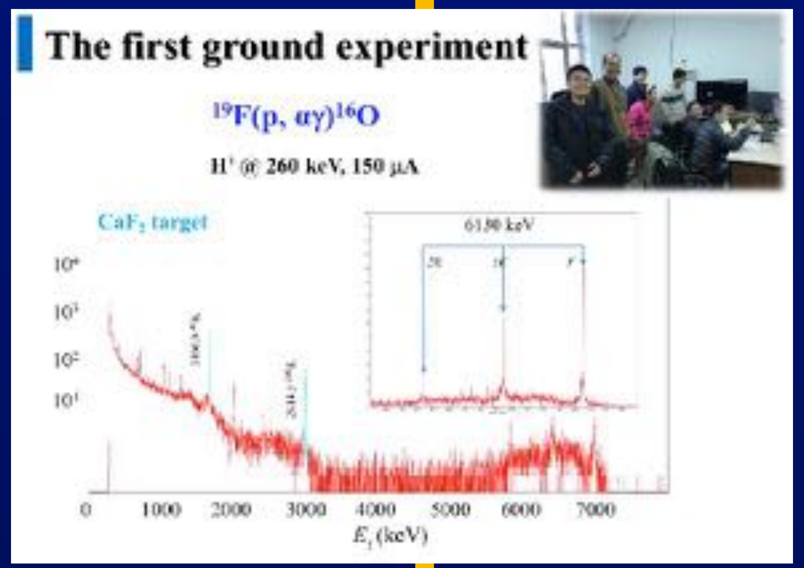
Aug. 2013
Startup meeting

Jan. 2015
Project inauguration

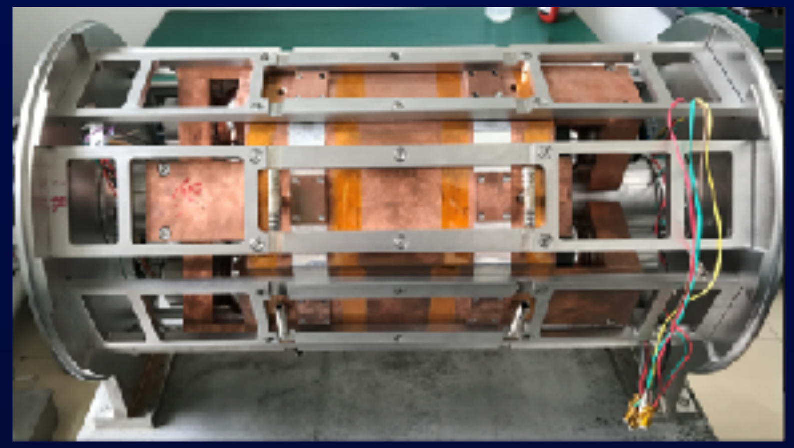
Mar. 2016
On site start

May 2017
Beam on ground

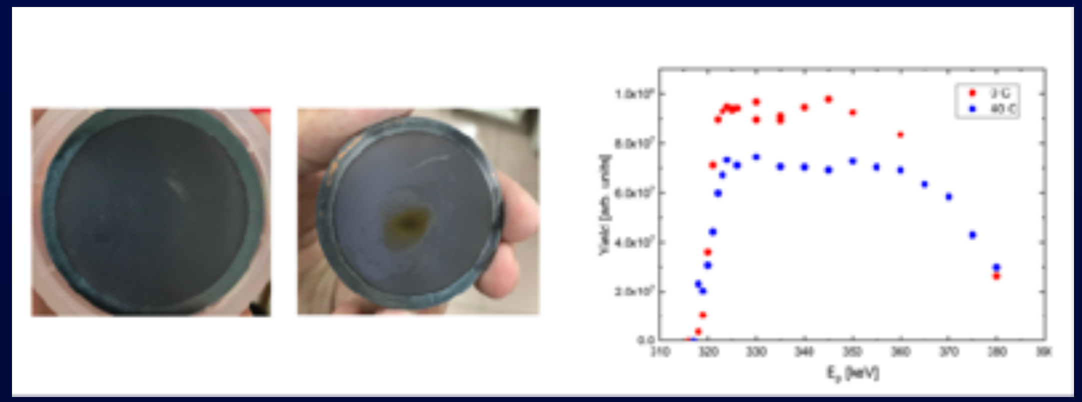
Dec. 2017
3 mA on ground



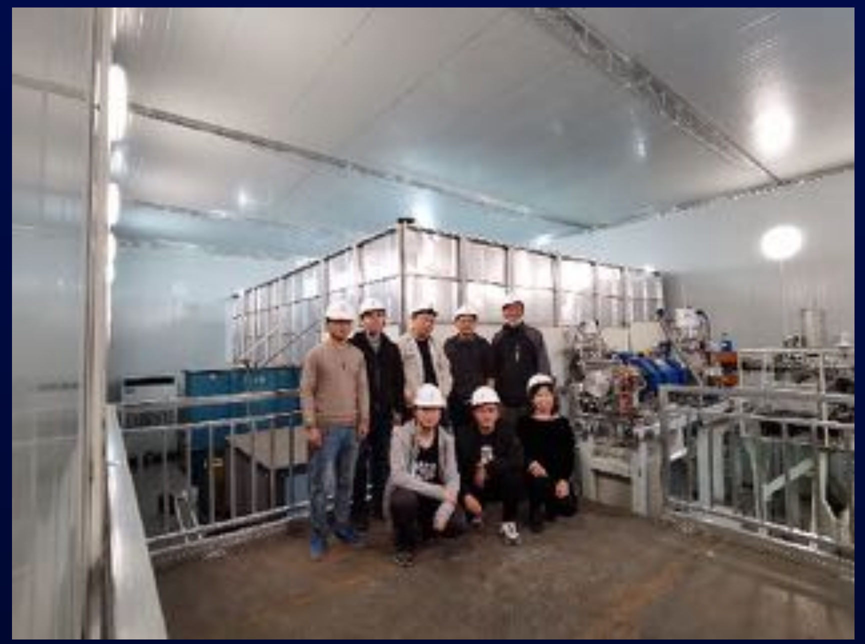
Dec. 2018
Der. Ready
Beam 10 mA



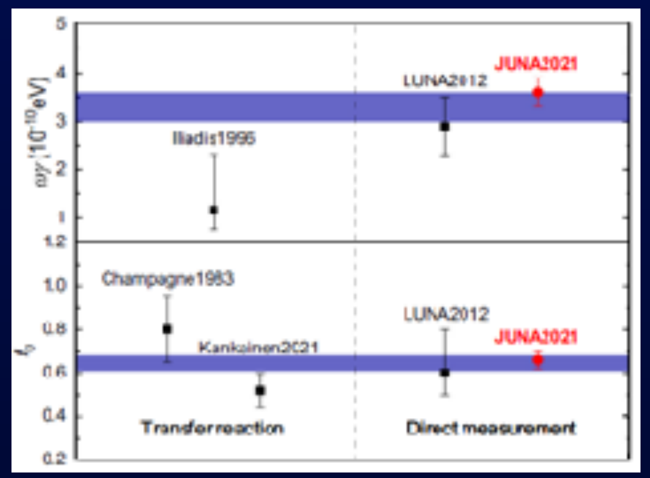
April 2019
Target ready
Acc. Ready



Dec. 2020
Beam underground



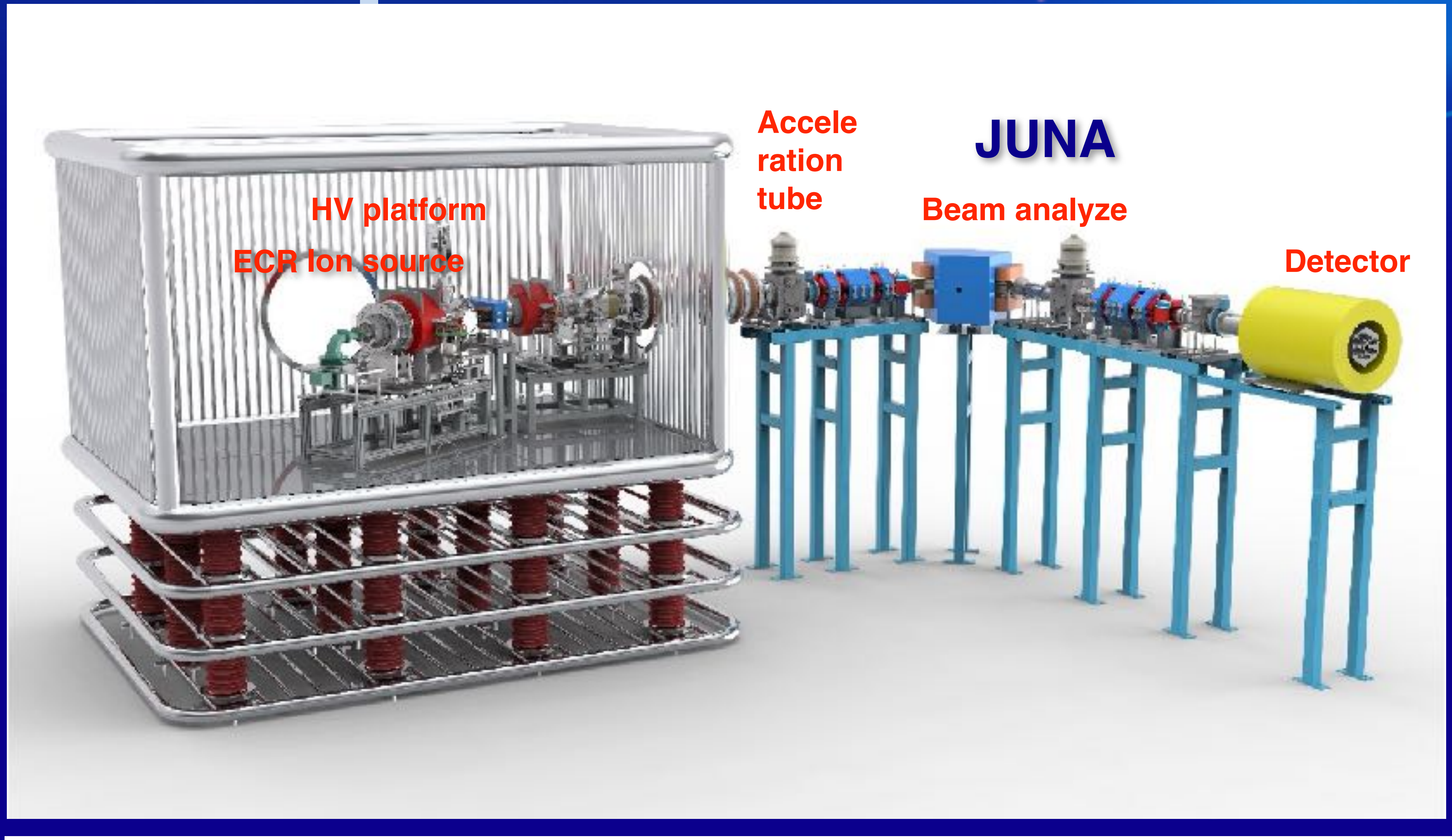
Mar. 2021
^{25}Mg, ^{19}F and ^{13}C data ready



Dec. 2021
Project commission

WPL et al., Sci. China 59(2016)2

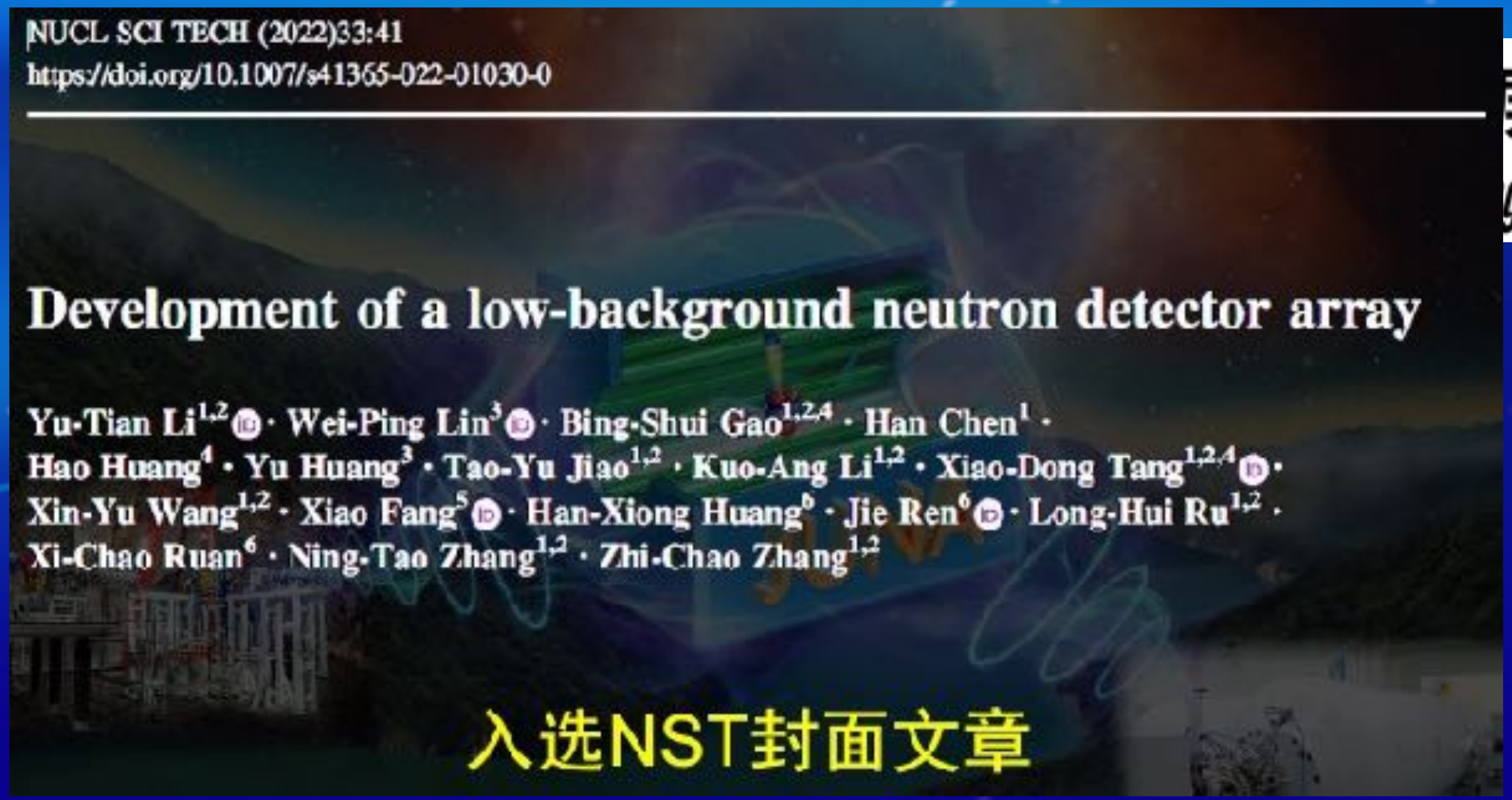
JUNA accelarator



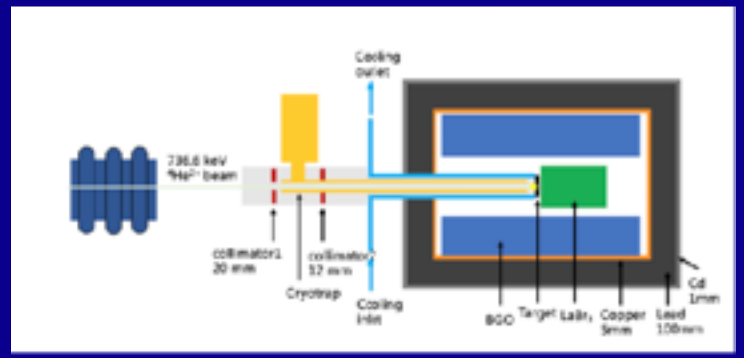
	lab depth m	cosmic μ bkg (cm^{-2} s^{-1})	beam energy (keV)			beam intensity (emA)			energy stability
			H^+	He^+	He^{2+}	H^+	He^+	He^{2+}	
LUNA	1400	2×10^{-8}	50-400	50-400	3.5 MV	0.3~1	0.3~0.8	---	0.05%
CASPAR	1500	4×10^{-9}	100-1000	100-1000	1 MV	0.1	0.1	---	0.05%
JUNA	2400	2×10^{-10}	50-400	50-400	100-800	2-10	2-10	1-2	0.04%
Felsenkeller	45	$\sim10^{-7}$			5 MV		30 uA		

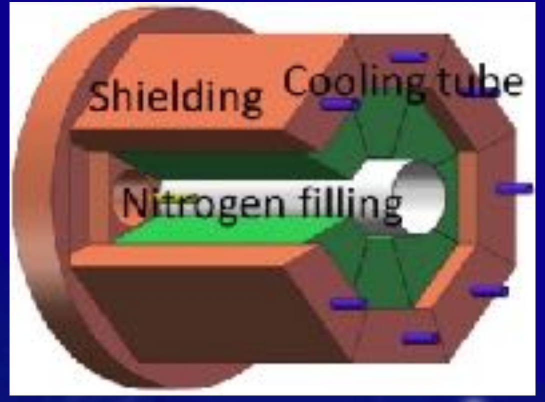
WPL et al., Sci. China 59(2016)2

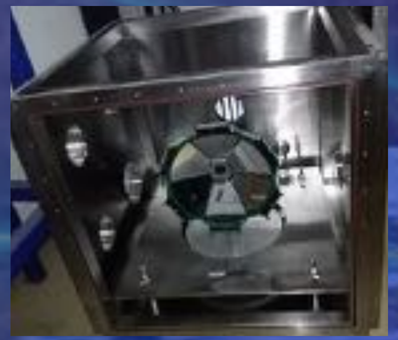
Detector tech.



屏深地核天体物理实验
g Underground Nuclear Astrophysics Experiment







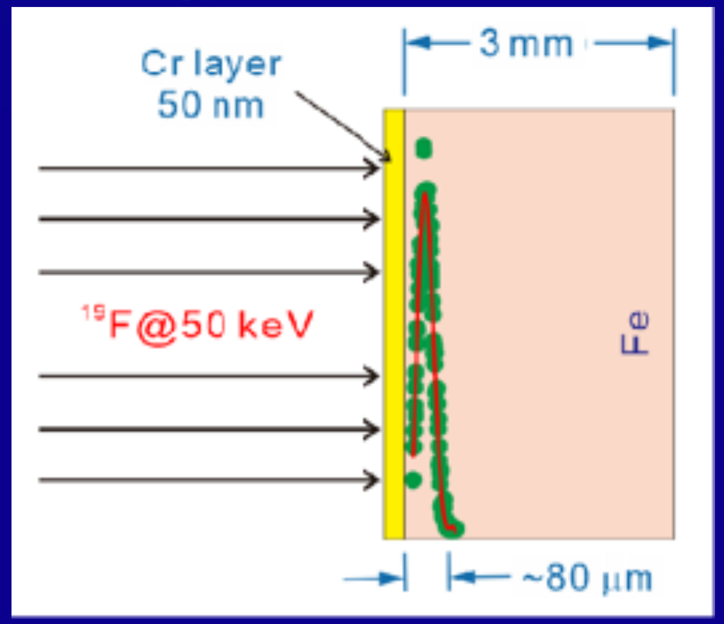
reaction	technology	publication	world best	JUNA
^{12}C	BGO+LaBr		down to 891 keV	down to 552 keV
^{25}Mg	BGO array X8	Atomic ST 52(2018)140	resolution17 %	11 %
^{13}C	^{3}He array X24	NST33(2022) 41, cover story	Exptripolation	Self consisitent
^{19}F	Charged particle array		170 keV	down to 100 keV

High durability target

3-10 times better than previous targets

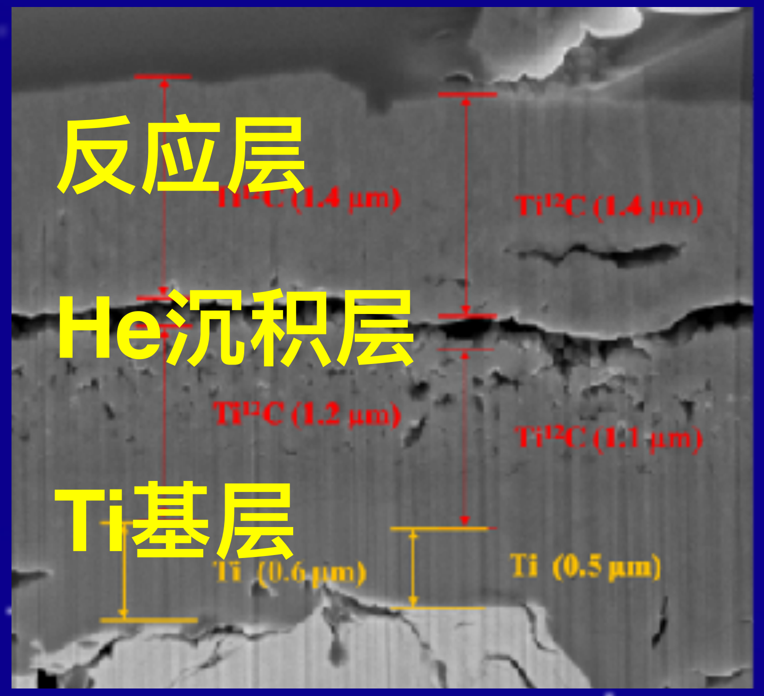
^{19}F Implantation

- High-purity iron substrate
- Magnetron chrome plating
- 100 C



^{12}C Deposit target

- FCVA
- 12C 99.99%
- 400 C

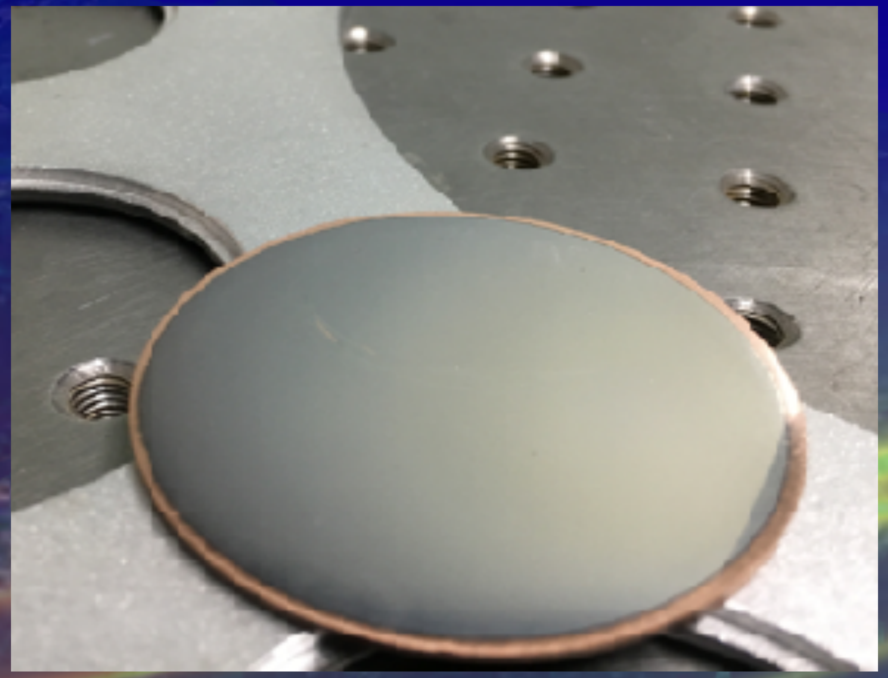


^{13}C Thick target

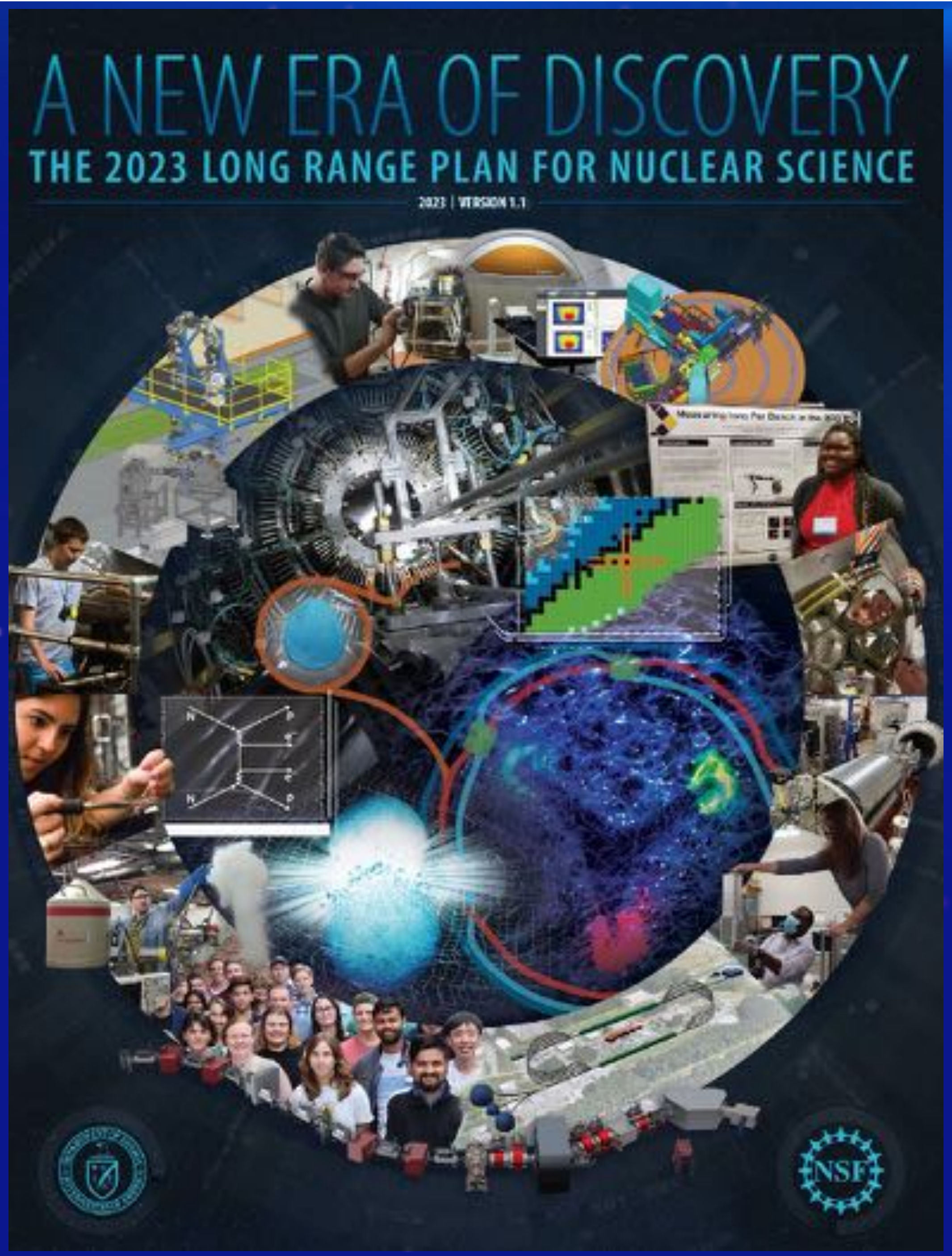
- High-temperature and high-pressure sintering
- 3550°C
- 0.5kW/cm^2

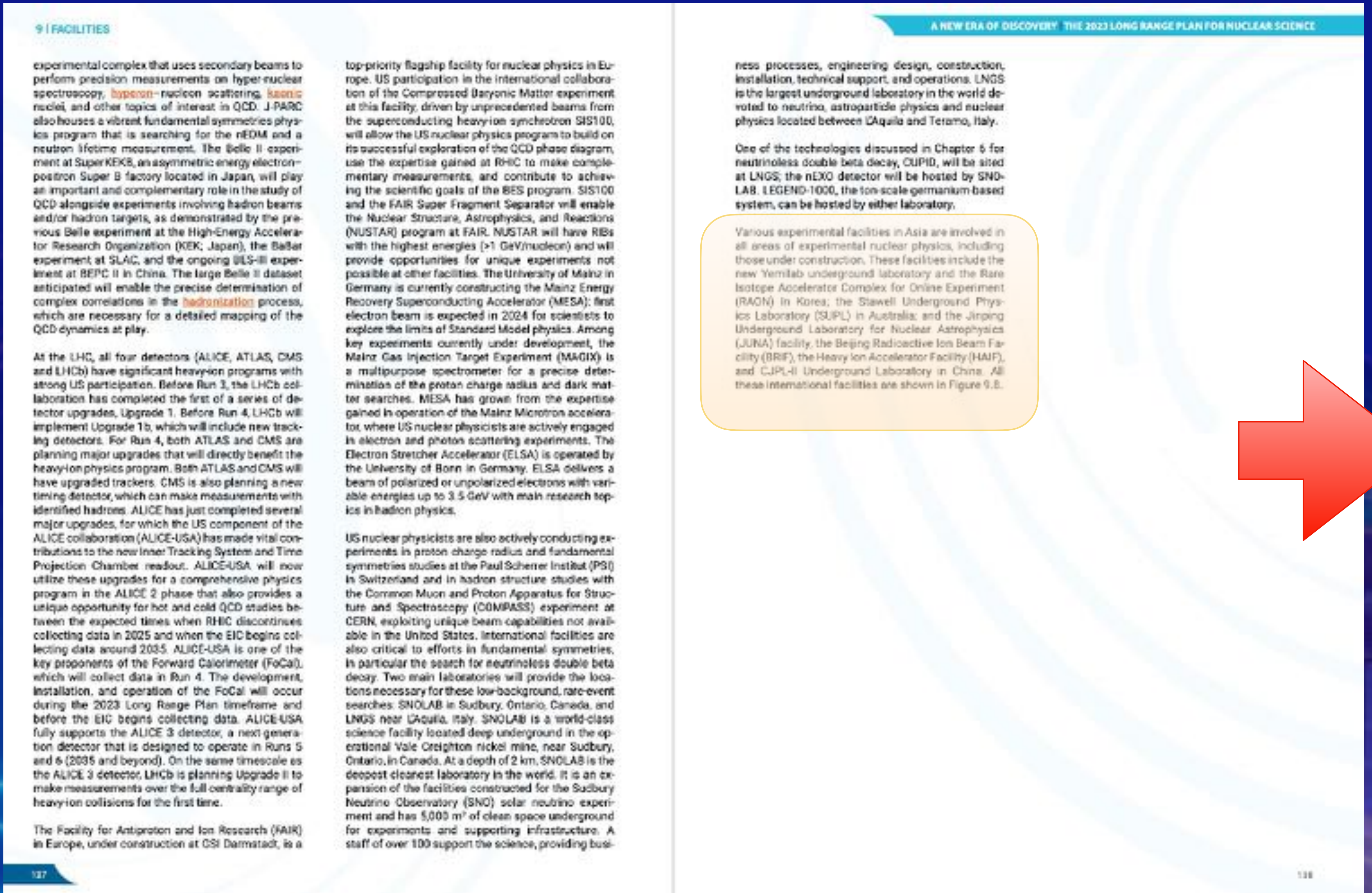
^{25}Mg Hybrid layers

- Cr+Mg+Cr
- rotating coating
- 300 C



JUNA in DOE long range plan 2023

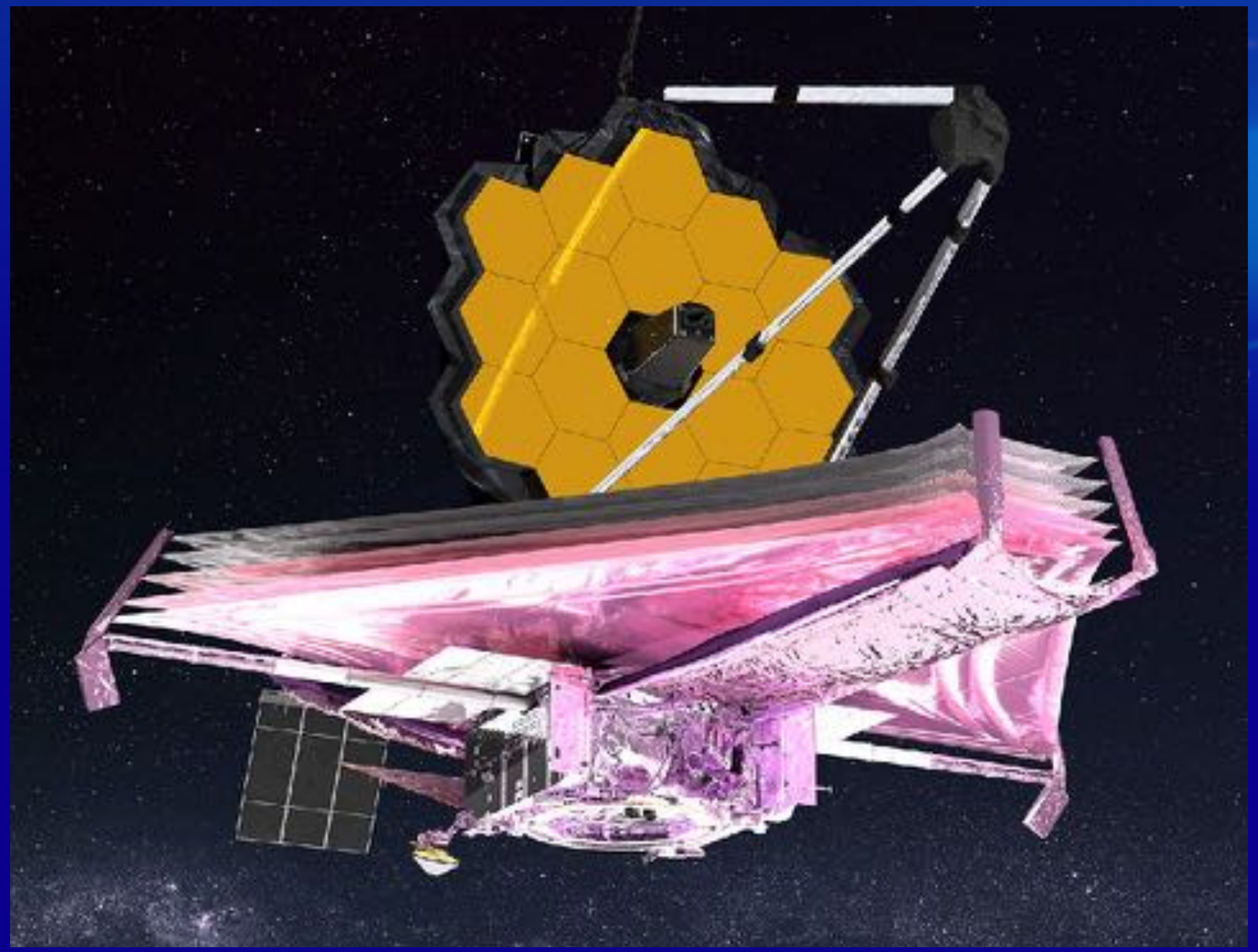




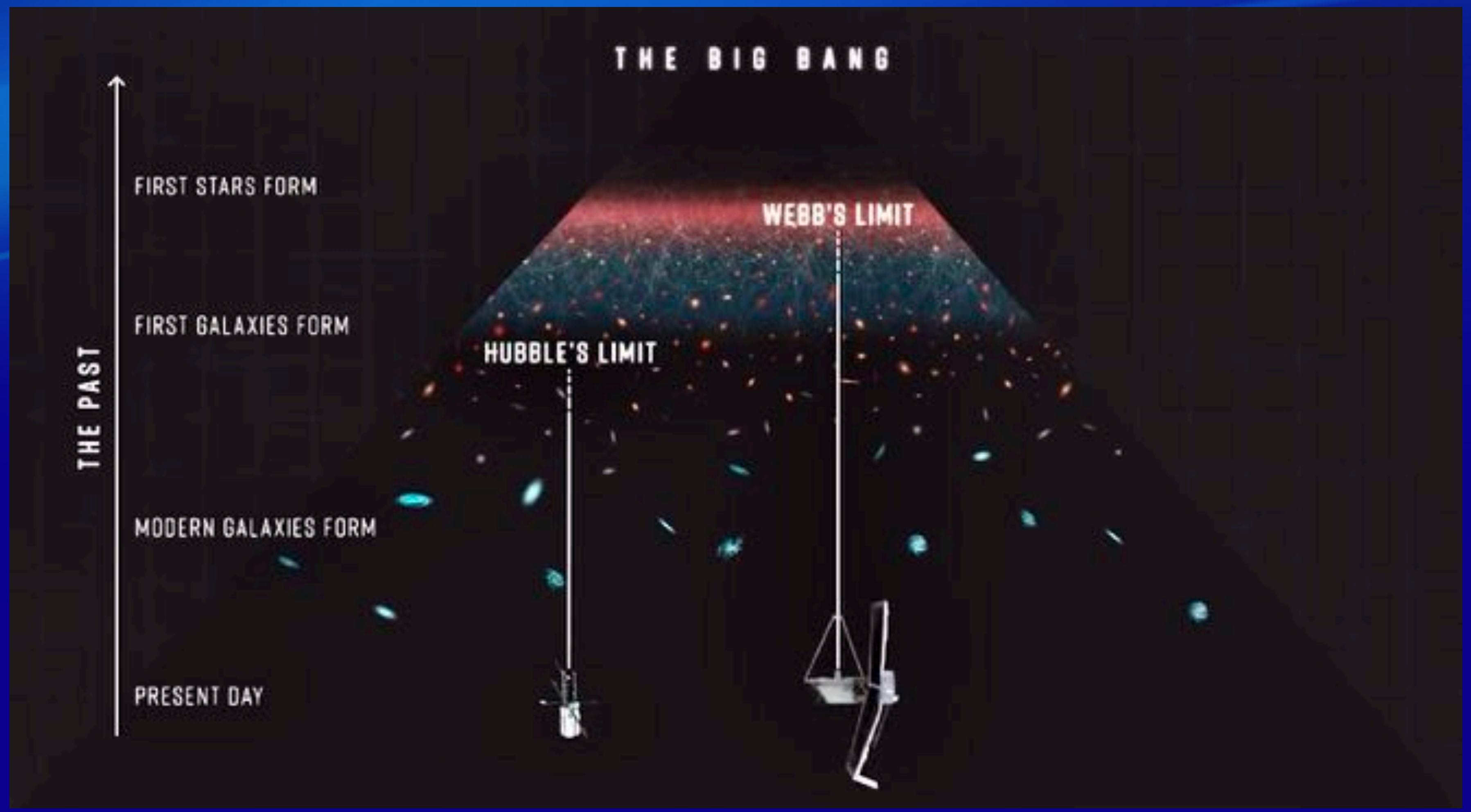
Various experimental facilities in Asia are involved in all areas of experimental nuclear physics, including those under construction. These facilities include the new Yemilab underground laboratory and the Rare Isotope Accelerator Complex for Online Experiment (RAON) in Korea; the Stawell Underground Physics Laboratory (SUPL) in Australia; and the Jinping Underground Laboratory for Nuclear Astrophysics (JUNA) facility, the Beijing Radioactive Ion Beam Facility (BRIF), the Heavy Ion Accelerator Facility (HAIF), and CJPL-II Underground Laboratory in China. All these international facilities are shown in Figure 9.8.

$^{19}F(p,\gamma)^{20}Ne$： confirm CNO break，explain Ca in oldest star

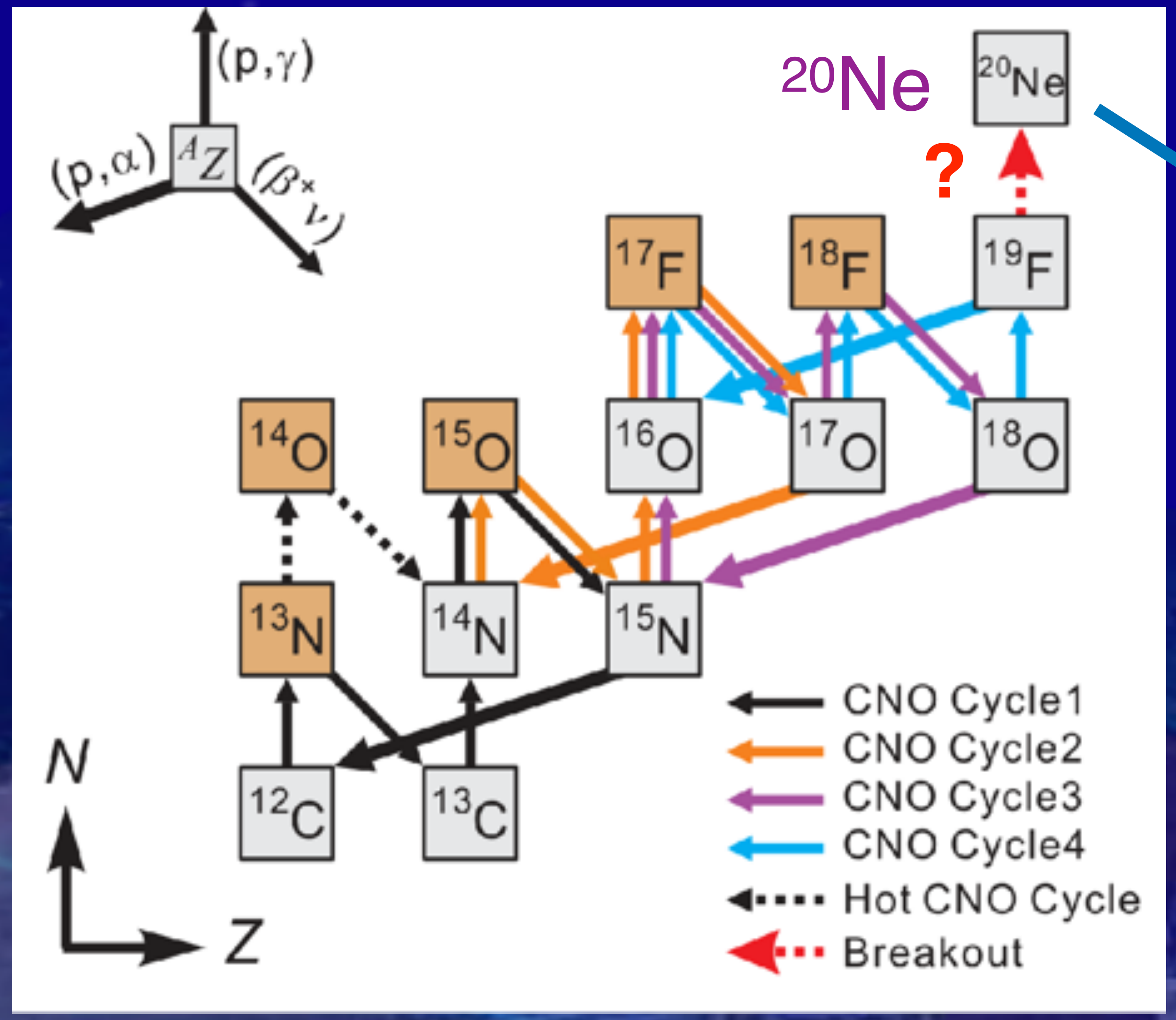
锦屏深地核天体物理实验
Jinping Underground Nuclear Astrophysics Experiment

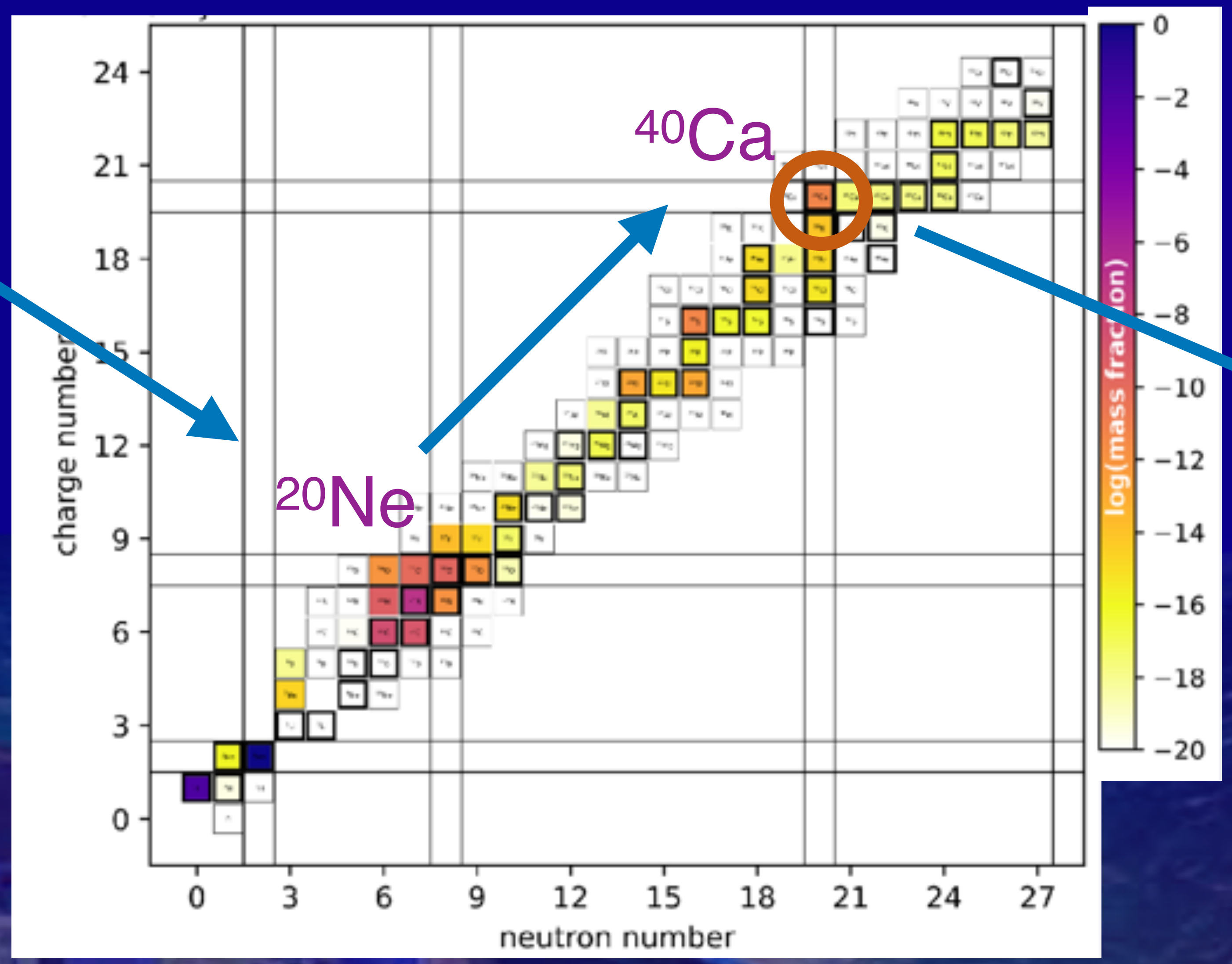


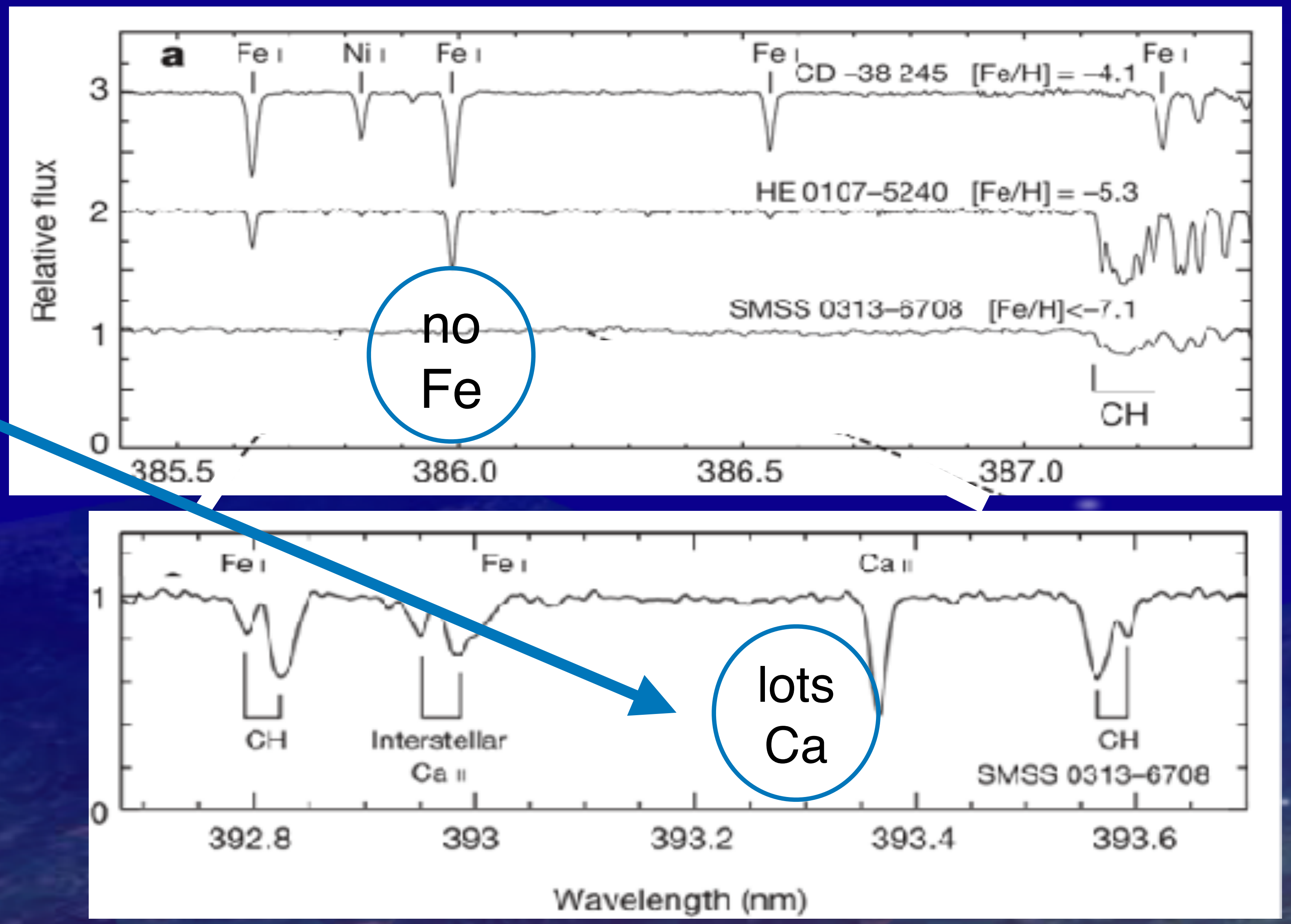
JWST



S.C. Keller et al., Nature 506 (2014) 463







Solution: $^{19}F(p,g)^{20}Ne$ rate one order large?

SMSS 0313-6708：lots of Ca, but no Fe?

Experiment setup

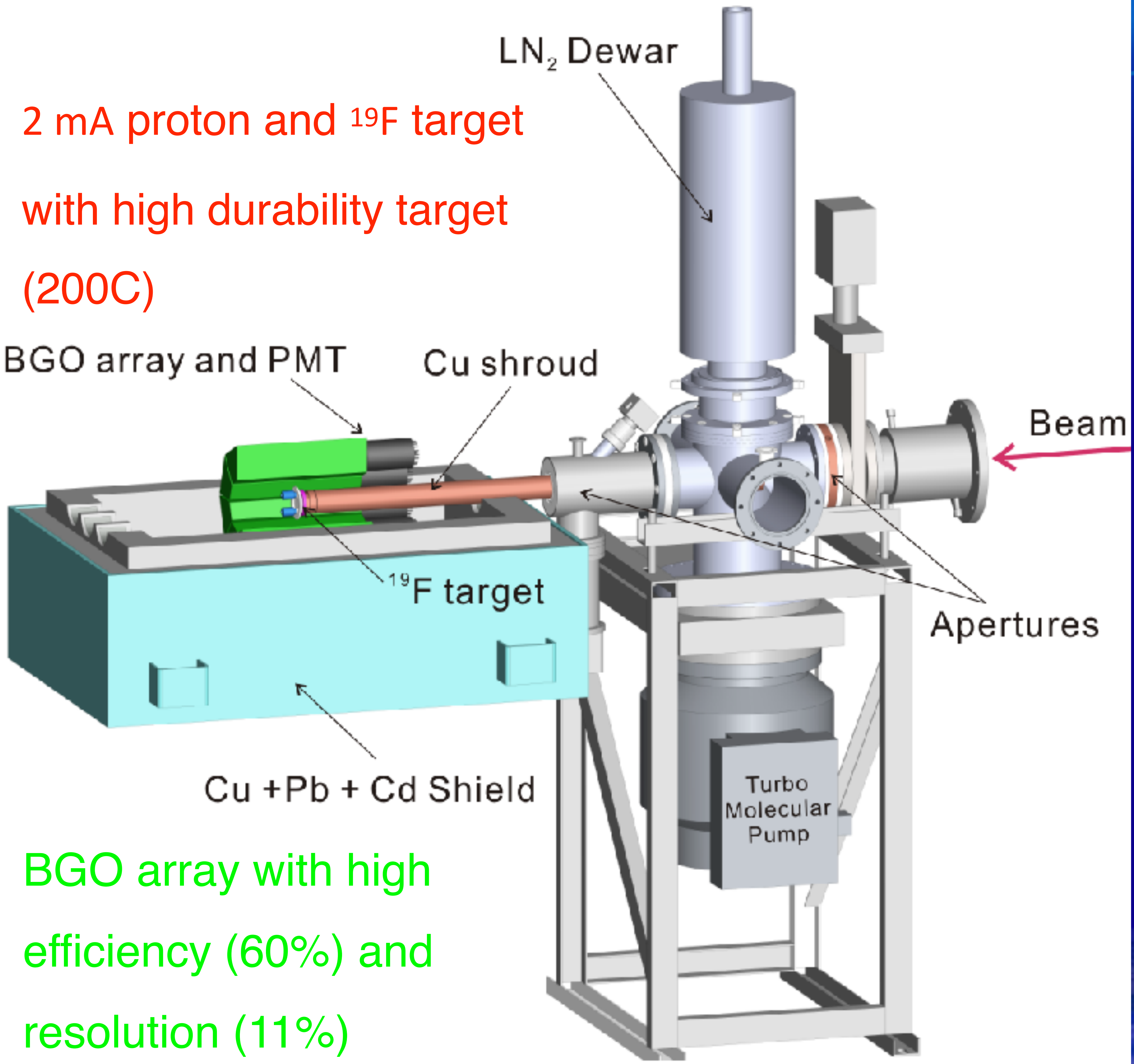
PI: J. J. He, BNU
with L. Y. Zhang, BNU

锦屏深地核天体物理实验
Jinping Underground Nuclear Astrophysics Experiment

2 mA proton and ^{19}F target

with high durability target

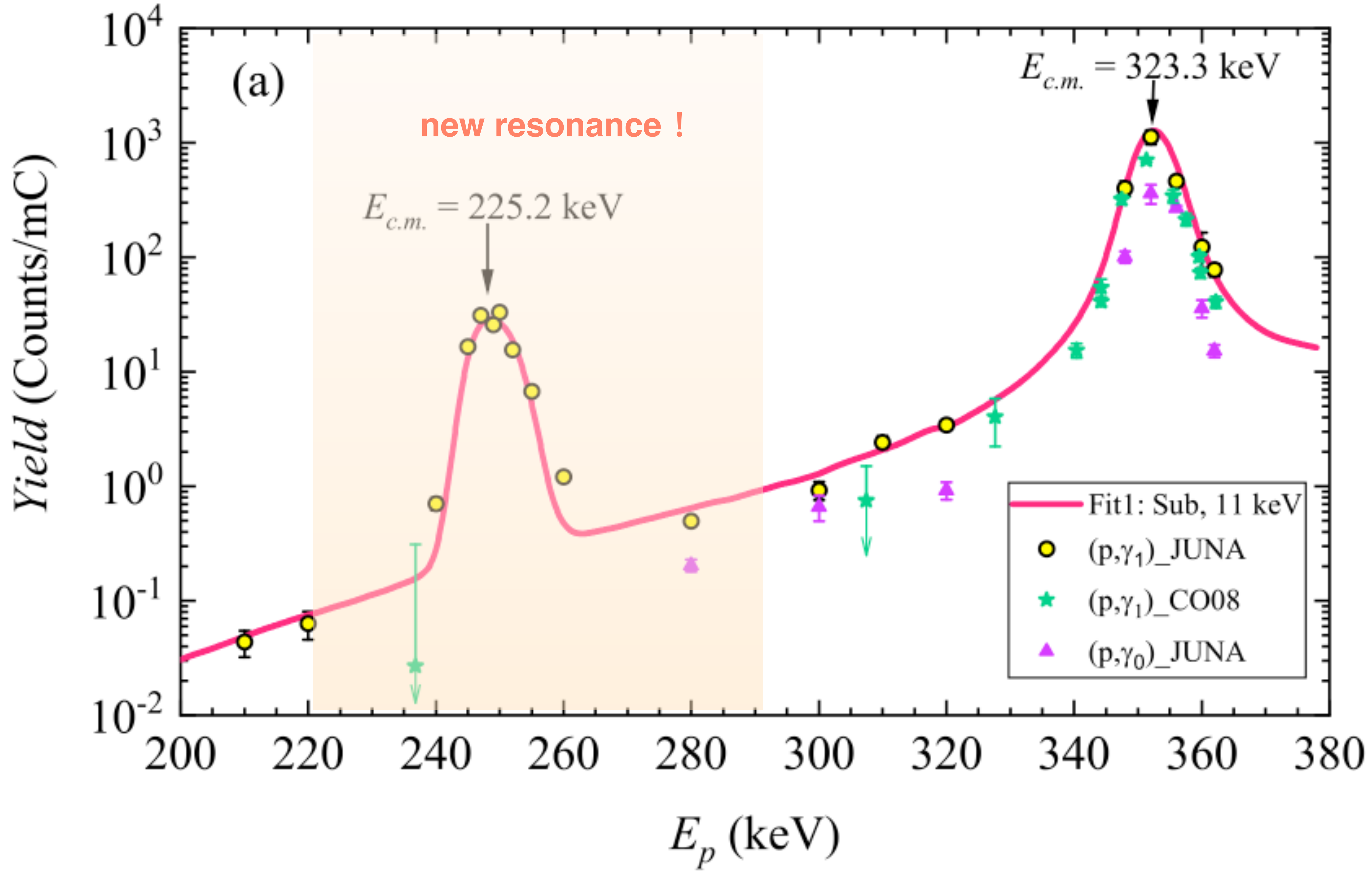
(200C)



BGO array with high

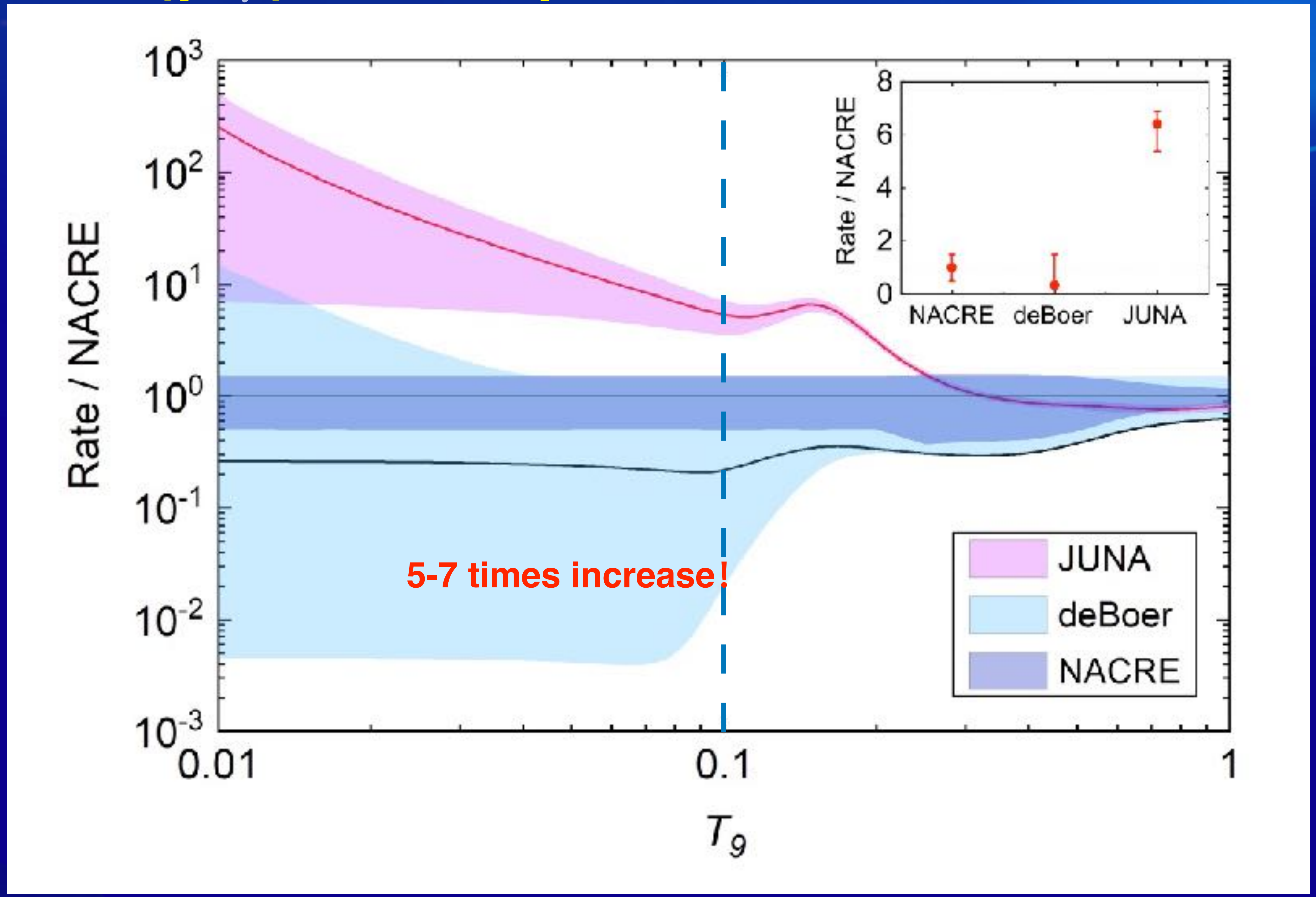
efficiency (60%) and

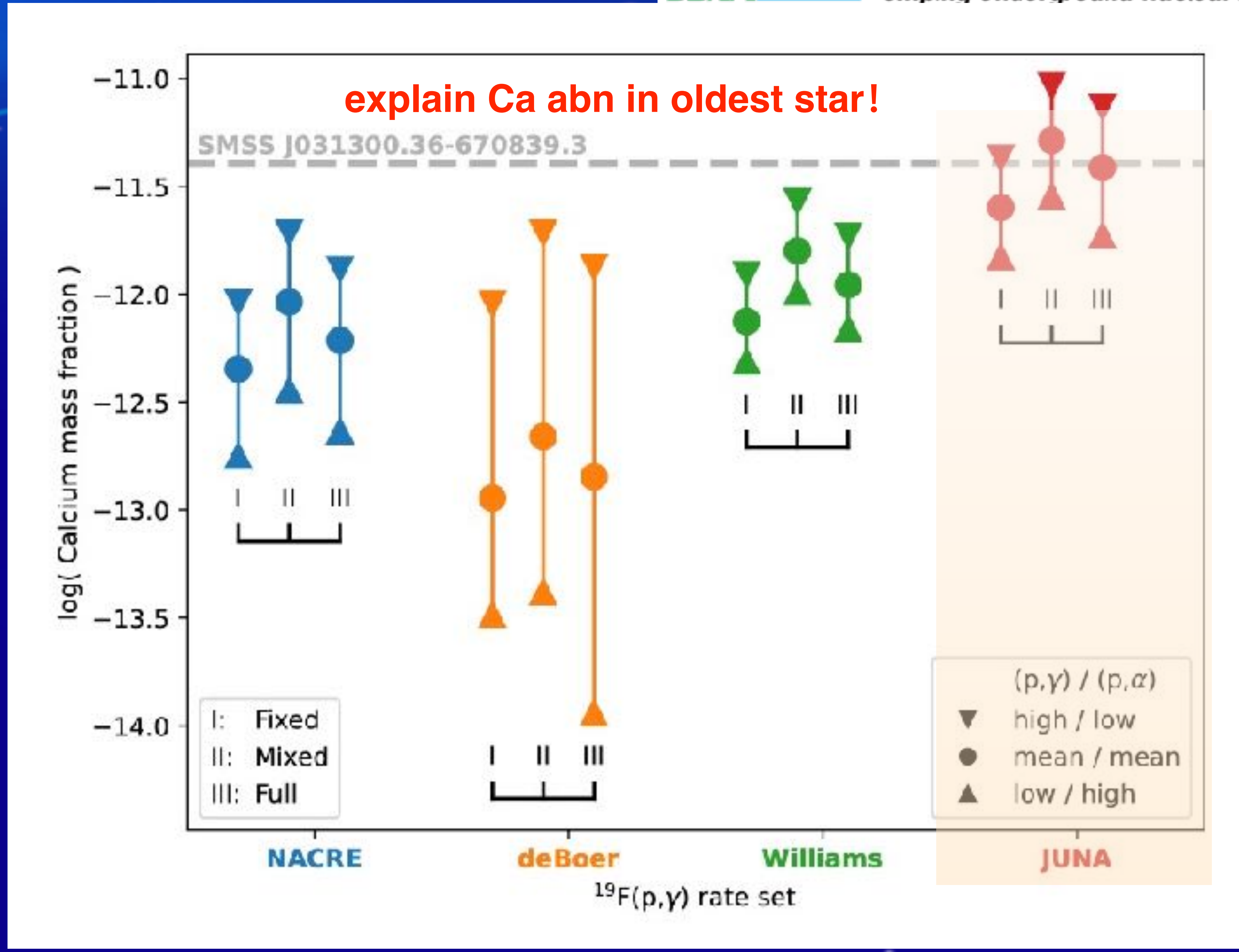
resolution (11%)



$^{19}F(p,\gamma)^{20}Ne$ implications

锦屏深地核天体物理实验
Jinping Underground Nuclear Astrophysics Experiment





Article

Measurement of $^{19}F(p, \gamma)^{20}Ne$ reaction suggests CNO breakout in first stars

https://doi.org/10.1038/s41586-022-05230-x
Received: 28 February 2022
Accepted: 11 August 2022
Check for updates

Liyong Zhang[1], Jianjun He[1✉], Richard J. deBoer[2], Michael Wiescher[2✉], Alexander Heger[3], Daid Kahl[4], Jun Su[1], Daniel Odell[5], Yinji Chen[1], Xinyue Li[1], Jianguo Wang[6], Long Zhang[7], Fuqiang Cao[7], Hao Zhang[1], Zhicheng Zhang[8], Xinzhi Jiang[1], Luohuan Wang[1], Ziming Li[1], Luyang Song[1], Hongwei Zhao[6], Liangting Sun[6], Qi Wu[6], Jiaqing Li[6], Baoqun Cui[7], Lihua Chen[7], Ruigang Ma[7], Ertao Li[8], Gang Lian[7], Yaode Sheng[1], Zhihong Li[7], Bing Guo[7], Xiaohong Zhou[6], Yuhu Zhang[6], Hushan Xu[6], Jianping Cheng[1] & Weiping Liu[7✉]

Nuclear astrophysics

Underground route to grasping the oldest stars

Marco Pignatari & Athanasios Psaltis

Nuclear-fusion experiments performed deep under Earth's surface reveal one possible scenario that could have resulted in the chemical abundances found in an ancient star in the Milky Way. **See p.656**

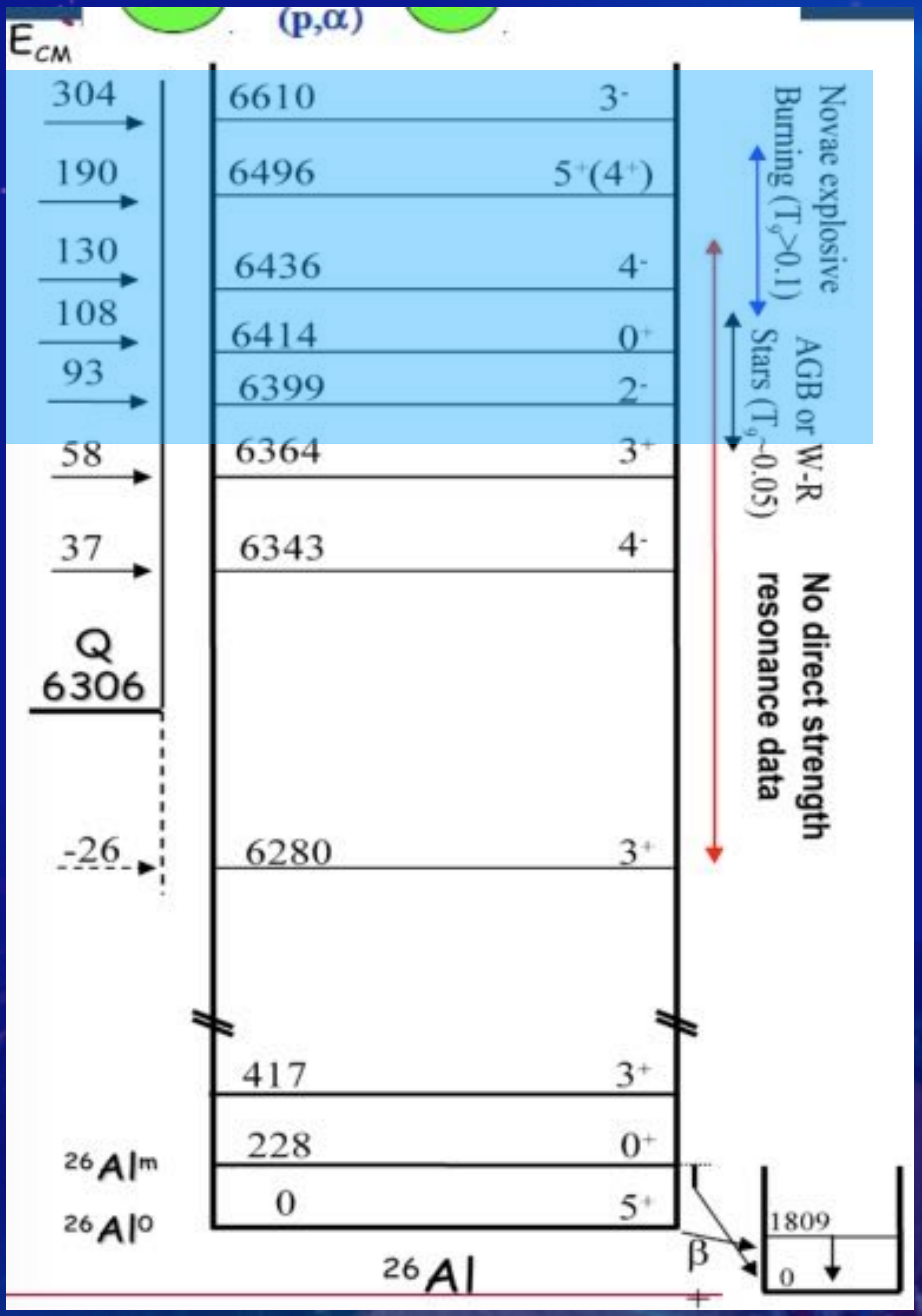
L. Y. Zhang, J. J. He*, ..., WPL*, Nature 610(2022)656，Selected as news and views

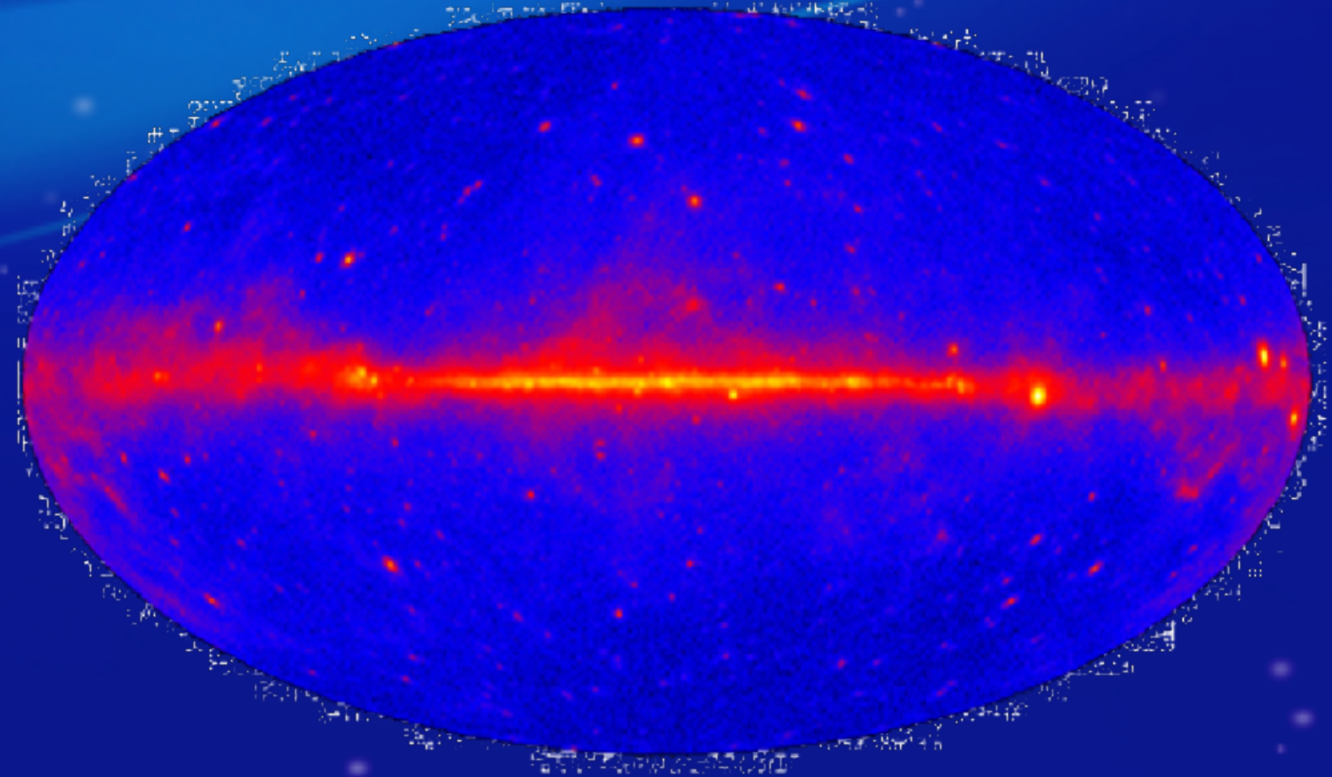
^{25}Mg(p,γ)^{26}Al：gamma astronomy

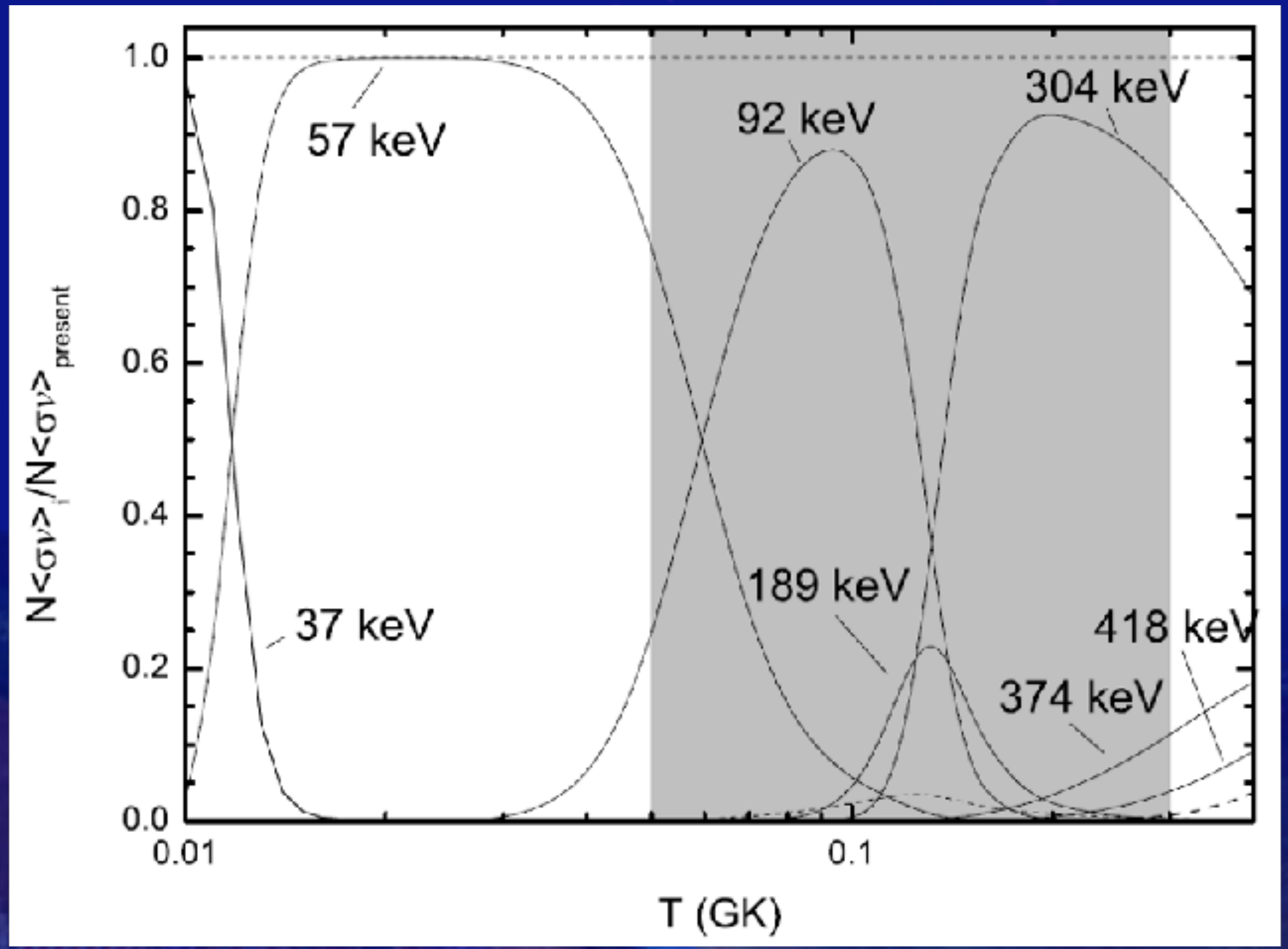
PI: Z. H. Li, CIAE

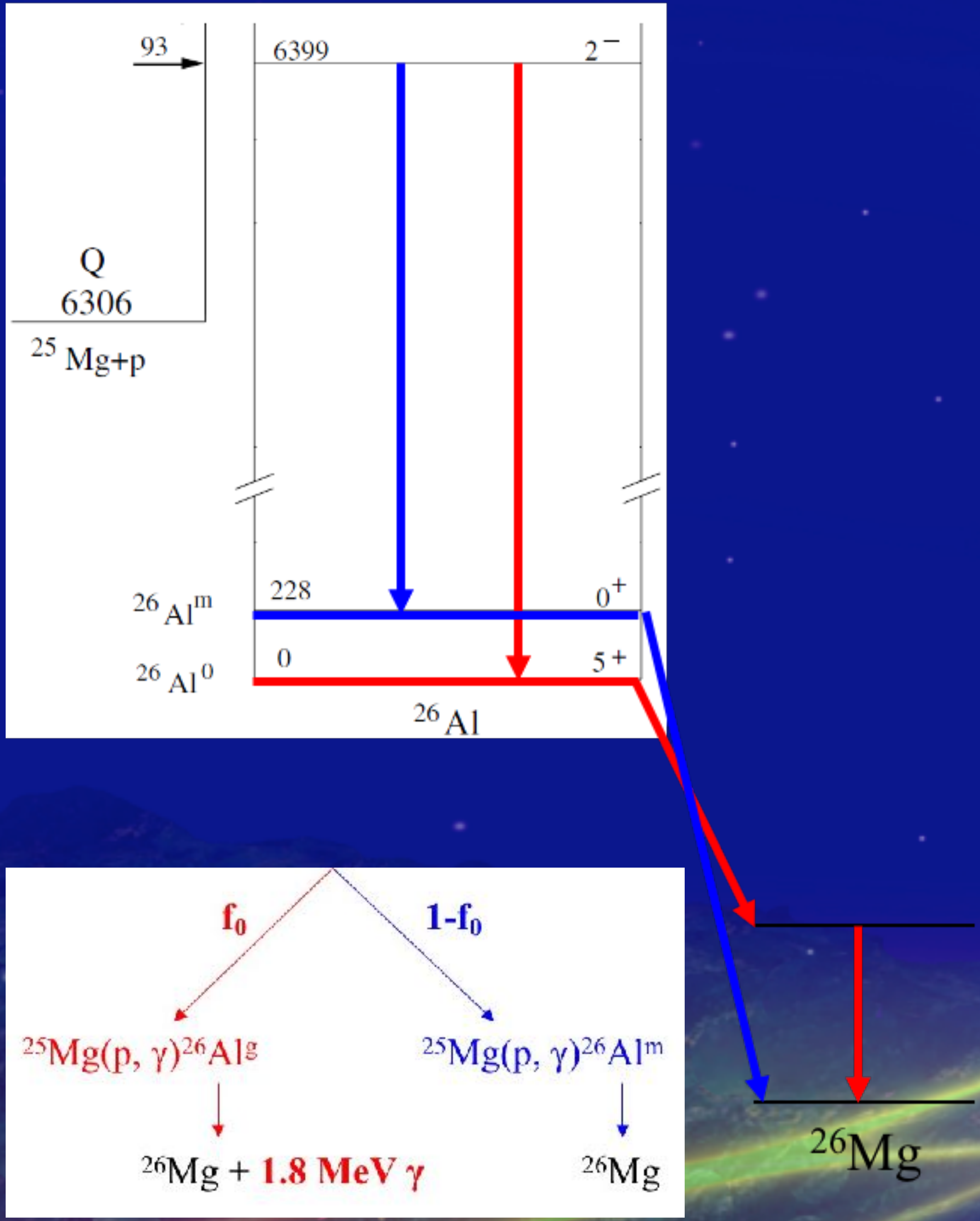
J. Su, CIAE/BNU

Exp.: Jan. 1-15, 2021



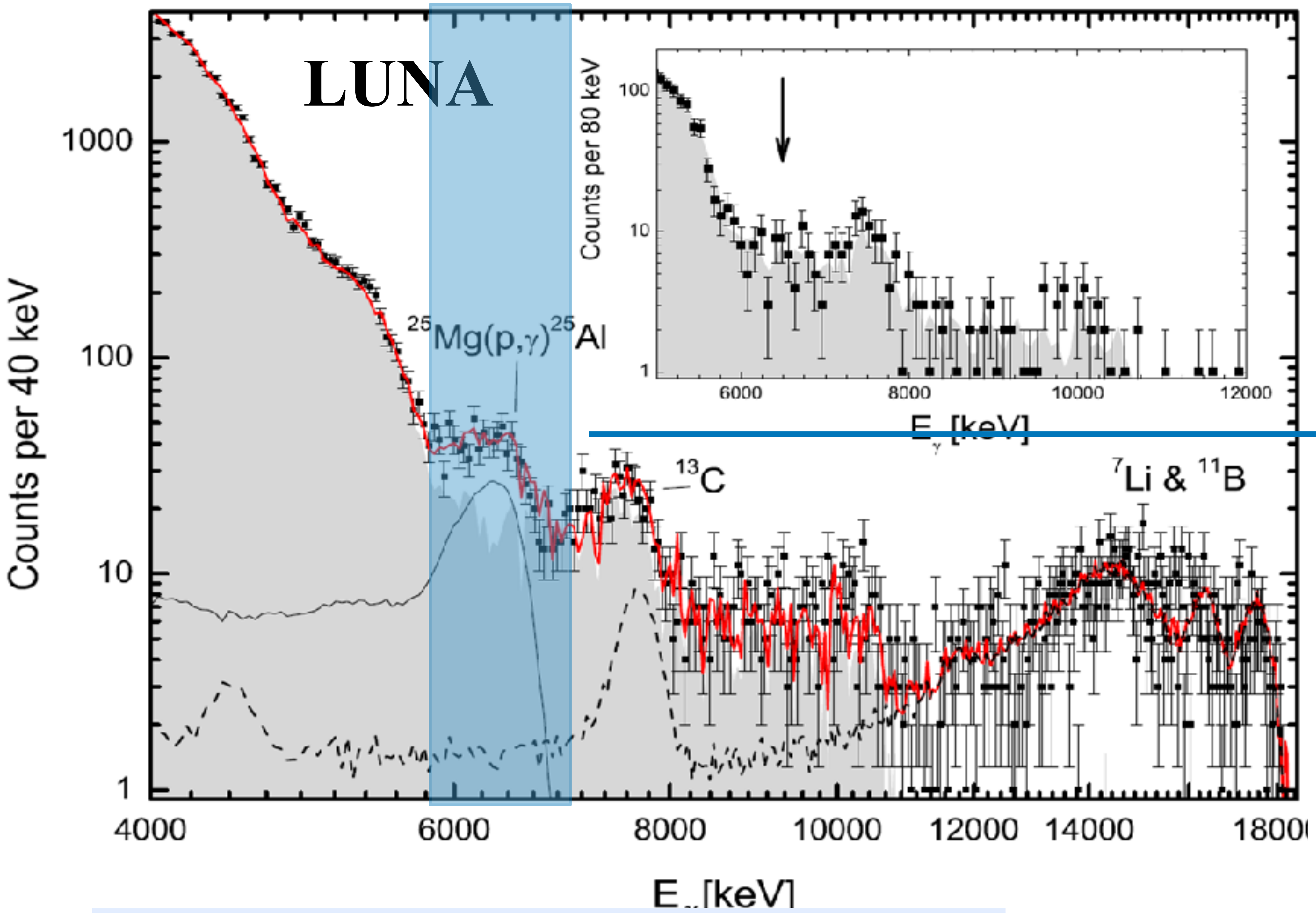


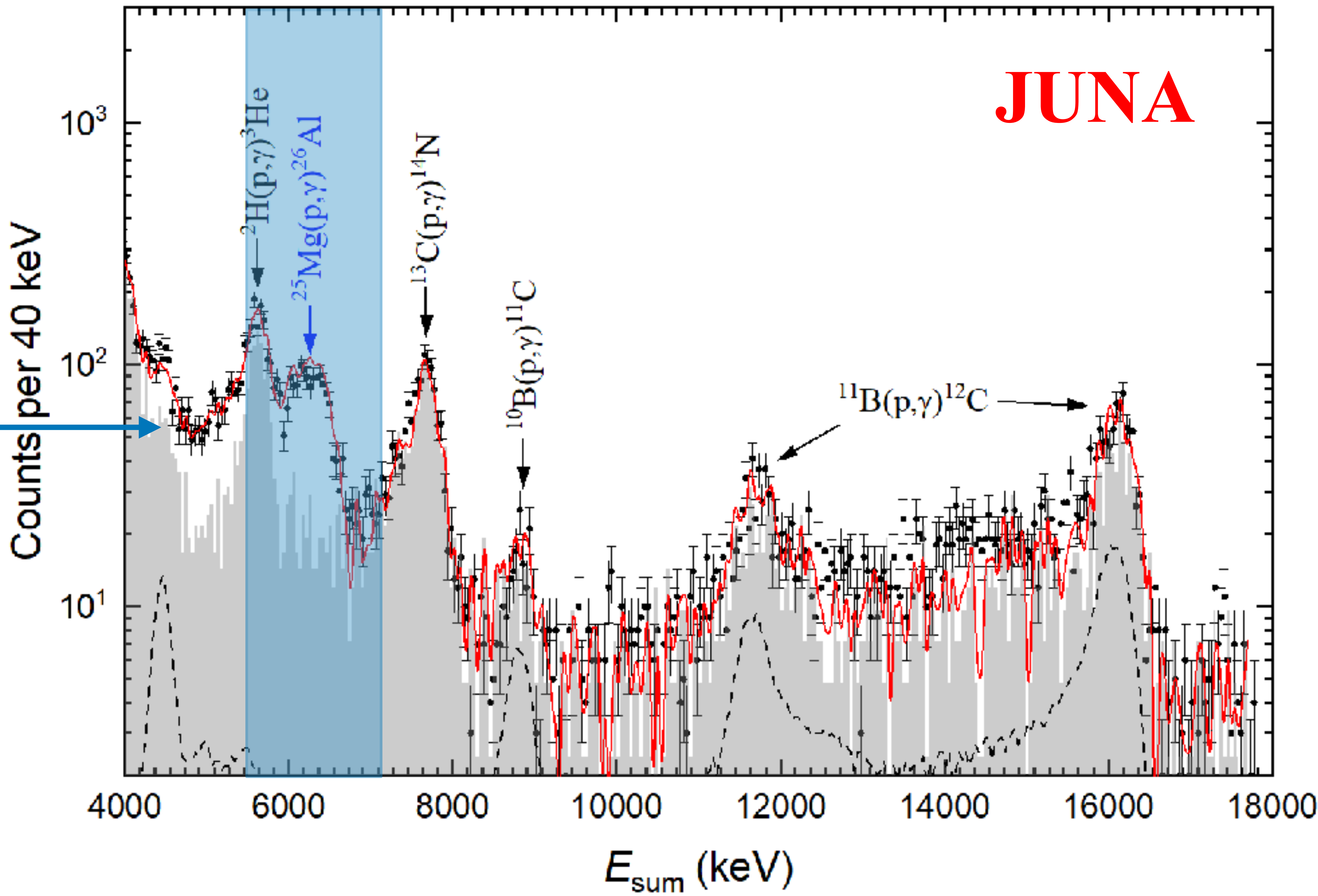




Z. H. Li et al., Sci. China Phys 58. 082002(2015), J. Su et al., Science Bulletin, 67(2022)2（封面文章）

JUNA vs. LUNA





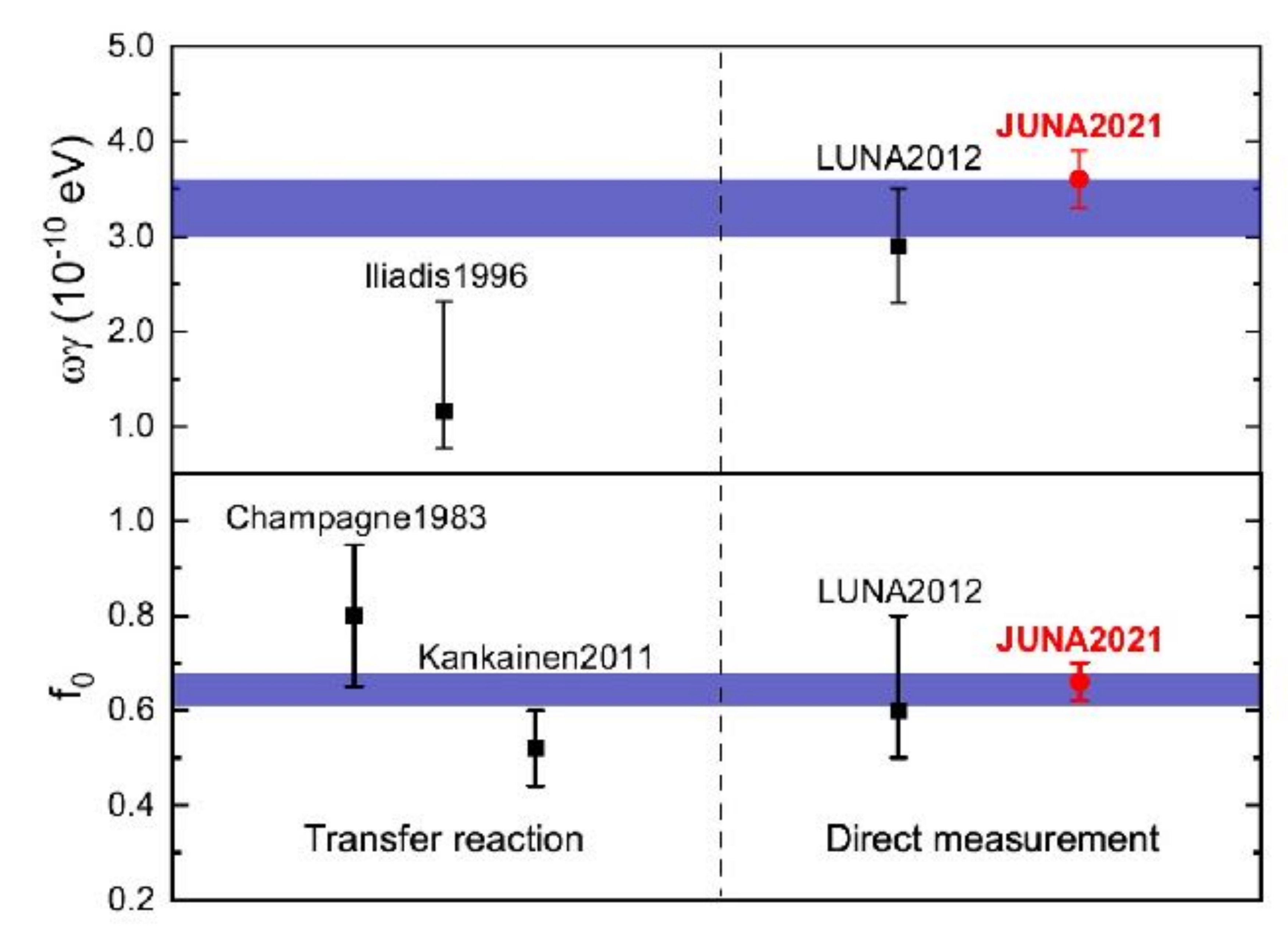
52 days， 370 C
signal：410
strength：2.9±0.6×10^{-10} eV

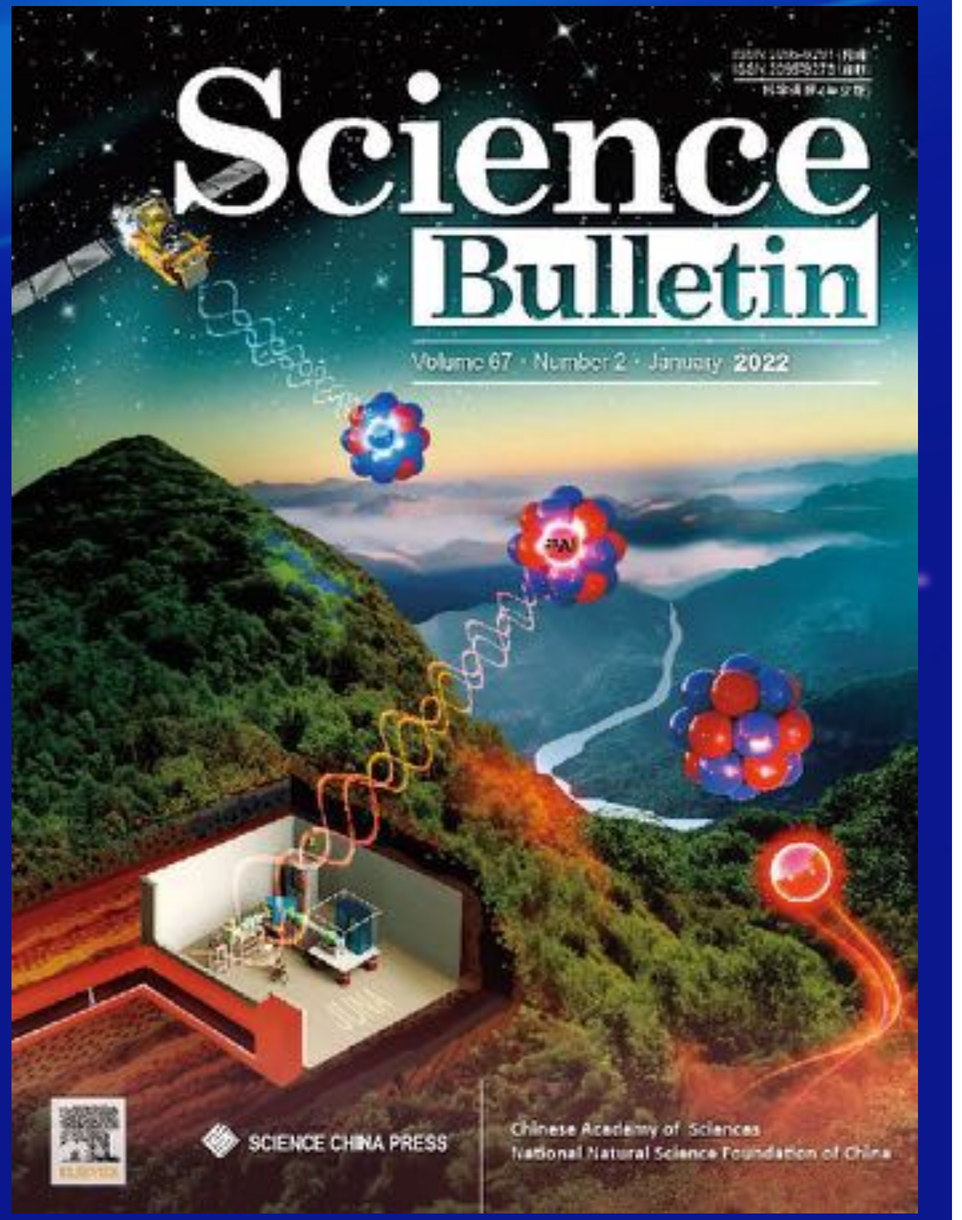
→

15days，1008 C
signal：1225
strength：3.8±0.4×10^{-10} eV

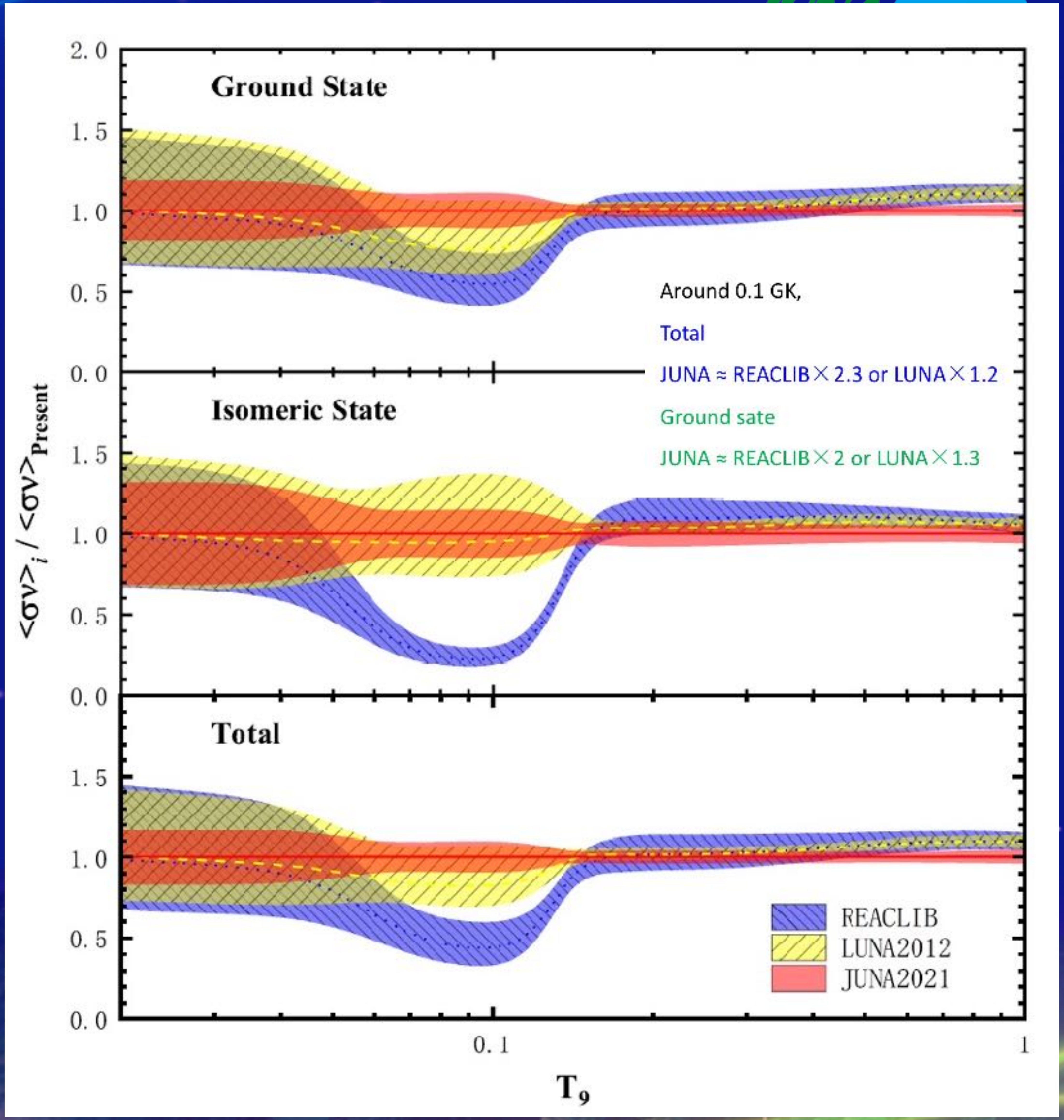
J. Su, Z. H. Li*,...., WPL*, Science Bulletin, 67(2022)2, cover paper

JUNA result of $^{25}Mg(p,\gamma)^{26}Al$





IF>20



BRIF in-direct

Y. J. Li, Z. H. Li, E. T. Li, X. Y. Li, T. L. Ma, Y. P. Shen, J. C. Liu, L. Gan, Y. Su, L.-H. Qiao, *et al.*, Phys. Rev. C **102**, 025804 (2020).

E_x (keV)[a]	$\omega\gamma$ (eV)	f_0
37.1 ± 0.1	$(4.5 \pm 1.8) \times 10^{-22}$[b]	0.79 ± 0.05[b]
57.7 ± 0.1	$(2.9 \pm 0.5) \times 10^{-13}$[c]	0.81 ± 0.05[b]
92.1 ± 0.2	$(3.8 \pm 0.3) \times 10^{-10}$[d]	0.66 ± 0.04[d]
189.6 ± 0.1	$(9.0 \pm 0.6) \times 10^{-7}$[b]	0.75 ± 0.02[b]
304.1 ± 0.1	$(3.1 \pm 0.1) \times 10^{-2}$[e]	0.859 ± 0.01[e]

JUNA underground

JUNA ground

J. Su, Z. H. Li*,…., WPL*, Science Bulletin, 67(2022)2,cover paper

$^{13}C(a,n)^{16}O$ neutron source reaction for heavy elements


PI: X. D. Tang, IMP

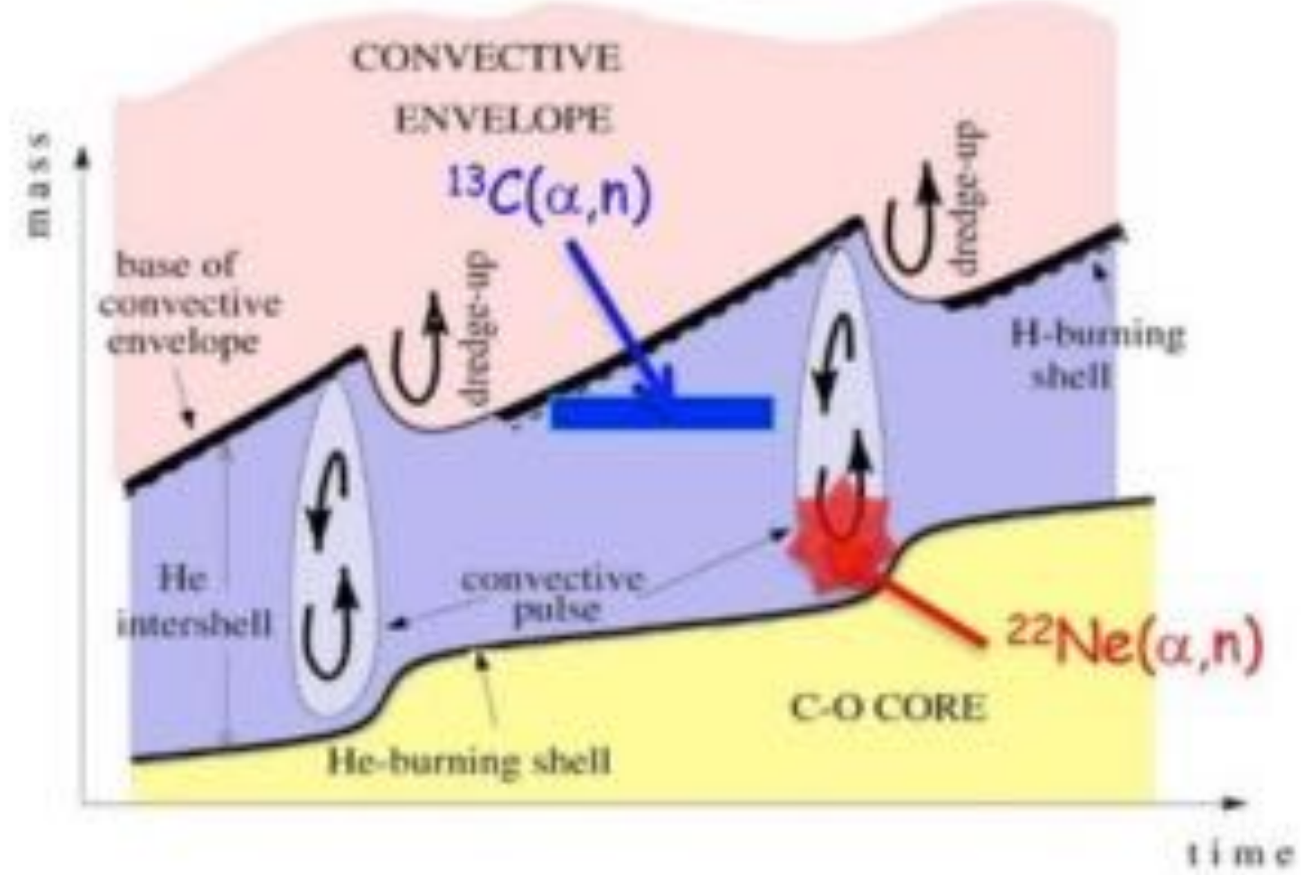
B. Gao, IMP

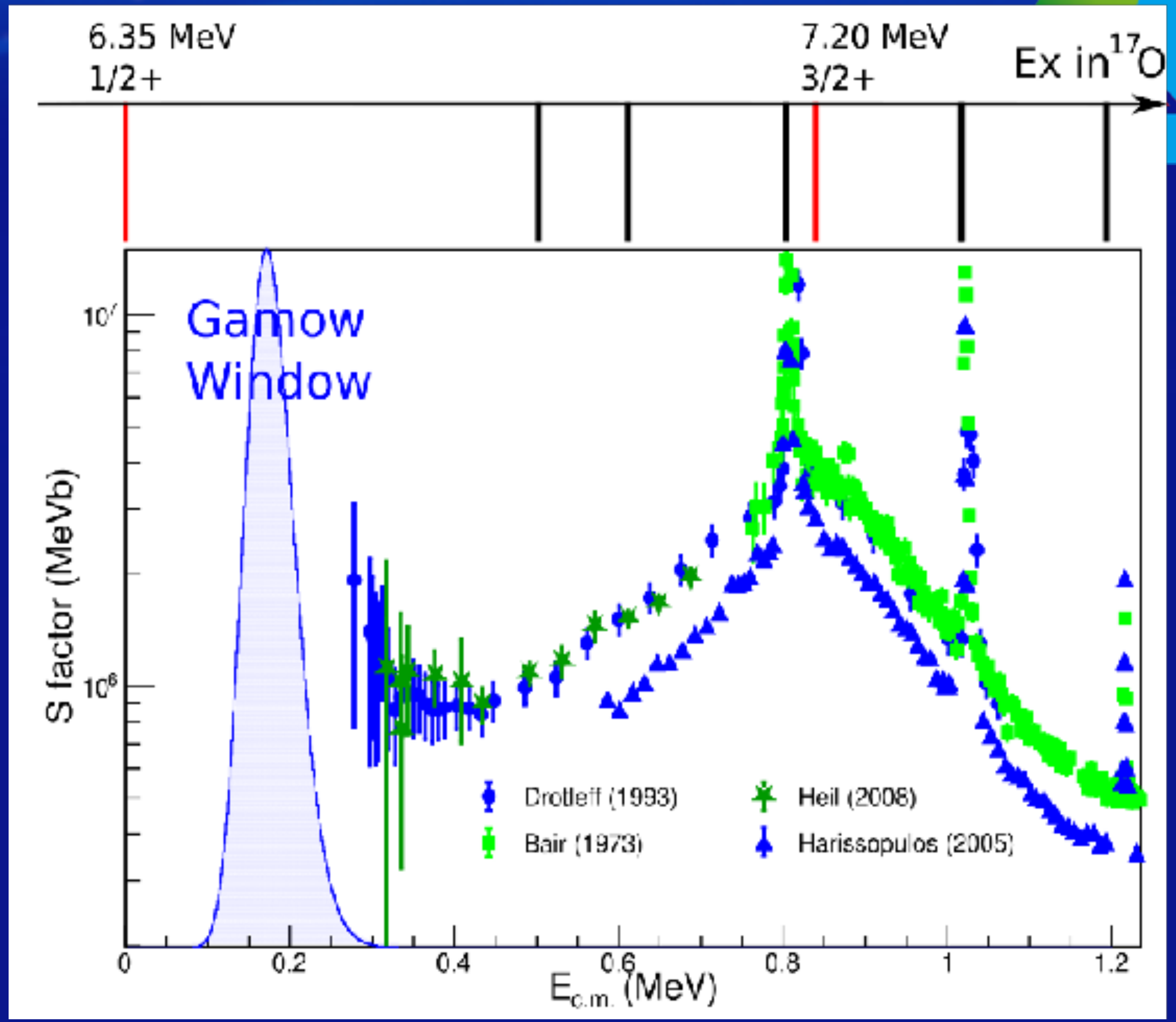
Exp.: Jan. 27-Feb. 16, 2021

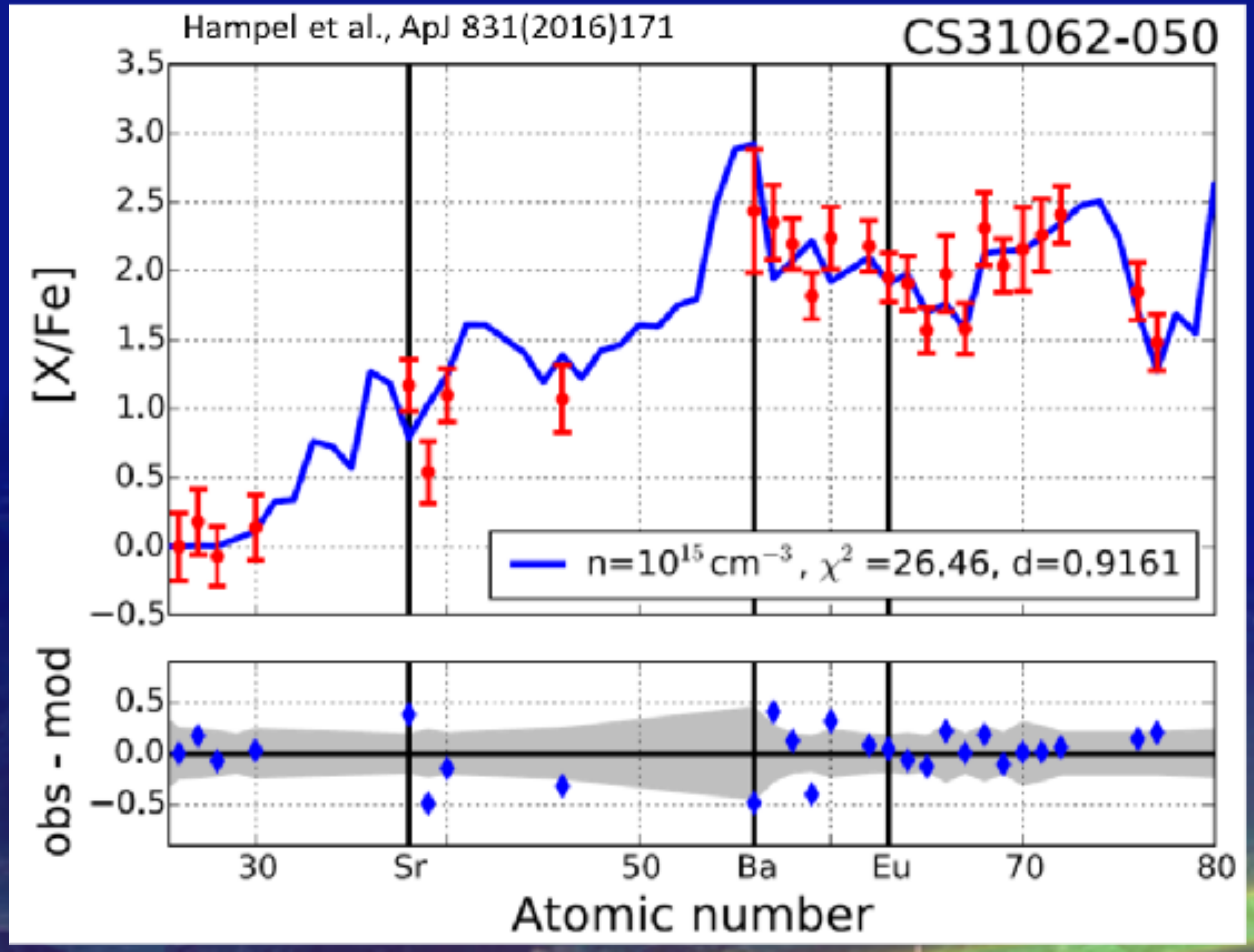
Main s-process ~90<A<210

TP-AGB stars

shell H-burning	He-flash
$T_9 \sim 0.1$ K	$0.25 \leq T_9 \sim 0.4$ K
10^7-10^8 cm^{-3}	10^{10}-10^{11} cm^{-3}
$^{13}C(\alpha,n)^{16}O$	$^{22}Ne(\alpha,n)^{25}Mg$



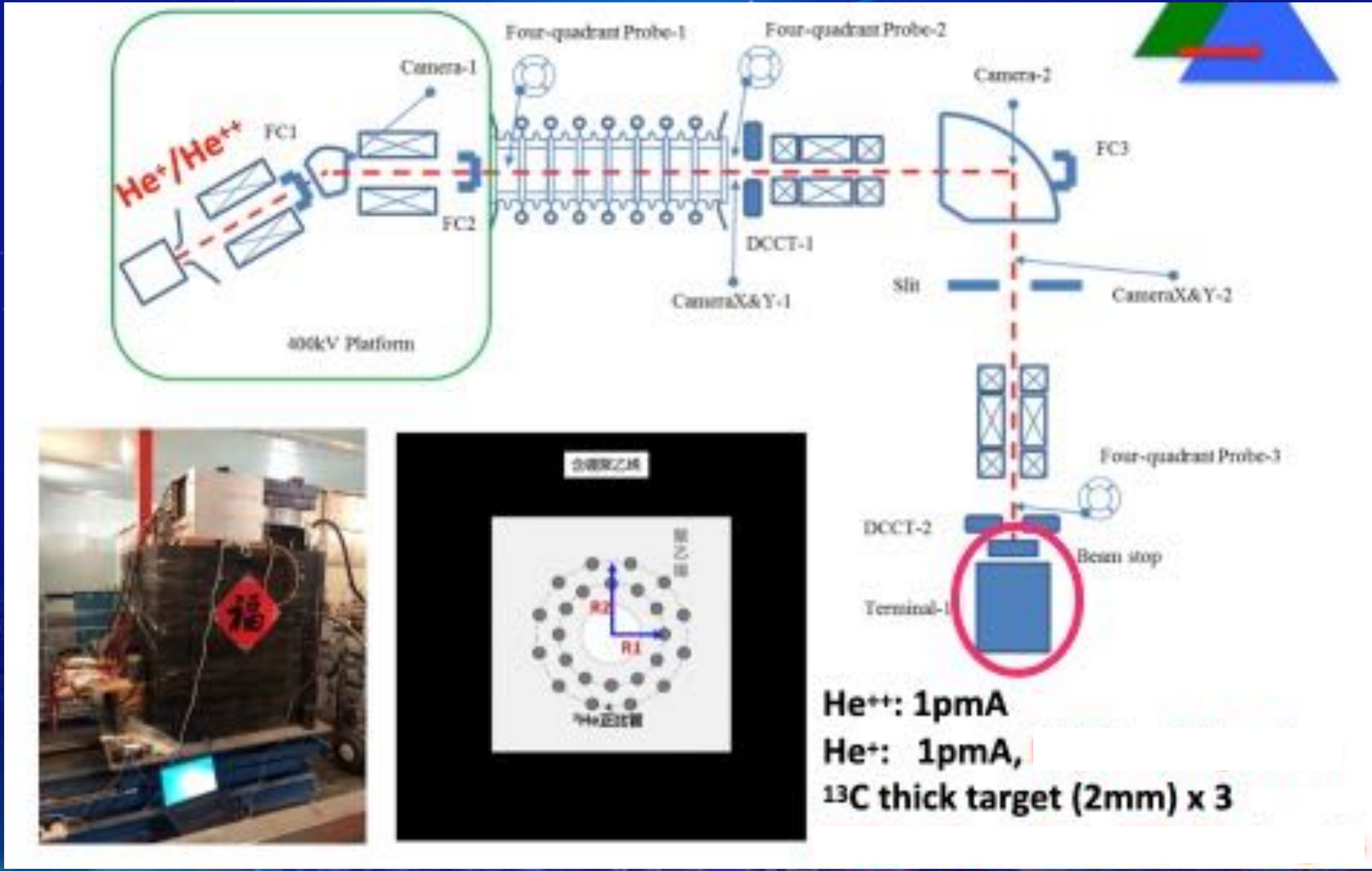


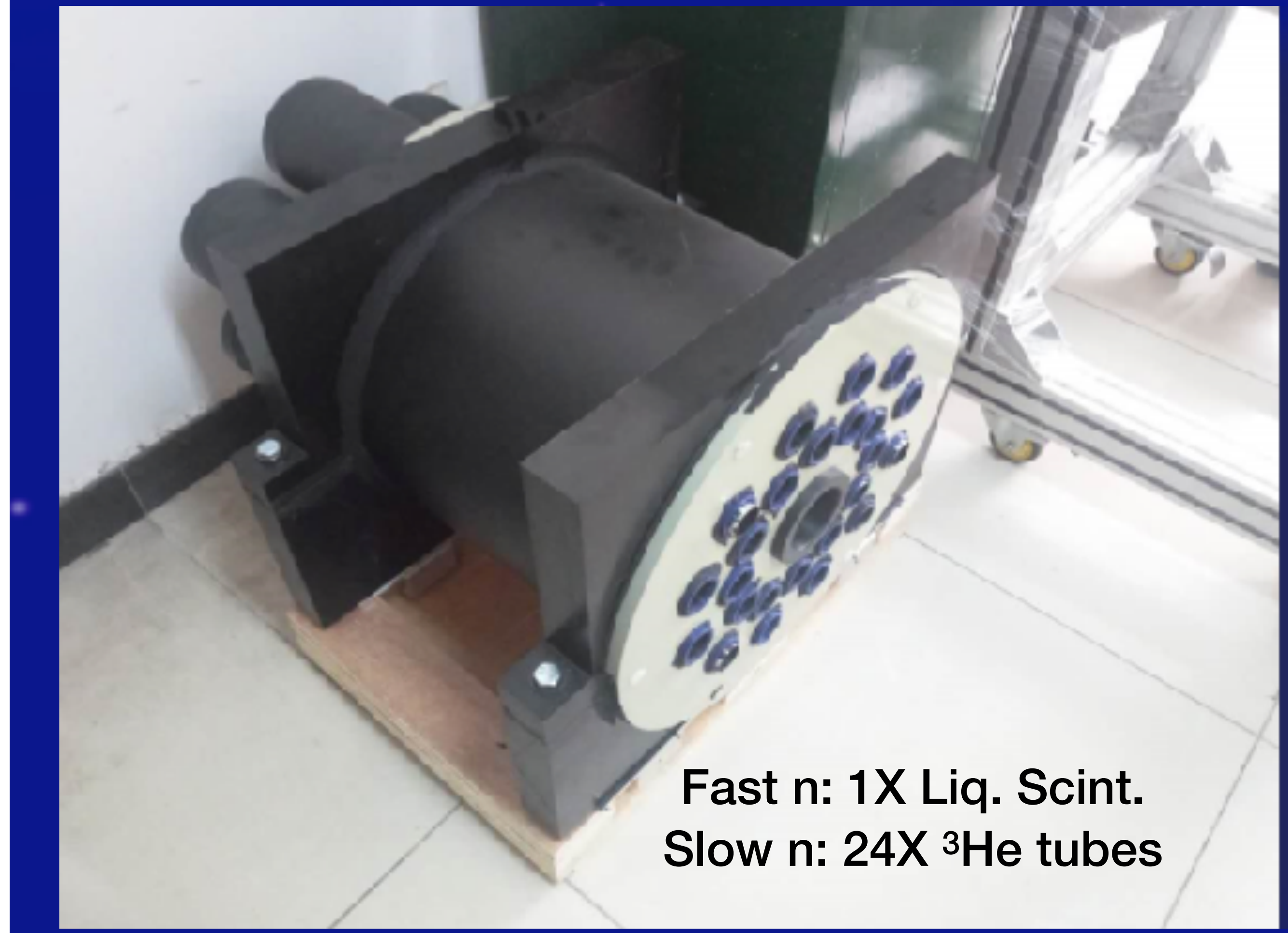


B. Gao, …, Y. D. Tang*,…, WPL*, $^{13}C(\alpha,n)^{16}O$, PRL 129(2022)132701

Lin, Nov. 3, 11:10-11:35
this workshop

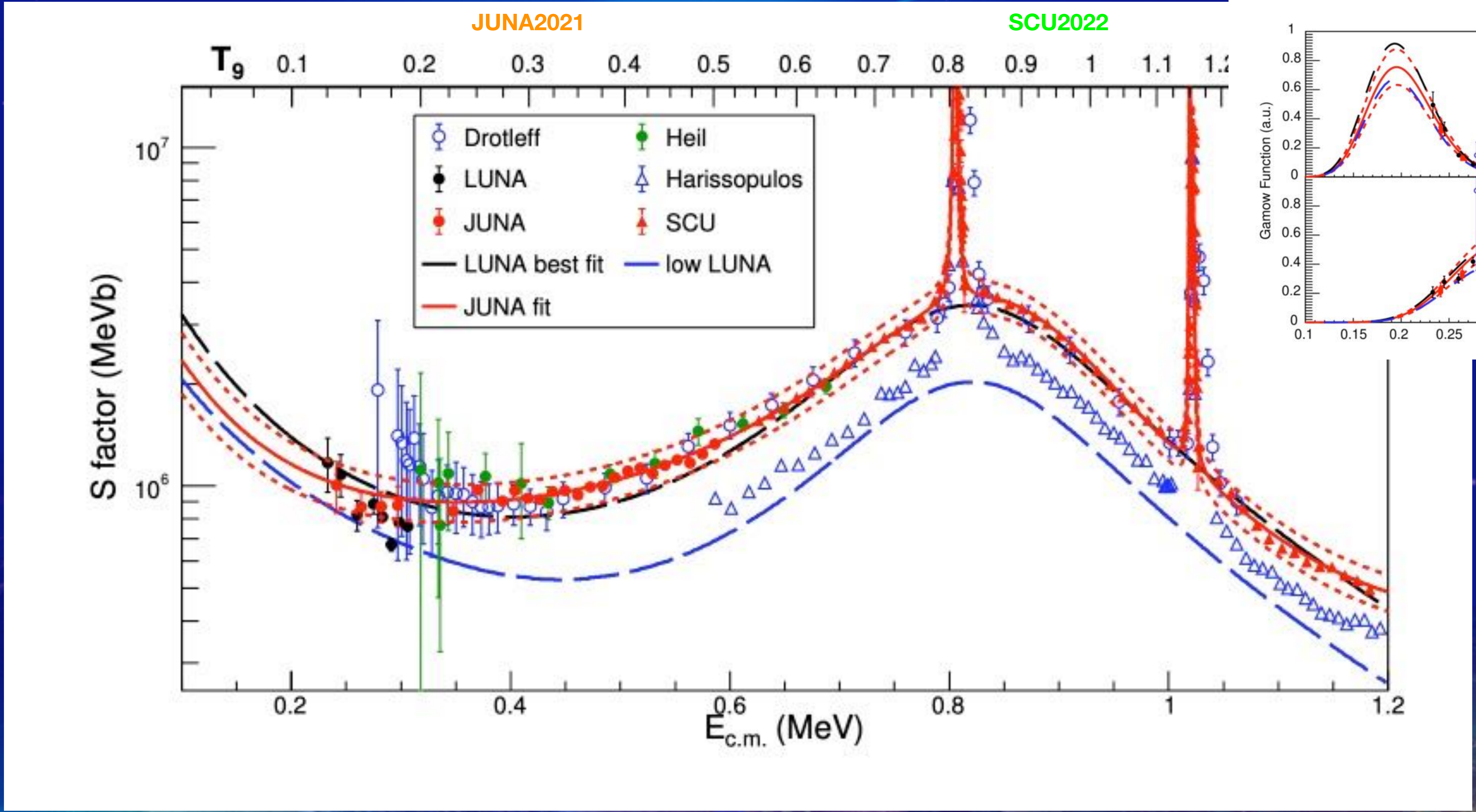
$^{13}C(\alpha,n)^{16}O$ neutron detection

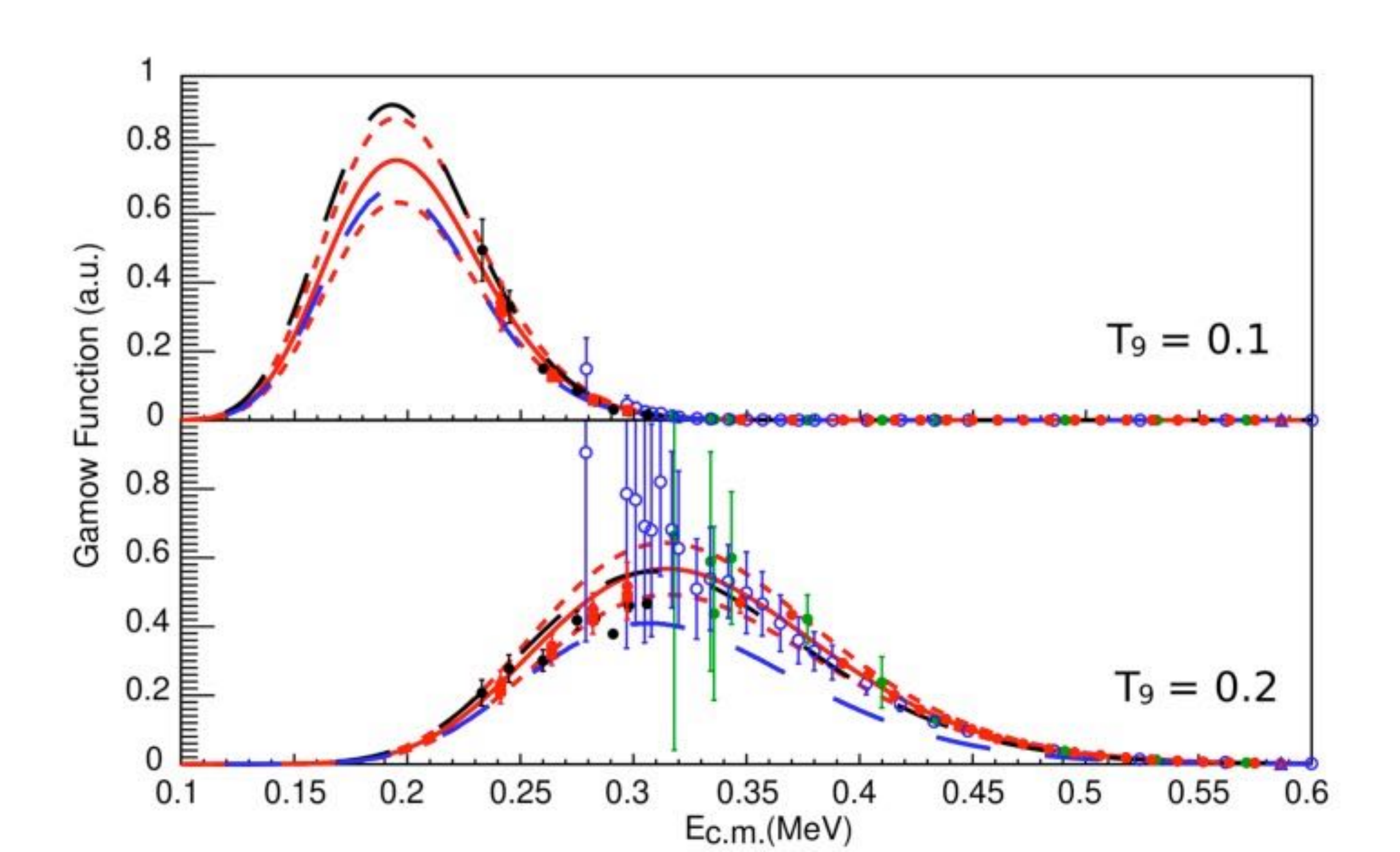




B. Gao, ..., Y. D. Tang*,..., WPL*, $^{13}C(\alpha,n)^{16}O$, PRL 129(2022)132701

$^{13}C(\alpha,n)^{16}O$ results





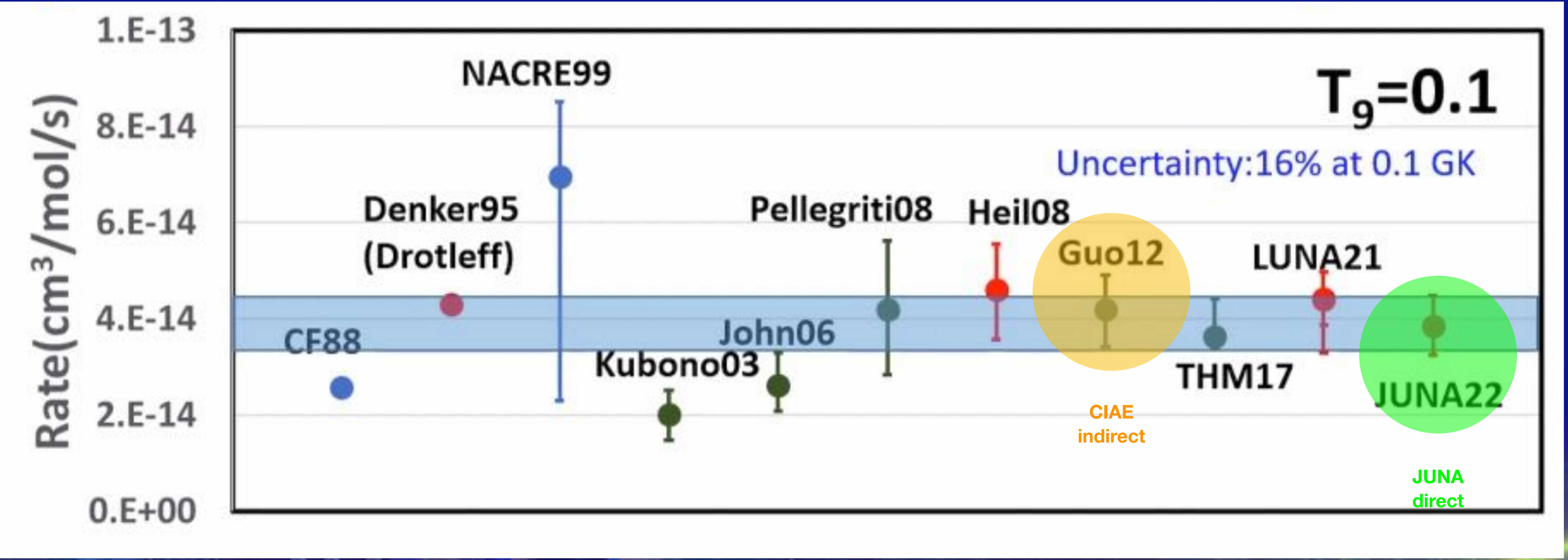
- mA thick target, differential method to pin down thickness
- magnetic removal of He^{2+}, cover 0.4 MeV to 0.8 MeV (JUNA), cover i-process; to 1.2 MeV tandem, calibration of eff. , cross check other date
- n background 5/ hour, 2.5 MeV eff. 25%, good S/N

B. Gao, …, Y. D. Tang*,…, WPL*, $^{13}C(\alpha,n)^{16}O$, PRL 129(2022)132701

Resolve conflict in30 years research

Annu. Rev. Nucl. Part. Sci. 2023. 73:315–40

Recent studies of the low-energy range of the $^{13}C(\alpha,n)^{16}O$ reaction rate at the deep underground accelerator facilities of LUNA in Italy (83) and JUNA in China (84) have removed most of the uncertainties in the extrapolation of the previous higher-energy data [the NACRE II compilation (85)]. The low-energy data match well the prediction of a recent R-matrix analysis (86)



B. Gao, …, Y. D. Tang*,…, WPL*, $^{13}C(\alpha,n)^{16}O$, PRL 129(2022)132701

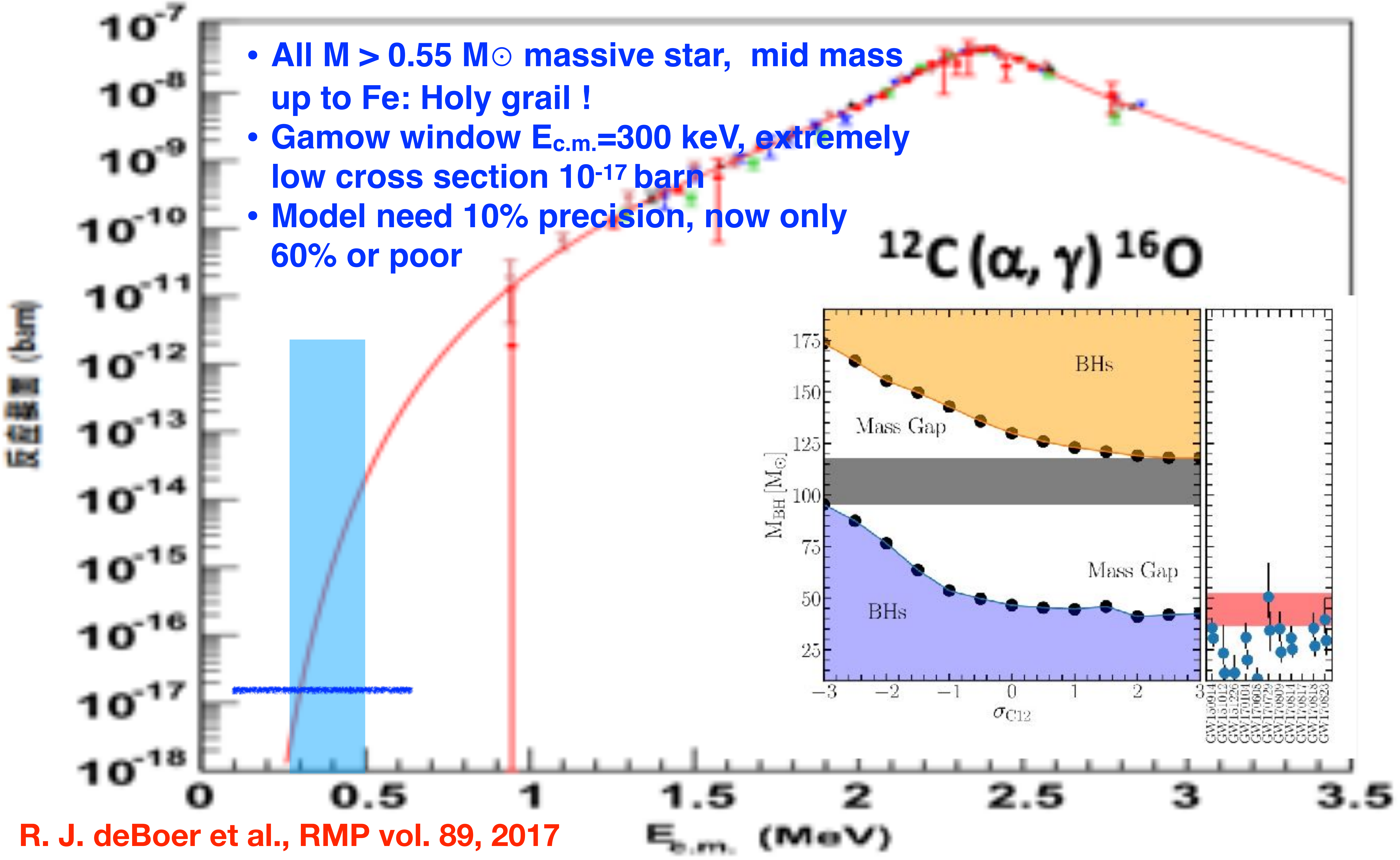
B. Guo*, Z. H. Li,…, WPL*, Astrophys. J. 756(2012)193.

Big question, big impact, big challenge

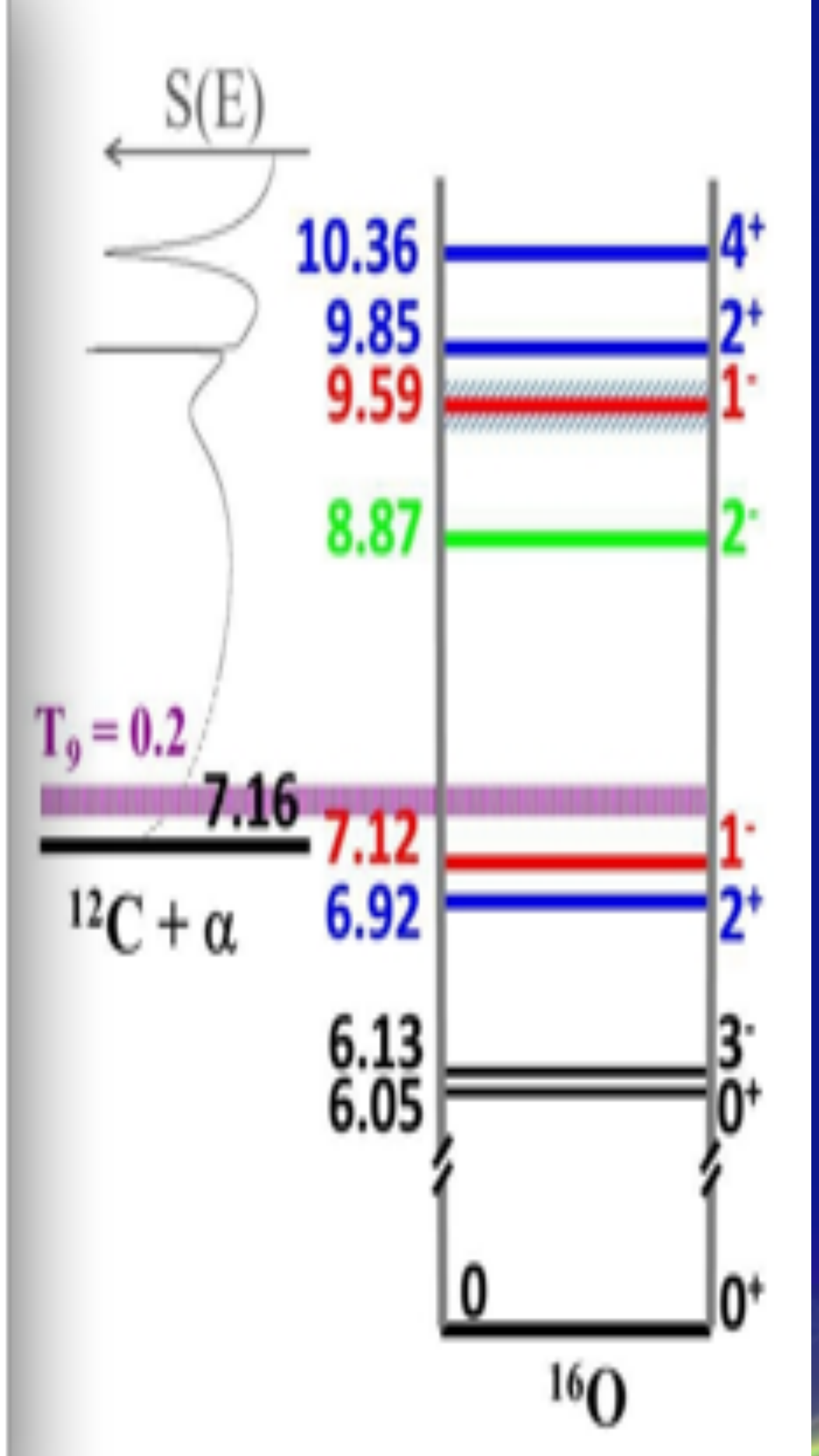
PI: WPL/Y. P. Shen, CIAE

Exp.: Feb. 26-Apr. 18, 2021

- All M > 0.55 M⊙ massive star, mid mass up to Fe: Holy grail !
- Gamow window $E_{c.m.}$=300 keV, extremely low cross section 10^{-17} barn
- Model need 10% precision, now only 60% or poor



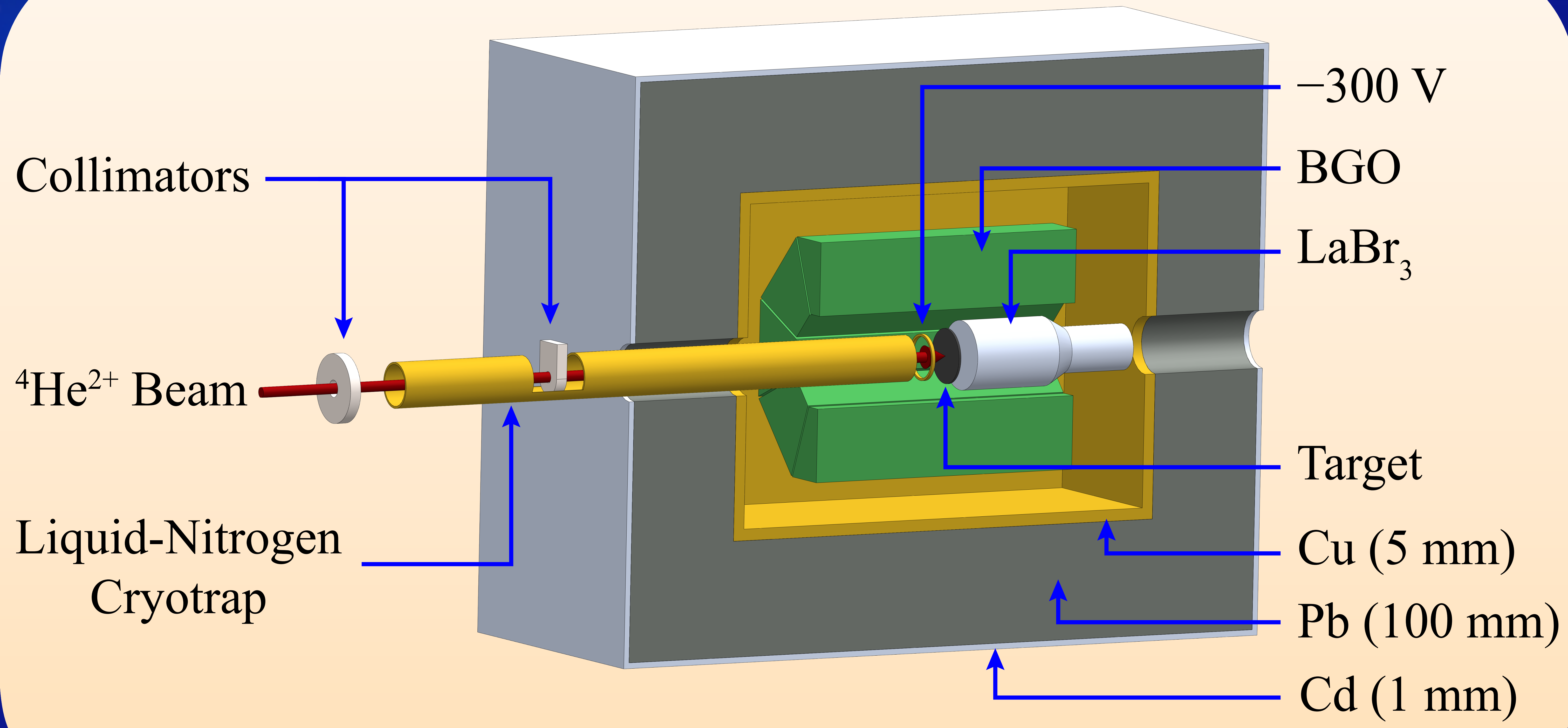
R. J. deBoer et al., RMP vol. 89, 2017



B. Guo, Z. H. Li, …, WPL,, APJ 756, 193 (2012); Y. P. Shen, B. Guo, …, WPL, PRL 124, 162701(2020)

R. Farmer1,2 , M. Renzo3 , S. E. de Mink et al., APJ 902, L36, 2020

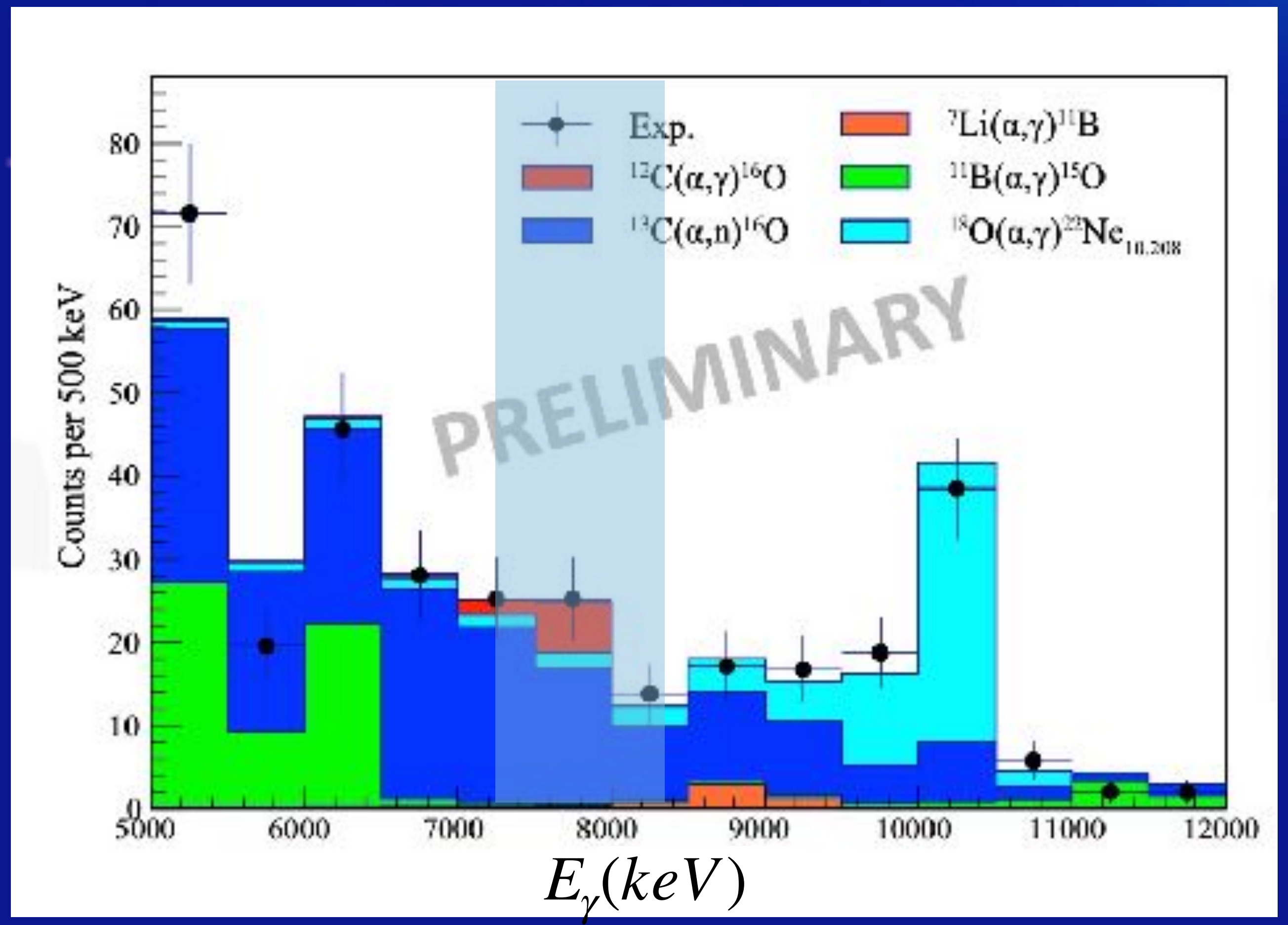
$^{12}C(\alpha,\gamma)^{16}O$ ： more sensitivity



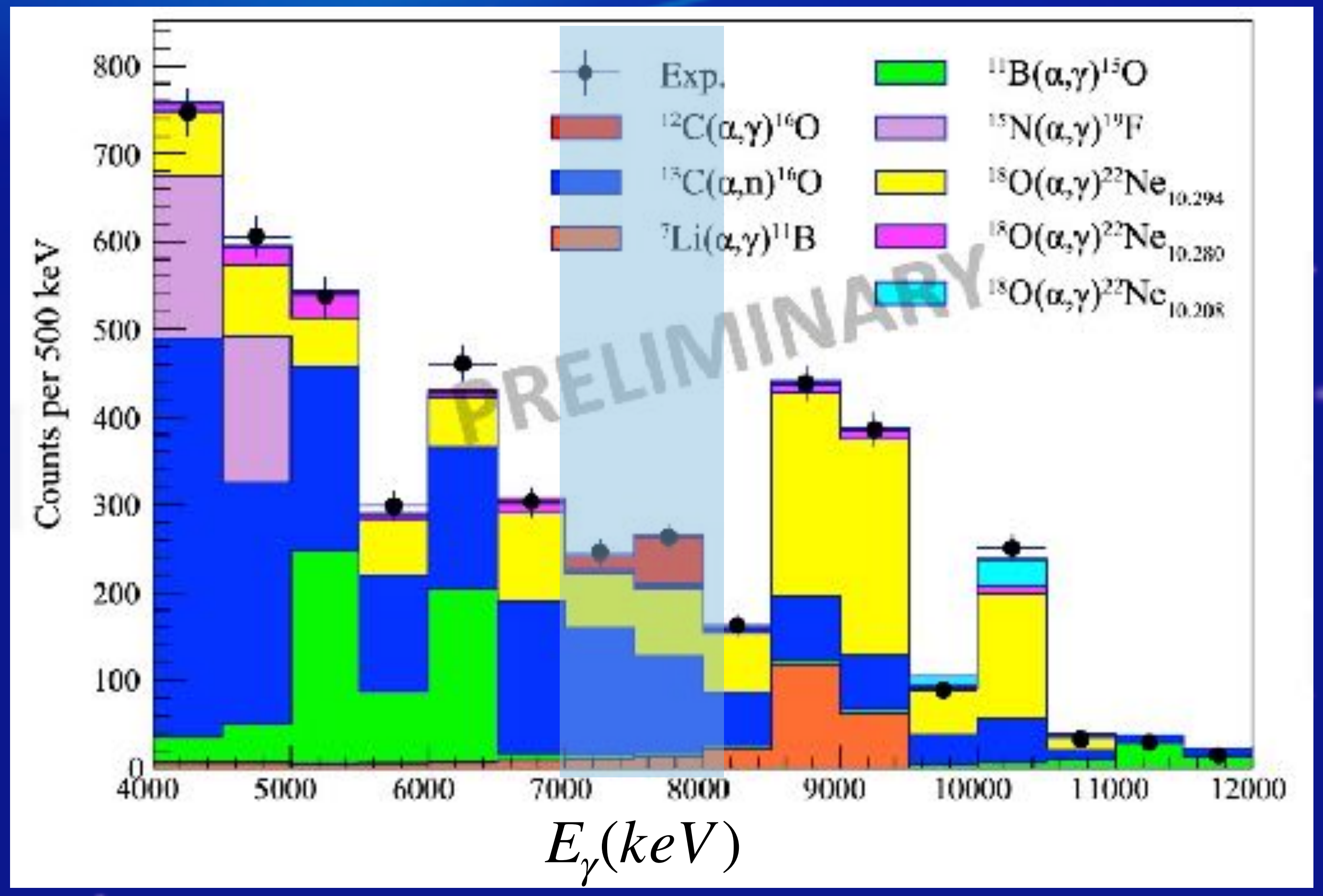
- FCVA implantation CTi thick targets
- durability >280 C @800 keV He^{2+}, with only 25% loss
- BGO+$LaBr_3$ (Lanthanum bromide) veto
- wide energy search for best S/N, 552 keV is best, other suffer from $^{18}O(\alpha,\gamma)^{22}Ne$ contaminations
- sensitivity of 10^{-12} b @$E_{c.m.}$ = 552 keV

$^{12}C(\alpha,\gamma)^{16}O$ ： submitted

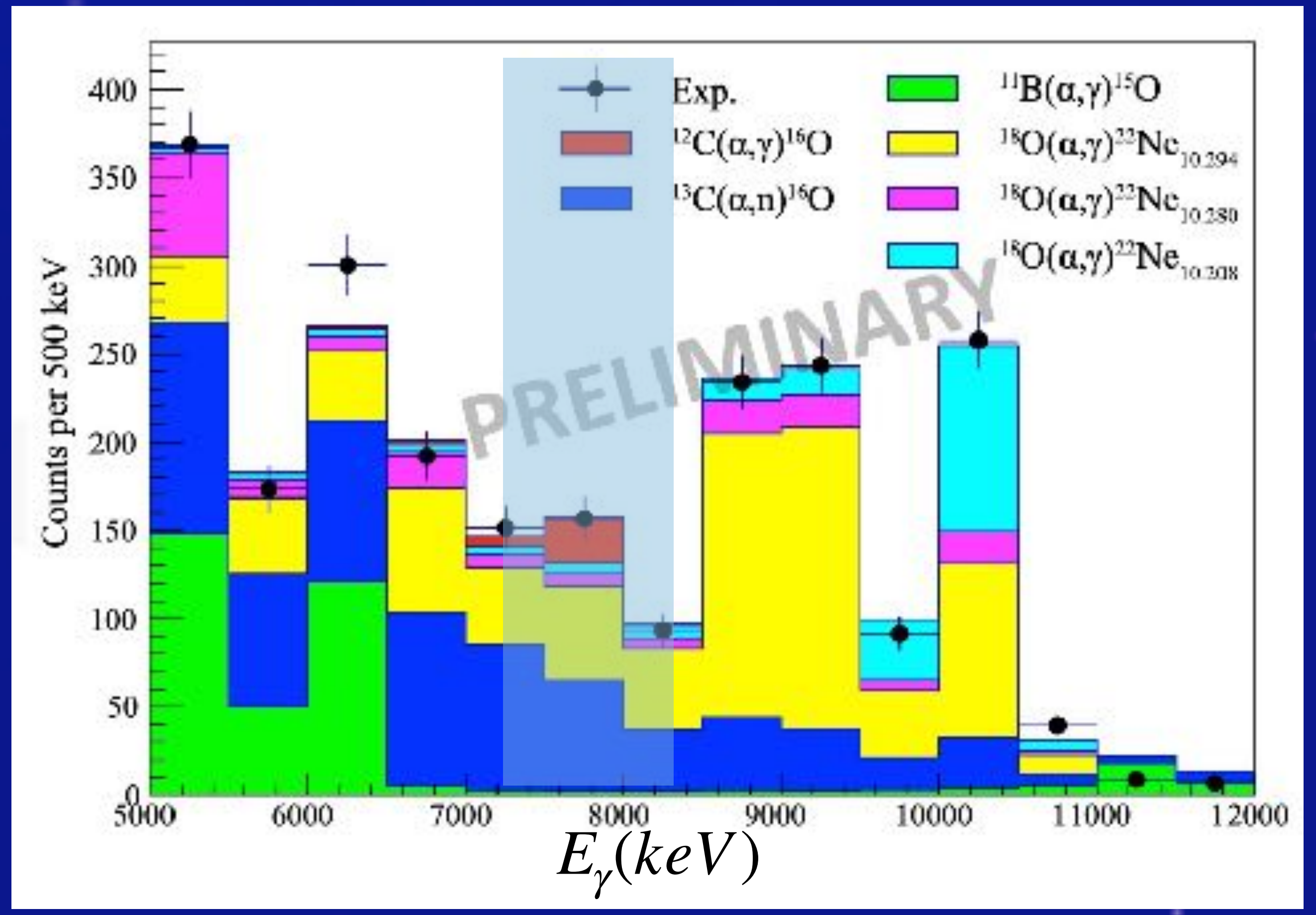
552 keV, 126 C

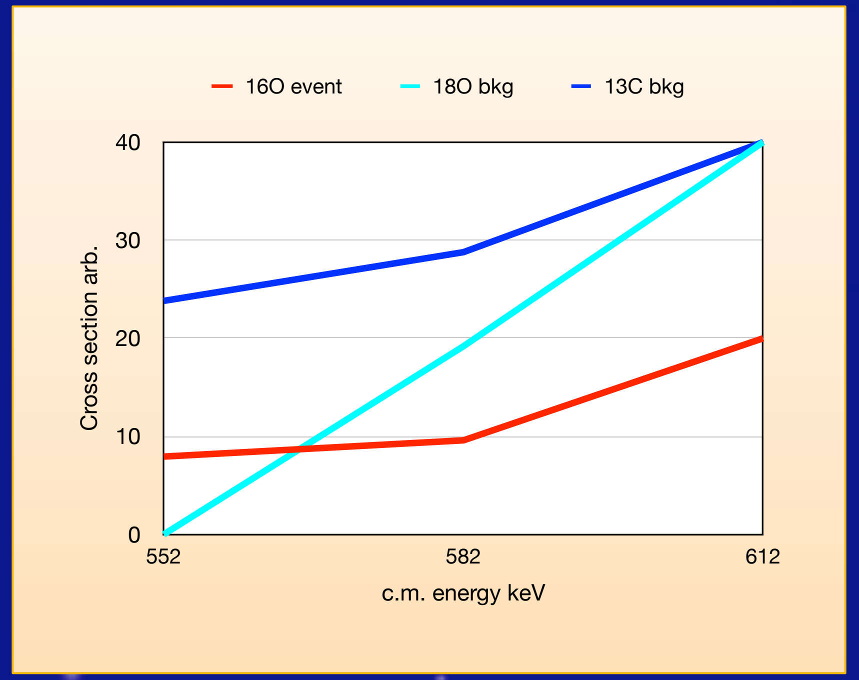


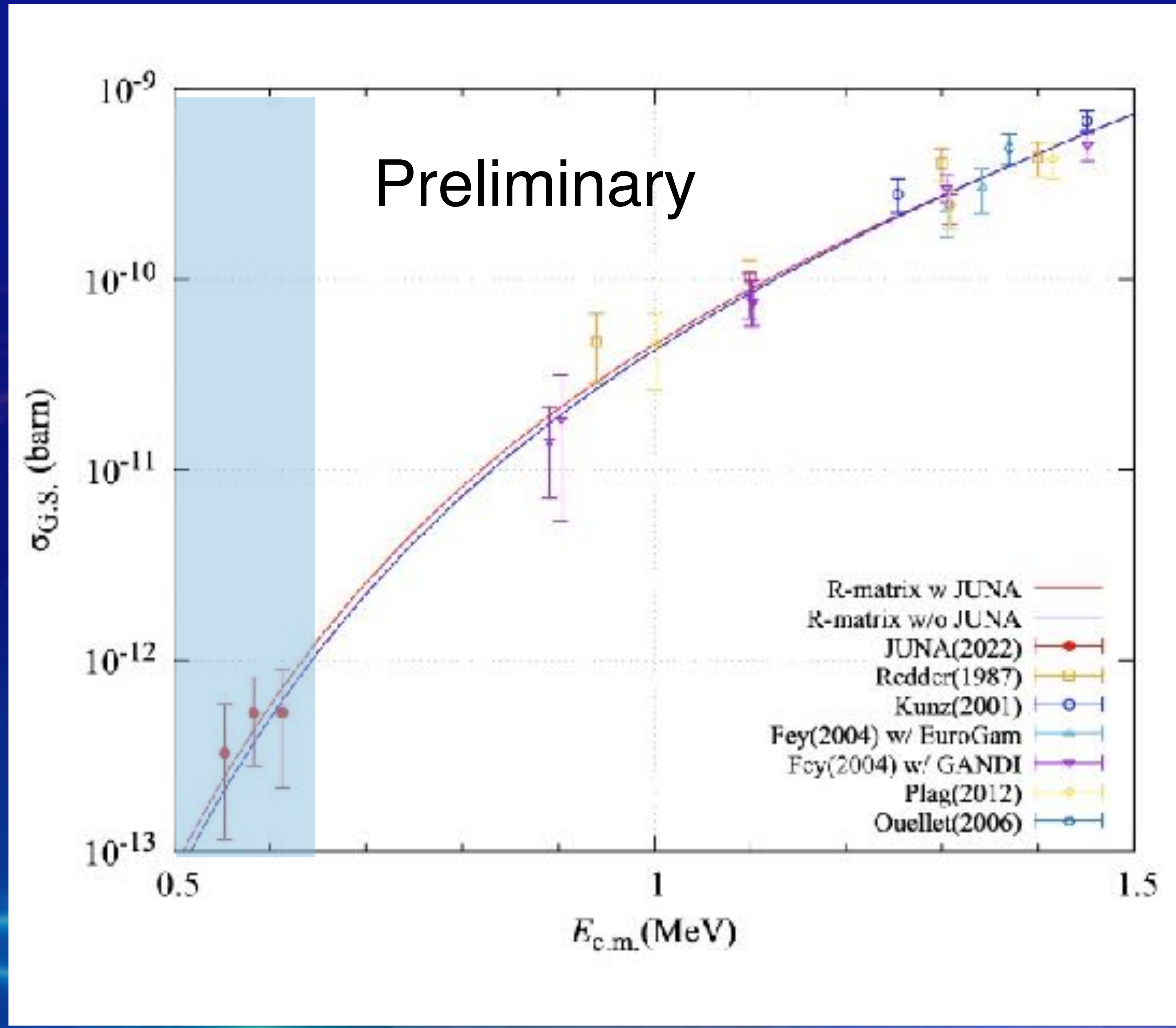
582 keV, 417 C

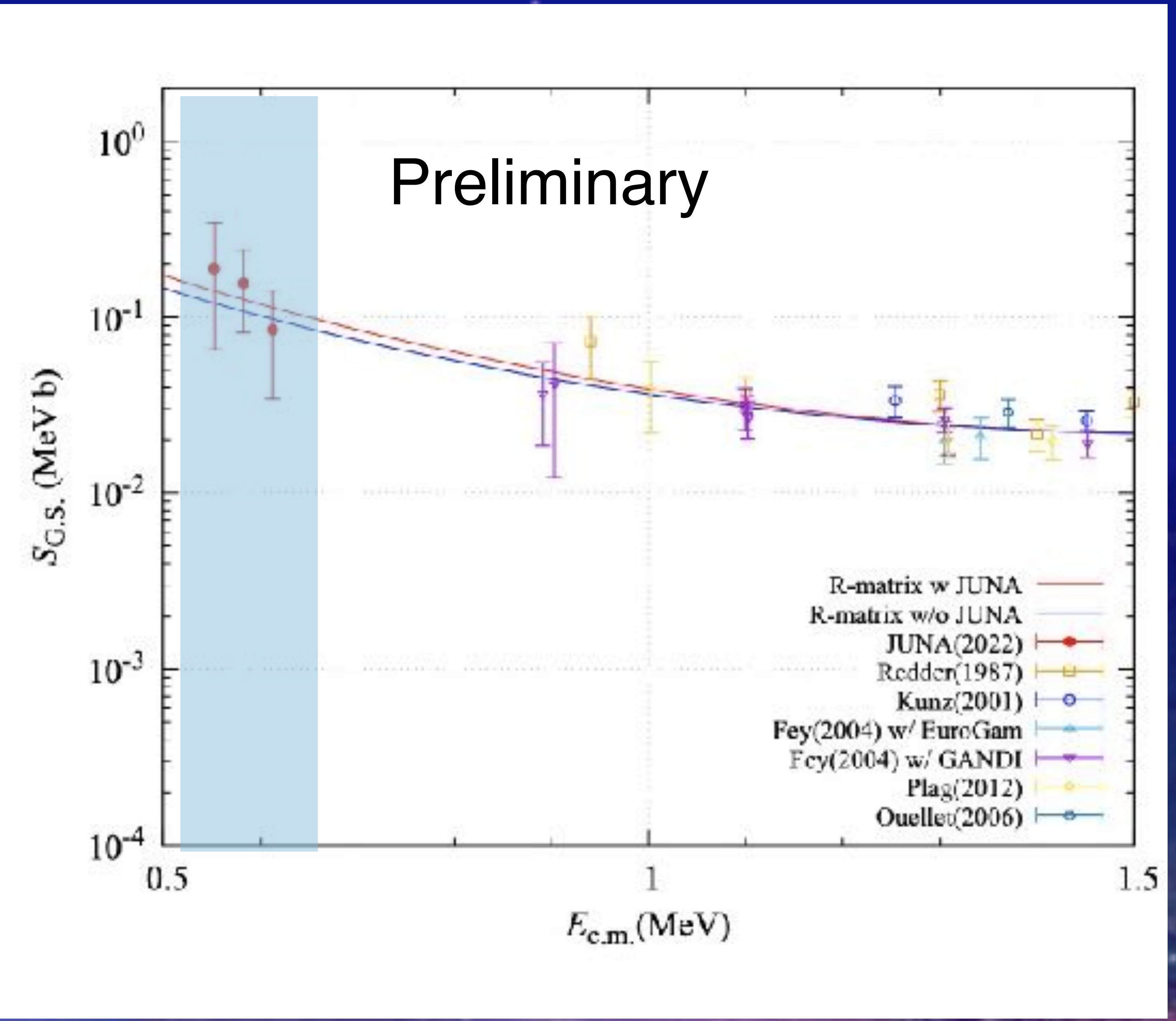


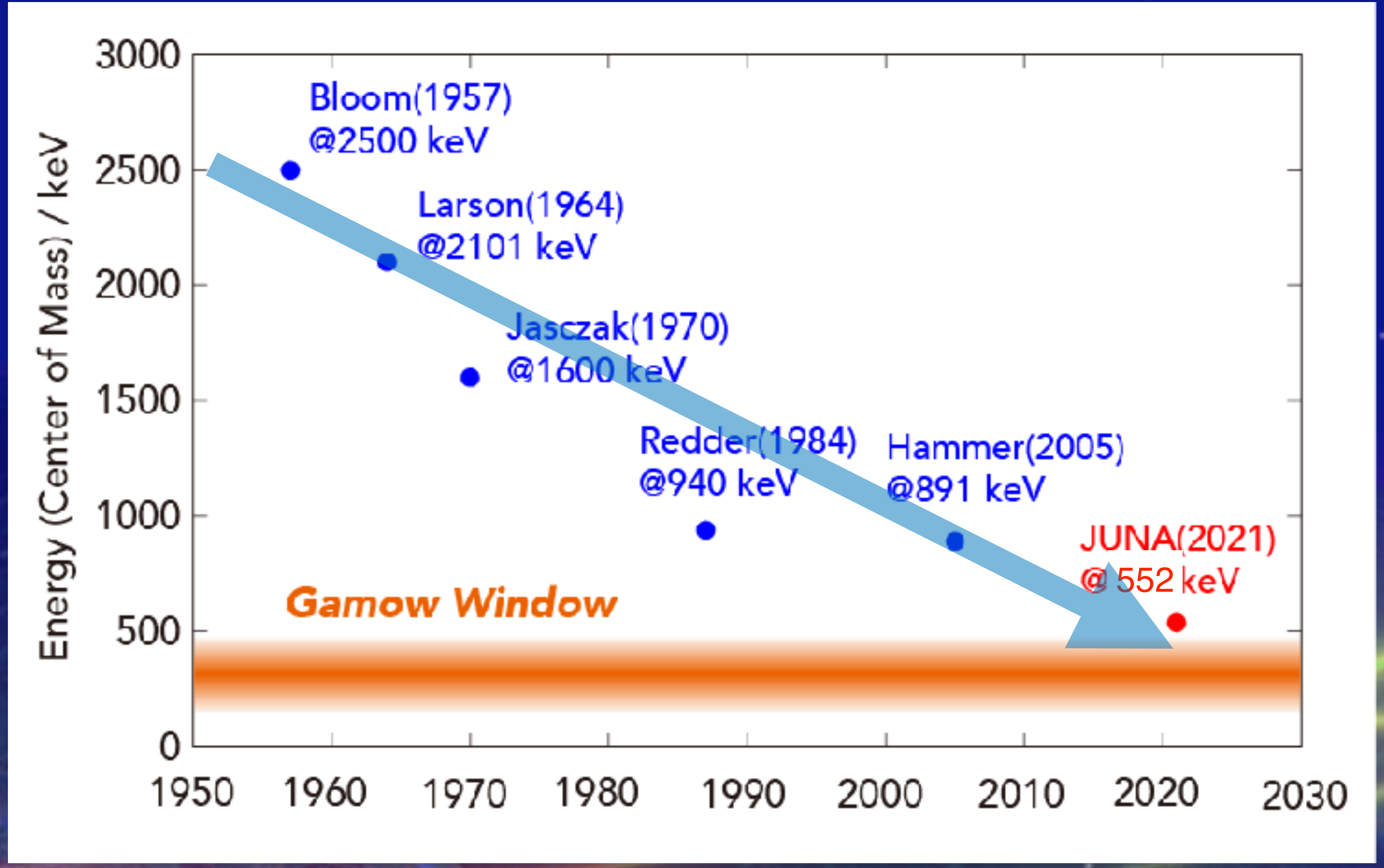
612 keV, 200 C











JUNA results from Run-1

Reaction	Quantities	Best data before	JUNA data	Publication
Holy grail $^{12}C(\alpha,\gamma)^{16}O$	Lowest energy/keV	891	**552**	**In preparation**
	Cross section/b	10^{-11}	**10^{-12}**	
Neutron source $^{13}C(\alpha,n)^{16}O$	Energy range/keV	230-300	**240-1900**	**PRL 129(2022)132701**
	s-process	50%	**20%**	
^{26}Al abundance $^{25}Mg(p,\gamma)^{26}Al$	Uncertainty	21%	**8%**	**Science Bulletin 67(2022)2 cover paper**
F abundance $^{19}F(p,\alpha\gamma)^{16}O$	Lowest energy/keV	189	**72**	**PRL 127(2021)152702 Editor suggestion**
	Uncertainty	80%	**5%**	
Ne isotope ratio $^{18}O(\alpha,\gamma)^{22}Ne$	Uncertainty	472±18 keV	**474.1±1.1 keV**	**PRL 130(2023)092701**
CNO breakout $^{19}F(p,\gamma)^{20}Ne$	Lowest energy/keV	300	**200**	**Nature 610(2022)656 news and views**

Recent development in underground nuclear astrophysics

锦屏深地核天体物理实验
Jinping Underground Nuclear Astrophysics Experiment
JUNA
400 kV
800 kV
Super-JUNA
25
20
15
10
5
0
0.1
1
$I(^4He)$ / mA
$E(^4He)$ / MeV
$^{12}C(\alpha,\gamma)^{16}O$
Super JUNA
JUNA
$^{22}Ne(\alpha,n)^{25}Mg$
$^{13}C(\alpha,n)^{16}O$
JUNA
Super JUNA

JUNA and Super JUNA coverage

H burning

$^{7}Be(p,\gamma)^{8}B$

$^{12}C(p,\gamma)^{13}N$

$^{14}N(p,\gamma)^{15}O$

$^{15}N(p,\gamma),(p,\alpha)^{16}O,^{12}C$

$^{17}O(p,\gamma),(p,\alpha)^{18}F,^{14}N$

$^{18}O(p,\gamma),(p,\alpha)^{19}F,^{15}N$

$^{19}F(p,\alpha)^{16}O$

$^{3}He(\alpha,\gamma)^{7}Be$

He burning

$^{12}C(\alpha,\gamma)^{16}O$

$^{16}O(\alpha,\gamma)^{20}Ne$

$^{20}Ne(\alpha,\gamma)^{24}Mg$

$^{18}O(\alpha,\gamma)^{22}Ne$

$^{22}Ne(\alpha,\gamma)^{26}Mg$

$^{24}Mg(\alpha,\gamma)^{28}Si$

N source

$^{13}C(\alpha,n)^{16}O$

$^{22}Ne(\alpha,n)^{25}Mg$

$^{25}Mg(\alpha,n)^{28}Si$

$^{26}Mg(\alpha,n)^{29}Si$

C\O burning

$^{12}C+^{12}C$

$^{12}C+^{16}O$

$^{16}O+^{16}O$

γ astronomy

$^{25}Mg(p,\gamma)^{26}Al$

$^{26}Al(p,\gamma)^{27}Si$

JUNA achieved

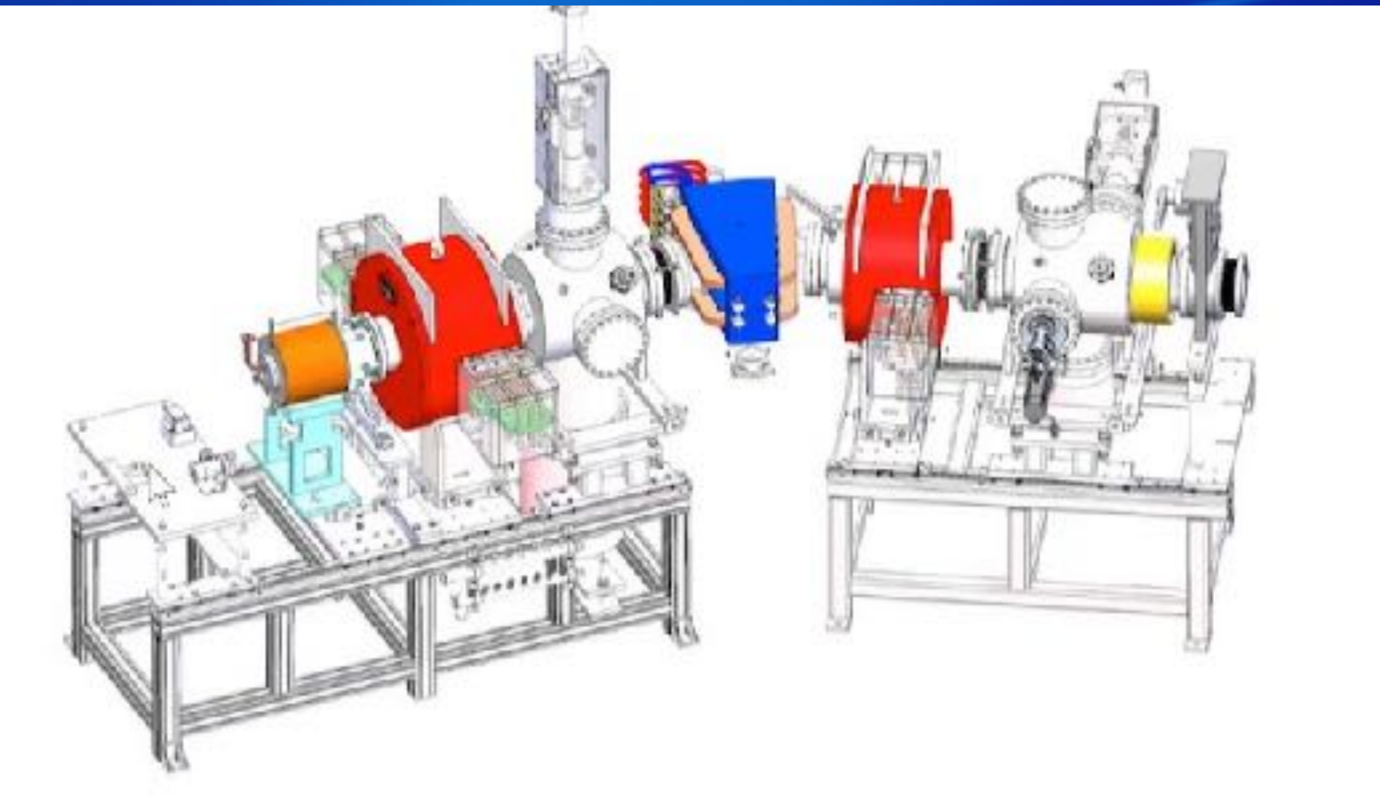
JUNA and Super JUNA proposed

2024-2030 JUNA and Super-JUNA

	JUNA	Super-JUNA	JUNA Exp.	Super-JUNA Exp.
2024	ground test run Run-2 exp.	R&D	(p,g)	
2025	Run-2	Ground test and fabrication	(a,n) and (a,g) gas target	
2026	Run-2	Setup and test	cont.	
2027	Upgrade Run-3 exp.	Test run Run-4 exp.	Test run	(p,g) test
2028	Run-3	Run-4	cont.	(a,g), (a,n) Exp.
2029	Run-3	Run-4	cont.	(a,g), (a,n) Exp.
2030	Run-3	Run-4	cont.	(a,g), (a,n) Exp.

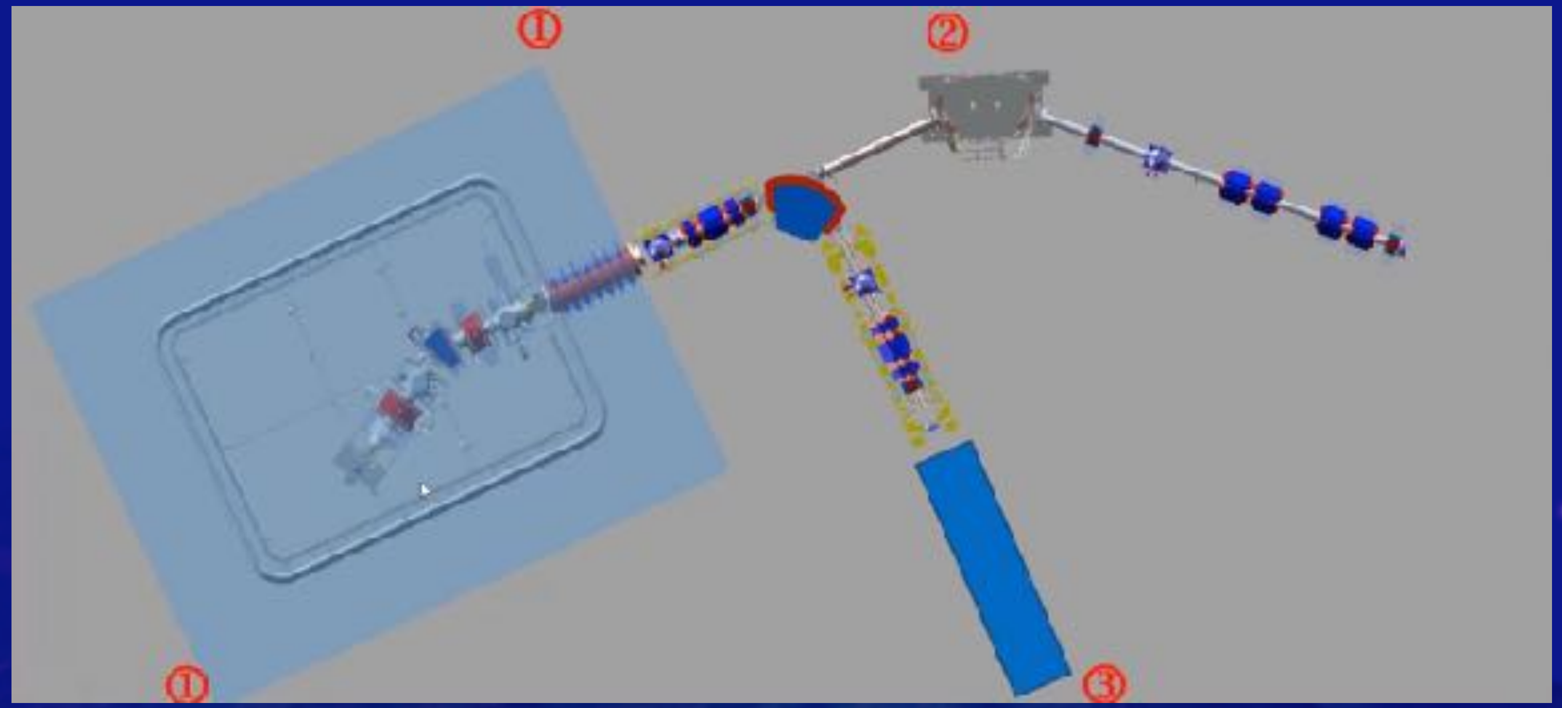
Green lights for JUNA Run-2

- CJPL IAC highly recommend JUNA and gave green lights for next 5 years and support for A1 fine tuning
- High density radiation hard target and gas target
- High resolution and efficiency neutron and gamma detector



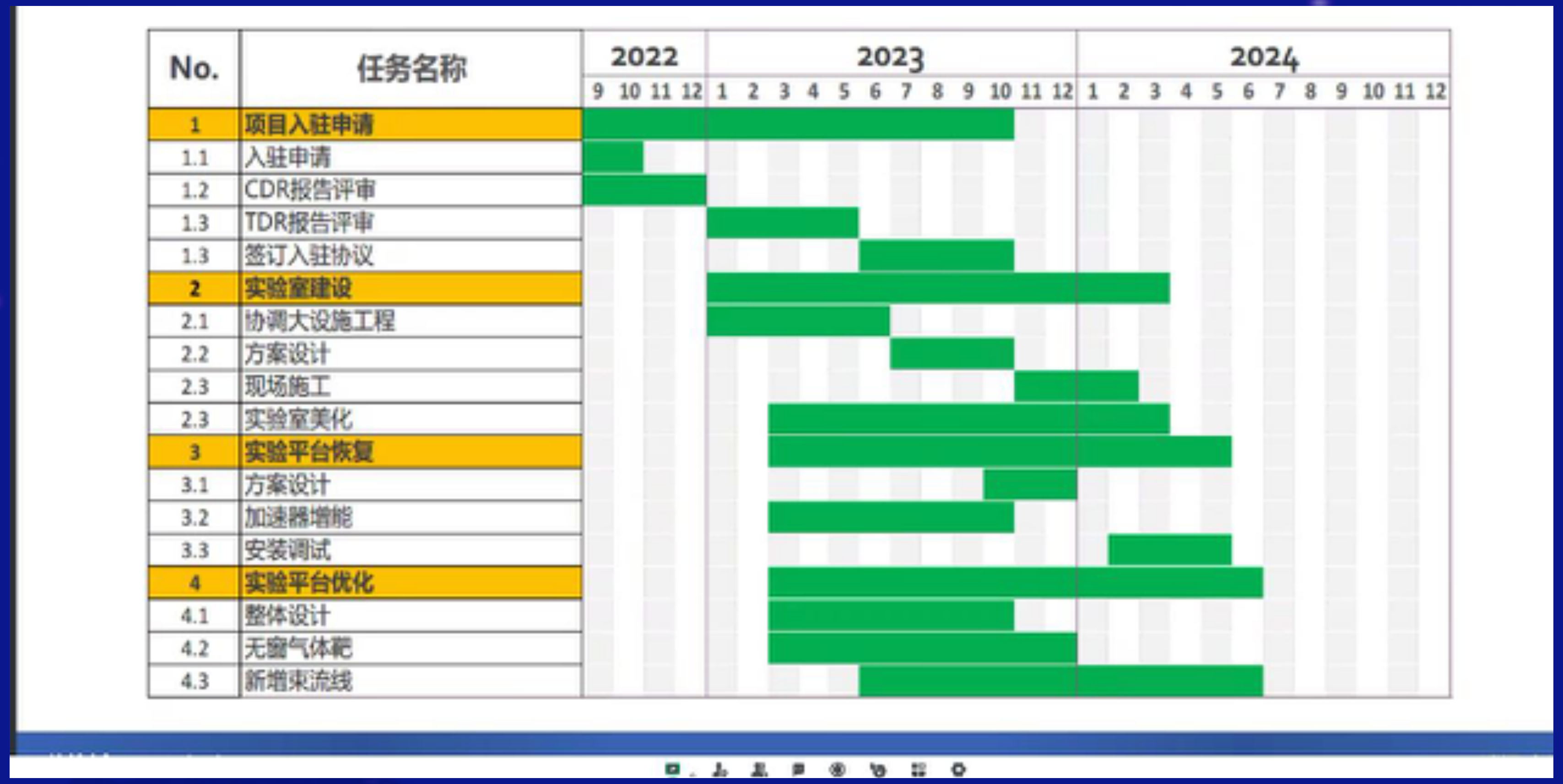
Improved ion source

Run-2 kickoff meeting April 24, 2024

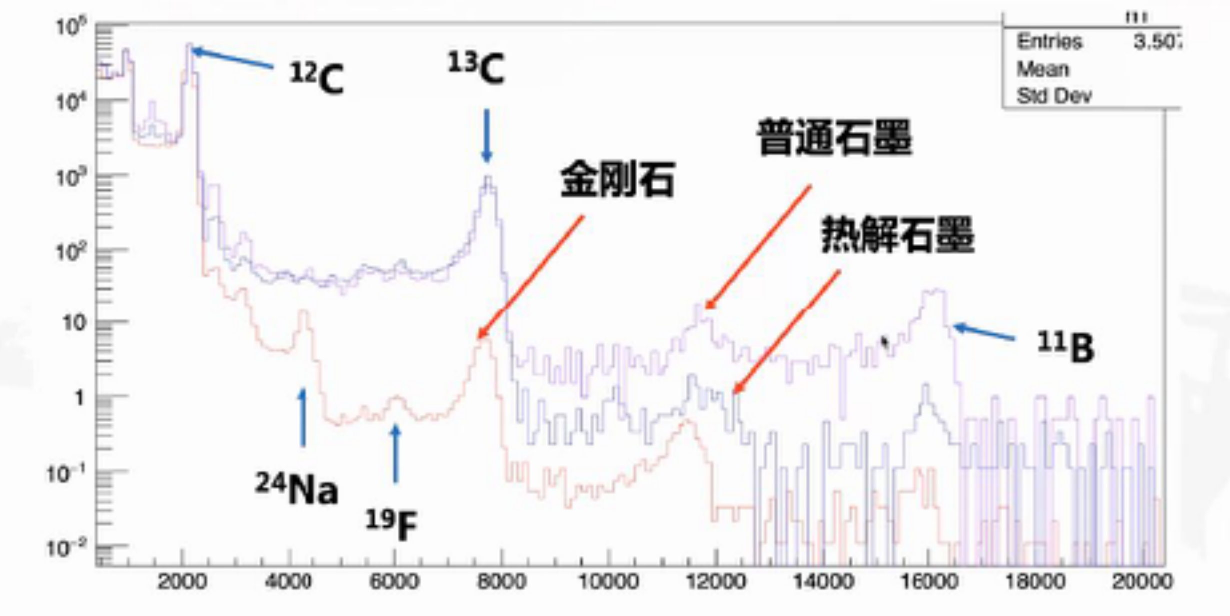


floor plan for JUNA Run-2

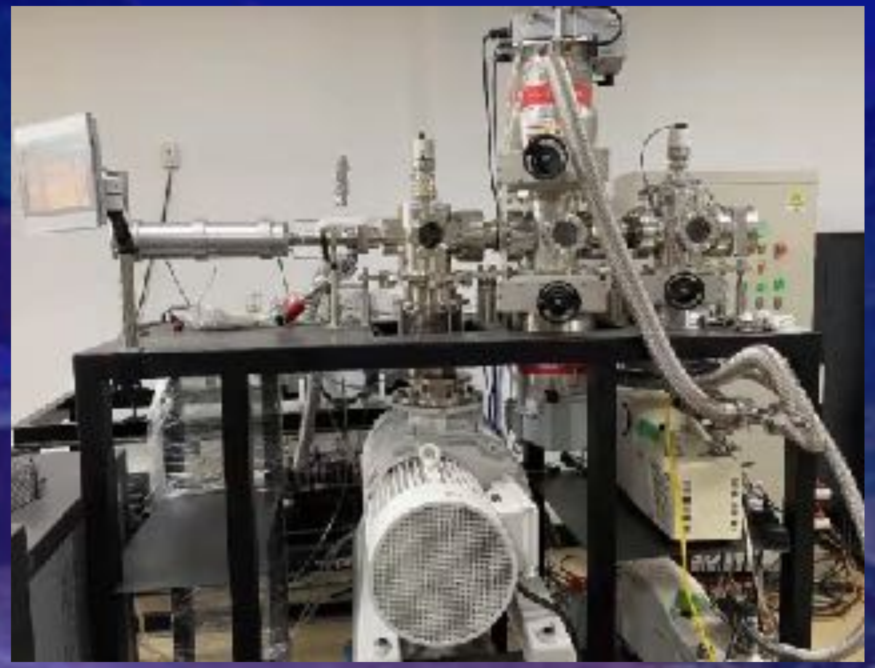
CJPL-II A1 for JUNA: March, 2024



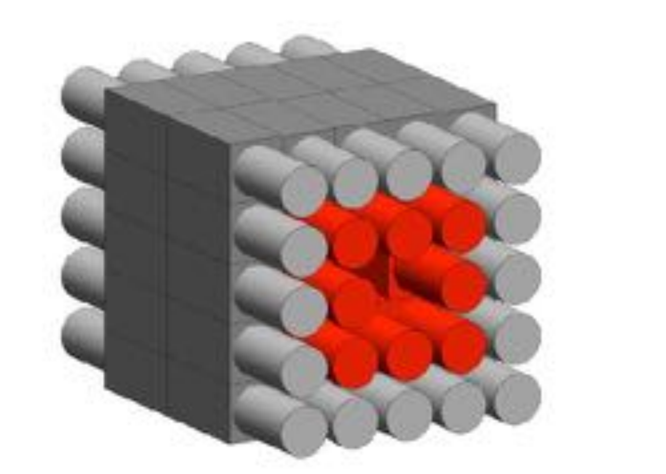
Run-2 schedule April 24, 2024



test result for diamond target



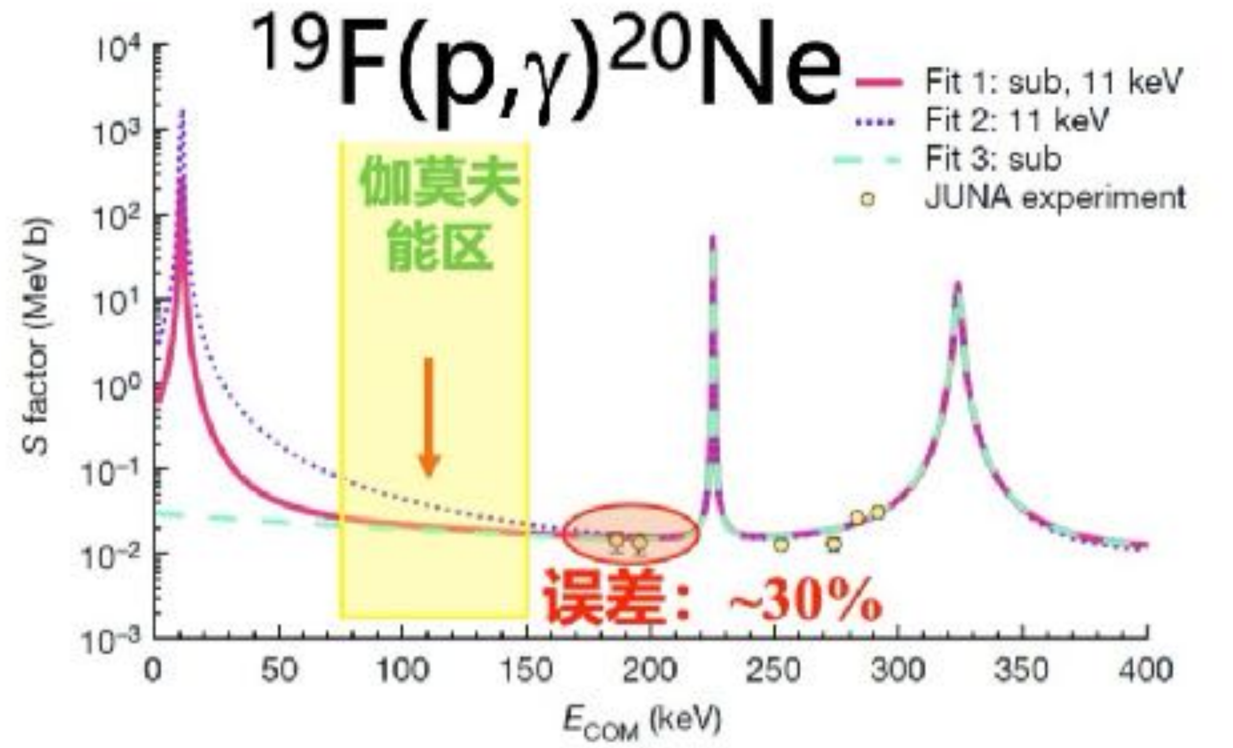
gas jet target

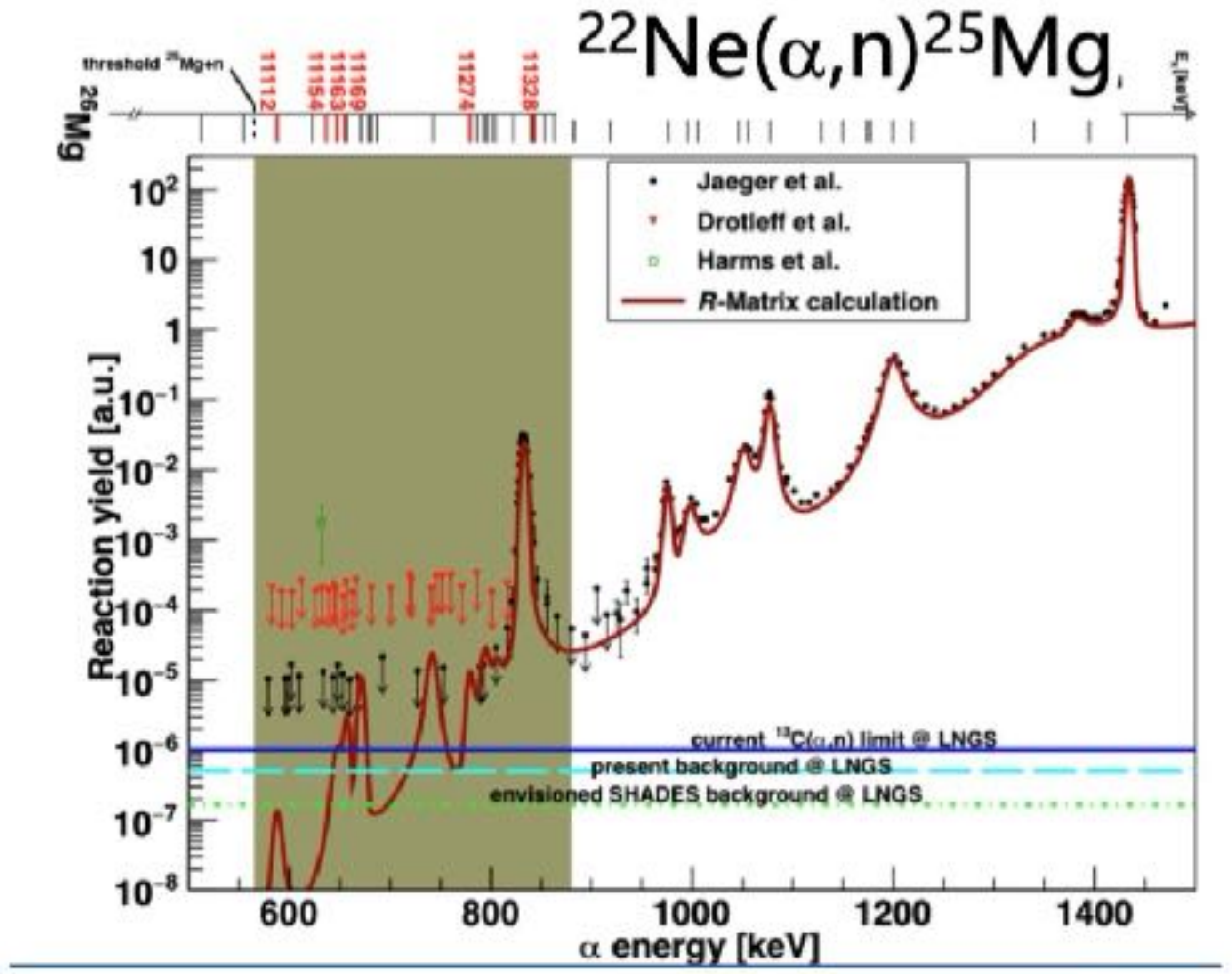


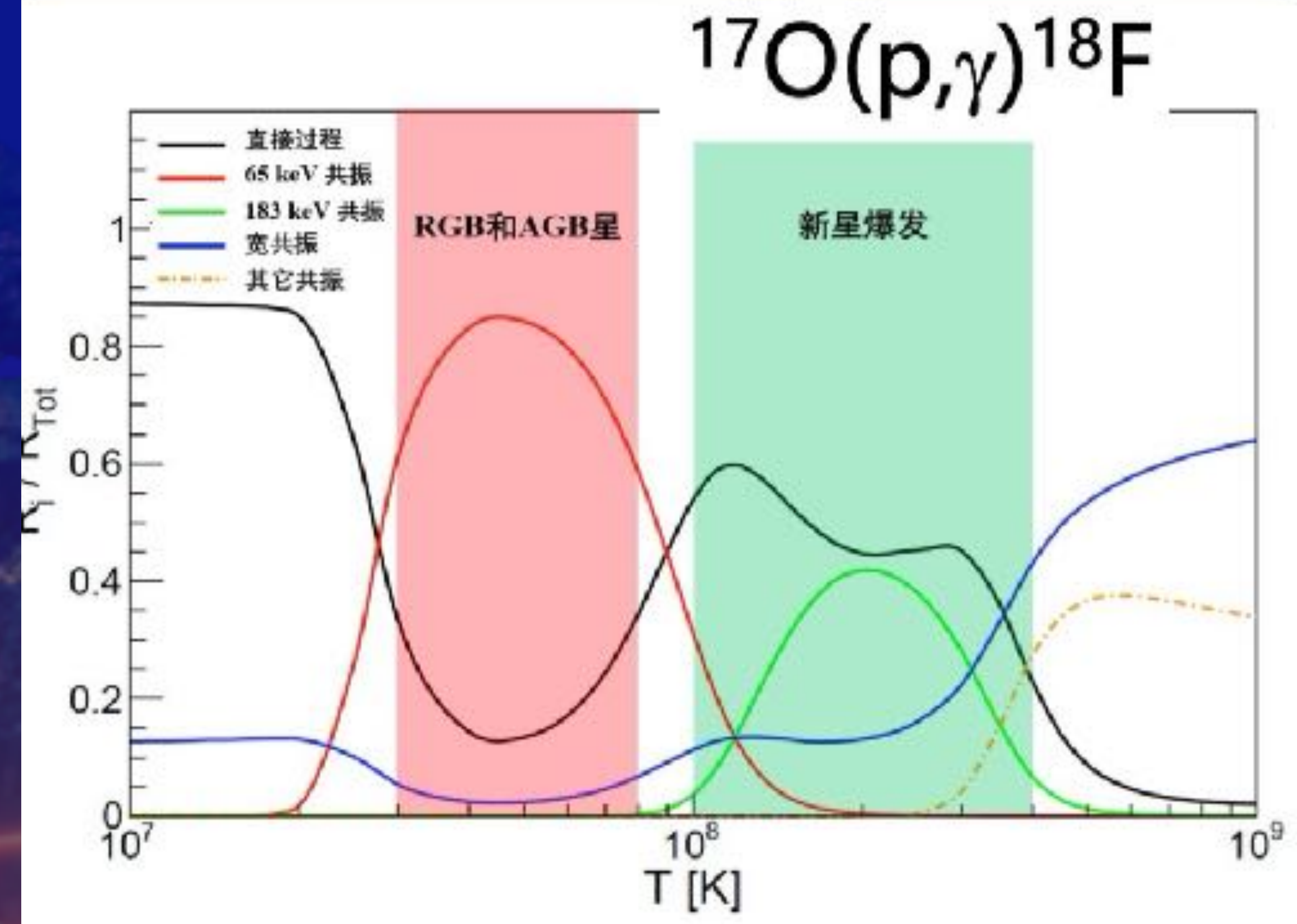
enlarged BGO array

JUNA Run-2 Exp.: 2024-2026

- From Run-1 to lower energy
  - $^{12}C(\alpha,\gamma)^{16}O$, precision from 1s to 3s
  - $^{13}C(\alpha,n)^{16}O$, full coverage of s-process
  - $^{19}F(p,\gamma)^{20}Ne$, cover 80-150 keV with high precision
  - $^{14}N(p,\gamma)^{15}O$, Solar neutrino
- Using gas target
  - $^{22}Ne(\alpha,n)^{25}Mg$, weak s-process n source
  - $^{3}He(\alpha,\gamma)^{7}Be$, solar neutron, Li problem, 80-380 keV
- Others
  - $^{17}O(p,\alpha)^{14}N$, ^{17}O over abundance
  - $^{17}O(p,\gamma)^{18}F$, H isotope ratio, 65 keV resonance
  - $^{26}Al(p,\gamma)^{27}Si$, BRIF ISOL ^{26}Al implantation target
  - $^{10}B(\alpha,n)^{13}N$, search for new resonance







JUNA summary

- JUNA is an advanced deep astrophysics platform. China, follow Italy and United States and others, started to carry out direct measurement of key astrophysical reactions, which leading the nuclear astrophysics to the stage of precision numerical simulation stage

- JUNA accurately measured key nuclear astrophysical reactions, compared with previous experiment, beam intensity is higher, detector efficiency, target exposure, sensitivity and energy coverage are greatly improved

- From JUNA Run-1, Gamma-ray astronomical reaction has reached the highest precision, and the astrophysical holy grail reaction has achieved the highest sensitivity, new resonances revealing the origin of heavy element abundance in the oldest stars, and the discrepancies of neutron source reactions was resolved

- JUNA Run-2 will start by the end of 2024, welcome to join JUNA collaboration and submit your proposals deep underground!

JUNA Team

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